

# HIP 10725: The First Solar Twin/Analogue Field Blue Straggler<sup>★</sup>

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## ABSTRACT

**Context.** Blue stragglers are easy to identify in globular clusters, but are much harder to identify in the field. Here we present the serendipitous discovery of one field blue straggler, HIP 10725, that closely matches the Sun in mass and age, but with a metallicity slightly lower than the Sun's.

**Aims.** To characterise the solar twin/analogue HIP 10725 in order to assess whether this star is a blue straggler.

**Methods.** We employ high resolution ( $R \sim 10^5$ ) high S/N (330) VLT/UVES spectra to perform a differential abundance analysis of the solar analogue HIP 10725. Radial velocities obtained by other instruments were also used to check for binarity. We also study its chromospheric activity, age and rotational velocity.

**Results.** We find that HIP 10725 is severely depleted in beryllium ( $[\text{Be}/\text{H}] \leq -1.2$  dex) for its stellar parameters and age. The abundances relative to solar of the elements with  $Z \leq 30$  show a correlation with condensation temperature and the neutron capture elements produced by the s-process are greatly enhanced, while the r-process elements seem normal. We found its projected rotational velocity ( $v \sin i = 3.3 \pm 0.1$  km/s) to be significantly larger than solar, and incompatible with its isochrone-derived age. Radial velocity monitoring shows that the star has a binary companion.

**Conclusions.** Based on the high s-process element enhancements and low beryllium abundance, we suggest that HIP 10725 has been polluted by mass-transfer from an AGB star, probably with initial mass of about  $2 M_{\odot}$ . The radial velocity variations suggest the presence of an unseen binary companion, probably the remnant of a former AGB star. Isochrones predict a solar-age star, but this is in disagreement with the high projected rotational velocity and high chromospheric activity. We conclude that HIP 10725 is a field blue straggler, rejuvenated by the mass transfer process of its former AGB companion.

**Key words.** Sun: abundances – stars: fundamental parameters — stars: abundances – blue stragglers

## 1. Introduction

Blue stragglers are traditionally recognised as main-sequence stars, which are significantly bluer than the main-sequence turn-off of the population to which they belong (Ryan et al., 2001); this phenomenon has been also observed in redder stars (e.g. Sills et al., 2000). Although first identified observationally in clusters (Sandage, 1953), they are also present in the field (e.g., Carney & Peterson, 1981). Field blue stragglers are harder to identify than their cluster counterparts, since it is a lot more difficult to reliably establish which stars share a common origin in this case (since field stars which initially belonged to a cluster have already dispersed out around the galaxy). However, there are other clues left behind by the blue straggler formation processes which can aid in their identification, especially in the case of cooler stars.

Chromospheric activity in blue stragglers tends to be incompatible with their isochrone-derived ages; they present significant rotational velocities, thus appearing to be younger (Fuhrmann & Bernkopf, 1999). Severe Li depletion is also observed, a sign of old age (Monroe et al., 2013; Meléndez et al., 2014a) which is also in contradiction with the young age scenario. When the phenomenon occurs through mass transfer events (i.e., a mass exchange between a now "dead" star during its AGB phase, which pollutes its companion's photosphere, in the process also transferring angular momentum (McCrea, 1964)), enhancement of s-process elements is predicted for a sufficiently massive AGB donor, as well as abundance anomalies of light elements (Desidera et al., 2007). In this case, one also expects to find radial velocity variations due to what is now a white dwarf companion. Although these properties could be found individually in normal stars, their presence alone is not an indication of the blue straggler status; their combination, however, strongly suggests so (Rocha-Pinto et al., 2002).

In the case of solar analogs the argument can be further strengthened by analysing Be abundances. Large beryllium depletion is not observed in solar twins, where its abundance

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is relatively constant with age (Tucci Maia et al., 2015)<sup>1</sup>, unlike lithium, which is continuously being depleted as the star evolves (Baumann et al., 2010; Meléndez et al., 2010, 2014a; Monroe et al., 2013) by extra-mixing below the convective zone. This suggests that when Be depletion is present in these stars, other mechanisms must be responsible for it. The blue straggler scenario is a possible explanation. A small Be depletion would be due to the mass transferred by the AGB companion, as it would already be devoid of this element. However, a larger Be depletion could be produced either by enhanced internal mixing due to the transfer of angular momentum, or to a complete destruction of Be as a result of a merger.

In this work, we present evidence for the blue straggler status of the solar twin/analogue HIP 10725, being the field blue straggler whose stellar parameters ( $T_{\text{eff}}$ ,  $\log g$  and metallicity) are the closest to solar. This star could be classified as a solar twin if the volatile elements (C, N, O) are used to characterize its metallicity, or as a solar analog if the depleted refractory elements (such as iron) are used instead.

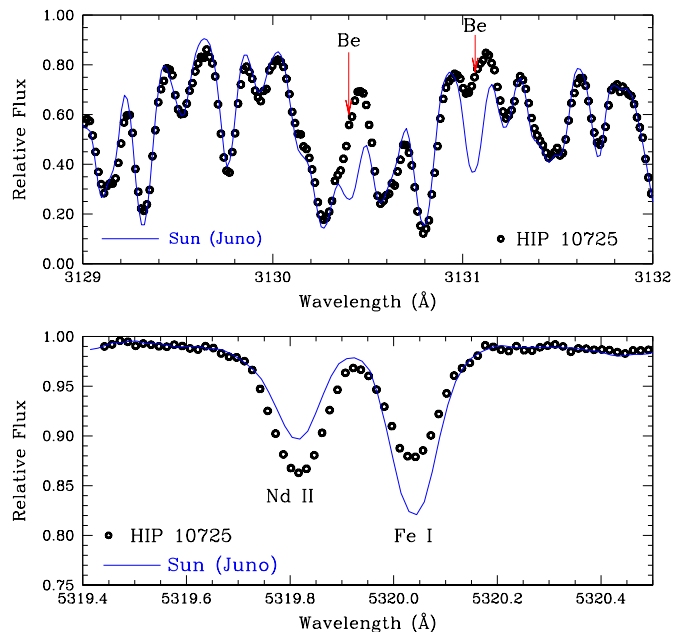
## 2. Observations

In our quest to identify new solar twins, we chose eight candidates to be studied at greater detail (HIP 1536, HIP 3238, HIP 10725, HIP 11514, HIP 106288, HIP 109381, HIP 114328 and HIP 117499), based on their colours and Hipparcos parallaxes (i.e., their position on the Colour-Magnitude Diagram). One of those stars (HIP 10725) showed an unusually low Be abundance, which sparked further analyses.

The observations for this sample are described in Meléndez et al. (2014a). Briefly, we observed with UVES@VLT in dichroic mode, with the 346 nm setting (306-387 nm) in the blue arm and the 580 nm setting (480 - 682 nm) in the red arm. Most employed spectral lines are located in the red arm, where we achieved  $R = 110\,000$  (0.3 arcsec slit) and S/N of about 330 per pixel at 600 nm. In the UV we used a slit of 0.6 arcsec, resulting in  $R = 65\,000$ . A reference solar spectrum was obtained using the asteroid Juno with identical setup. The echelle orders were extracted and wavelength calibrated using IRAF, and further processing was performed with IDL.

We also monitored radial velocities for this star to detect possible variations, and were able to acquire four measurements at different epochs. The first one was the value published by Jenkins et al. (2011) based on a FEROS/ESO spectrum taken in October 2008. Our UVES/VLT data described above was obtained in August 2009. Furthermore we obtained two more recent measurements, one taken in July 2014 with UVES/VLT, and another one in August 2014 using the Coudé spectrograph at the Observatório Pico dos Dias (OPD) in Brazil. The radial velocities are reported in Table 1 and are discussed in Sect. 4.

Fig. 1 shows part of the reduced spectra of both the Sun and HIP 10725 around the beryllium feature (*top panel*) and around 5320 Å (*bottom panel*). Given the similarity in stellar parameters, the Be depletion in HIP 10725 is clearly visible, as well as the enhanced Nd abundance and lower iron content.



**Fig. 1.** Comparison of the spectra of HIP 10725 (circles) and the Sun (solid line). In the top panel, we focus on two BeII features in the UV region around 3030 Å. HIP 10725 shows extreme Be depletion compared to the Sun, which is not expected for solar analogues (Tucci Maia et al., 2015). The bottom panel shows an enhancement of the neutron capture element Nd in HIP 10725 as well as a substantially lower iron content relative to the Sun.

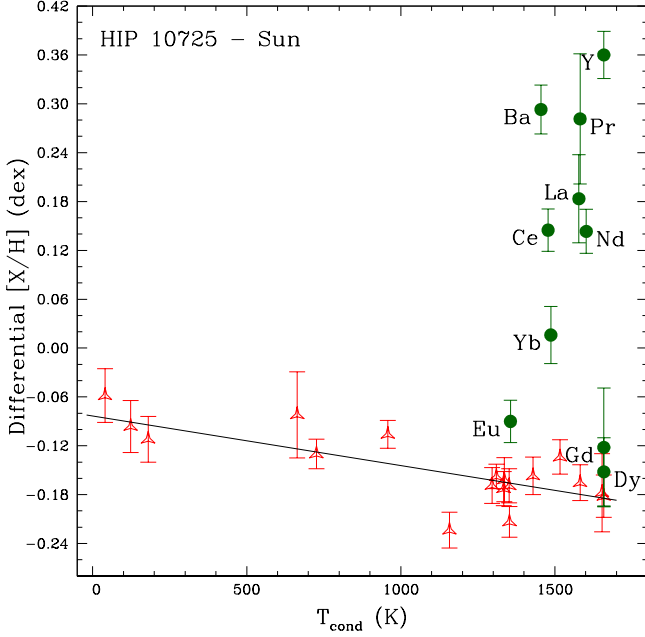
## 3. Abundance analysis

The same differential approach as in our previous papers (e.g. Meléndez et al., 2014a; Ramírez et al., 2014) was employed to obtain stellar parameters and chemical abundances, i.e., we followed a strictly differential line-by-line analysis, using ATLAS9 model atmospheres (Castelli & Kurucz, 2004) and the 2002 version of the chemical abundance analysis program MOOG (Snedden, 1973). The equivalent width (EW) measurements were performed manually using IRAF's *splot* task. The line list used is presented in Meléndez et al. (2014b).

The differential spectroscopic equilibrium of HIP 10725 relative to the Sun results in stellar parameters of  $T_{\text{eff}} = 5777 \pm 16$  K ( $\Delta T_{\text{eff}} = +5 \pm 16$  K),  $\log g = 4.45 \pm 0.05$  dex ( $\Delta \log g = +0.01 \pm 0.05$  dex),  $[\text{Fe}/\text{H}] = -0.17 \pm 0.01$  dex, and a microturbulent velocity of  $0.97$  km s<sup>-1</sup> ( $+0.07 \pm 0.04$  km s<sup>-1</sup> higher than solar). The uncertainties in the stellar parameters are based on the observational uncertainties. As the stellar parameters are interdependent of each other, we also took into account in the error budget this degeneracy.

With these stellar parameters, we computed differential abundances using the measured EWs, except for Li and Be which were analysed by spectral synthesis following the procedure described in Monroe et al. (2013). For Be the 2014 version of MOOG and the line list of Tucci Maia et al. (2015) were employed. Hyperfine structure was taken into account for V, Mn, Co, Y, Ba, La, Pr, Eu and Yb. The differential abundances are provided in Table 2 along with the uncertainties stemming from observations (standard error in abundances) and stellar parameters. The total error was obtained by adding in quadrature the observational and parameter uncertainties.

<sup>1</sup> Only 4 stars in the sample of 118 solar analogs studied by Takeda et al. (2011) are extremely depleted in Be. The origin of this depletion is under investigation (Viallet & Baraffe, 2012; Desidera et al., 2015).



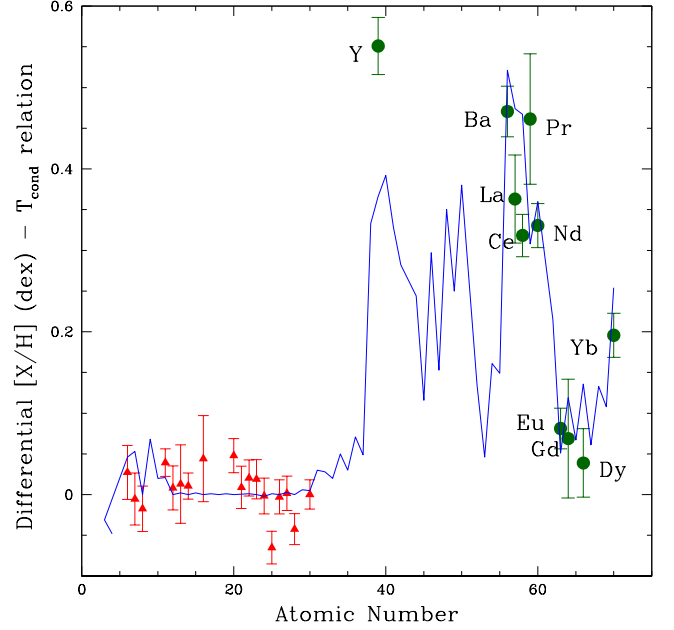
**Fig. 2.** Differential abundances of HIP 10725 relative to the Sun. The triangles represent the lighter elements ( $Z \leq 30$ ) while the filled circles represent neutron-capture elements with  $Z > 30$ . The solid line is a fit of  $[X/H]$  versus  $T_{\text{cond}}$ , taking into account only the lighter elements. The fit yields  $[X/H]_{(Z \leq 30)} = -0.0834 - 6.10 \times 10^{-5} T_{\text{cond}}$ .

Notice from the abundance pattern given in Table 2, that the average of the abundant volatile elements C, N, O, are within 0.1 dex of that of the Sun, but that the iron abundance is lower. This has implications regarding the classification of a star as either a solar twin or a solar analog. Solar twins are classified as stars having  $T_{\text{eff}}$  within 100 K of the solar effective temperature, and with  $\log g$  and  $[\text{Fe}/\text{H}]$  within 0.1 dex of the Sun's (Ramírez et al., 2009).  $[\text{Fe}/\text{H}]$  is used as a proxy of metallicity because it is easier to measure than other more abundant elements such as oxygen and carbon. The definition of solar twin could depend on the element chosen to characterise metallicity. If we consider the average of the abundant volatile elements (C, N, O), then HIP 10725 would be classified as a solar twin, as its metallicity would be within 0.1 dex of that of the Sun, but if the metallicity is based in the depleted refractories (e.g., Fe), then HIP 10725 would be a solar analog. In any case, HIP 10725 is a star that closely resembles the Sun.

We estimated the projected rotational velocity  $v_{\text{rot}} \sin i$  for this star using the same procedures and line list as in Tucci Maia et al. (2015). First, we obtained the macroturbulence velocity  $V_{\text{macro}} = 3.6 \text{ km s}^{-1}$  using the relation given in the aforementioned work. With the macroturbulence fixed, we determined  $v_{\text{rot}} \sin i = 3.3 \pm 0.1 \text{ km s}^{-1}$  by spectral synthesis of six lines (5 Fe I lines and 1 Ni I line). Thus the rotational velocity of HIP 10725 is significantly larger than the solar value ( $1.9 \text{ km s}^{-1}$ ).

#### 4. Discussion

The abundance pattern of HIP 10725 is peculiar in the sense that we observe an excess of s-process elements, no enhancement of the r-process elements, and also a trend with conden-

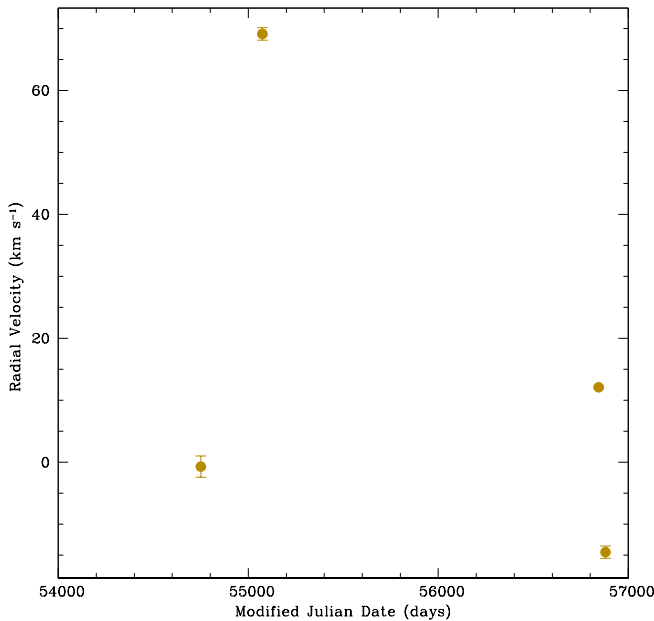


**Fig. 3.** Differential abundances of HIP 10725 relative to the Sun after the trend with condensation temperature has been subtracted. The heavier neutron-capture elements (filled circles) exhibit abundance enhancements, which are fairly well reproduced by mass transfer from an AGB star with initial mass of  $2 M_{\odot}$  (solid line), except for yttrium, that is underproduced by the models.

sation temperature ( $T_{\text{cond}}$ ), as can be seen in Fig. 2. Usually, this trend is positive in stars closely resembling the Sun, i.e. there is an overabundance of refractory elements relative to the Sun (Meléndez et al., 2009; Ramírez et al., 2009, 2010). In this star, however, we observe the opposite. In our previous works on solar twins, we speculated that the depletion of refractory elements was associated with the formation of rocky planets, but in the context of the present work, as suggested by Desidera et al. (2007), the correlation is more likely associated with dust formation in the outskirts of the former AGB companion (van Winckel, 2003; Waters et al., 1992): some of the refractory elements formed dust that was removed from the gas accreted onto HIP 10725.

Figure 2 shows that the elements that depart the most from the overall  $T_{\text{cond}}$  trend are those that are known to have important contributions from the s-process to their solar system abundances (e.g., Ba, La, Y), whereas those with the lowest s-process contribution to those abundances (e.g., Eu, Dy, Gd) fall very close to the  $T_{\text{cond}}$  trend (see for example Simmerer et al. (2004)).

In order to study this peculiar abundance pattern, we follow the procedure outlined in Meléndez et al. (2014b): we fit  $[X/H]$  versus condensation temperature (Lodders, 2003) for the lighter elements ( $Z \leq 30$ ), and subtract this from all abundances. The fit can be seen in Fig. 2 and yields  $[X/H]_{(Z \leq 30)} = -0.0834 - 6.10 \times 10^{-5} T_{\text{cond}}$ , with an element-to-element scatter of 0.03 dex. Fig. 3 shows the new abundance ratios after subtraction. Notice that while the s-process elements like Ba, Nd and Y are enhanced, the r-process elements Eu, Gd and Dy are not, suggesting that indeed the pollution is due to a former AGB companion.



**Fig. 4.** Plot showing radial velocity data obtained for HIP 10725 over the course of several years, evidence for the presence of the unseen white dwarf companion.

We modeled the pollution by the former AGB star using solar-metallicity models by Karakas (2014). We calculated the *s*-process nucleosynthesis using the same method outlined in Karakas (2010) and Lugaro et al. (2012). We mixed the AGB ejecta into the solar analog’s shallow surface convection zone, which we assumed  $0.023M_{\odot}$  (2.3% of the stellar mass), similar to the solar convection zone. An AGB stellar model with initial mass of  $2 M_{\odot}$  provides a reasonable fit to the observed abundance pattern, as seen in Fig. 3. The amount of accreted AGB material is 0.2% of the mass lost from the  $2 M_{\odot}$  AGB star over its lifetime, with most of the material lost at the tip of the AGB during the last couple of thermal pulses. This choice provides a fair fit to Ba, Nd, Eu, Gd, Yb, and Dy. The predicted Y abundance is lower than observed, but this discrepancy could be related to the formation of  $^{13}\text{C}$  pockets in AGB stars or to the activation of the  $^{22}\text{Ne}$  neutron source in the thermal pulses (see review by Karakas & Lattanzio (2014)).

As shown above, mass transfer from an AGB star can account for the enhanced *s*-process pattern observed, as expected for blue stragglers formed by this process (Desidera et al., 2007). The donor star would subsequently have evolved into an unseen white dwarf companion. Such a companion star should be detectable through radial velocity monitoring, which is indeed the case. As seen in Fig. 4, we found radial velocity variations of up to  $\sim 70 \text{ km s}^{-1}$ , good evidence of the binarity of the system, which is common in the case of field blue stragglers (Preston & Sneden, 2000; Carney et al., 2005). The radial velocities and observation dates are presented in Table 1.

In addition, we also determined Be abundances for HIP 10725 using spectral synthesis, and found a value of  $A(\text{Be}) \leq 0.2$  dex, or  $[\text{Be}/\text{H}] \leq -1.2$  dex, for  $A(\text{Be})_{\odot} = 1.38$  dex (Asplund et al., 2009), meaning that the beryllium abundance in this star is at least 15 times lower than in the Sun. This is also compatible with the AGB mass transfer scenario. In the AGB scenario we propose, Be depletion would be small (0.05 dex)

due to the Be-depleted material transferred by the AGB star, and large (1.15 dex) due to the transfer of angular momentum from the AGB star, which would cause the solar twin to spin up, as suggested by its higher rotational velocity.

In order to study the depletion of Be we use the Toulouse-Geneva Evolution Code (TGEC, Hui-Bon-Hoa, 2008) including standard physics and also extra mixing beyond that predicted by standard models. We take into account diffusion (including gravitational settling) and rotation-induced mixing, as described in our previous works (do Nascimento et al., 2009; Castro et al., 2011; Tucci Maia et al., 2015). We modeled the Be depletion using as parameters the diffusion coefficient at the base of the convective zone  $D_{bcz}$ , and the half-height width of the tachocline  $\Delta$ , which are both free parameters of stellar evolution models.

The mass transfer from the AGB star into the surface of HIP 10725 should have enhanced the angular momentum of the convective envelope, increasing the stellar rotation. Convection transfers this increase until the base of the convective zone, where exists a transition zone between the differential rotation of the convective zone and the interior that rotates as a solid body. This transition zone, called tachocline (Spiegel & Zahn, 1992; Richard et al., 1996; Brun et al., 1998), induces a mixing at the base of the convective zone which can be modelled by an exponential effective diffusion coefficient  $D_{bcz}$ . In order to reproduce the observed Be depletion, we multiplied  $D_{bcz}$  by 1.7 and increased the size of the tachocline  $\Delta$  by a factor of 2.3. Doing so we can match the observed destruction of Be. The above changes in  $D_{bcz}$  and  $\Delta$ , result in a surface velocity of  $3.7 \text{ km s}^{-1}$ , close to the observed  $v_{\text{rot}} \sin i = 3.3 \text{ km s}^{-1}$ .

Enhanced chromospheric activity ( $R_{HK} = -4.51$ , Jenkins et al. (2011)) and high rotational velocity (both suggesting a young star) are incompatible with the isochrone-derived age obtained for HIP 10725 (Kim et al., 2002; Demarque et al., 2004), using probability distribution functions as in Meléndez et al. (2012). Our determined value of  $v_{\text{rot}} \sin i = 3.3 \text{ km s}^{-1}$ , suggests an age of roughly 1 Gyr, while the Yonsei-Yale isochrones result in an age of  $5.2_{-2.1}^{+1.9}$  Gyr. Our upper limit for the Li abundance ( $A(\text{Li}) \leq 0.9$  dex) is also about 20 times smaller than the Li abundance expected for an age of 1 Gyr (Tucci Maia et al., 2015), showing that the star is not young. Thus, the high projected rotational velocity is most likely the consequence of angular momentum transfer during the accretion stage from the former AGB star (Fuhrmann & Bernkopf, 1999), that “rejuvenated” HIP 10725, enhancing both rotation and chromospheric activity.

## 5. Conclusions

We present an observational study demonstrating that the solar analog HIP 10725 has been polluted by *s*-process material. Its high *s*-process element abundance is in good agreement with mass transfer from a  $2\text{-}M_{\odot}$  AGB star. A white dwarf should remain from the now defunct AGB star, and we have been able to detect the presence of a companion, through radial velocity variations. While the isochrone-derived age shows that its age is similar to solar, its chromospheric activity and a higher than usual rotational velocity would instead suggest that the star is young. This “rejuvenation” process is a consequence of mass and angular momentum transfer via wind accretion or Roche lobe overflow during the accretion period. A large beryllium depletion compared with solar-type stars of similar mass and age is present as well, providing further evidence for transfer of angular momentum having taken place.

Thus, we identify HIP 10725 as a field blue straggler. Future acquisition of more radial velocity measurements will help better constrain orbital parameters, and determine the mass of the probable white dwarf companion.

Finally, our work may be relevant to explain the large depletion in beryllium observed in four of the 118 solar analogs observed by Takeda et al. (2011), as some of these stars could also be field blue stragglers. Indeed, in a recent work, Desidera et al. (2015) shows that those four Be-depleted solar analogs have binary companions, and for two of them the companion is a white dwarf, showing thus a connection with the blue straggler phenomenon.

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**Table 1.** Radial velocity measurements

| M.J.D. <sup>a</sup><br>(days) | Radial Velocity<br>(km s <sup>-1</sup> ) | Error<br>(km s <sup>-1</sup> ) | Date<br>(yyyy-mm-dd) |
|-------------------------------|--|--------------------------------|----------------------|
| 54750.270436                  | -00.70                                   | 1.7                            | 2008-10-11           |
| 55074.321678                  | 69.10                                    | 1.0                            | 2009-08-31           |
| 56844.436451                  | 12.10                                    | 0.2                            | 2014-07-06           |
| 56880.270833                  | -14.50                                   | 1.0                            | 2014-08-11           |

**Notes.** <sup>(a)</sup> Modified Julian Date

**Table 2.** Differential abundances\* of HIP 10725 relative to the Sun and errors in abundances  $\Delta A$  due to observational and systematic (stellar parameters) uncertainties

| Element | [X/H]<br>LTE<br>(dex) | $\Delta A/\Delta T_{\text{eff}}$<br>$\Delta T_{\text{eff}}$<br>+16K<br>(dex) | $\Delta A/\Delta \log g$<br>$\Delta \log g$<br>+0.05 dex<br>(dex) | $\Delta A/\Delta v_t$<br>$\Delta v_t$<br>+0.04 km s <sup>-1</sup><br>(dex) | $\Delta A/\Delta [\text{Fe}/\text{H}]$<br>$\Delta [\text{Fe}/\text{H}]$<br>+0.01 dex<br>(dex) | error<br>param <sup>a</sup><br>(dex) | error<br>obs <sup>b</sup><br>(dex) | error<br>total <sup>c</sup><br>(dex) |
|---------|-----------------------|--|---|--|---|--------------------------------------|------------------------------------|--------------------------------------|
| C       | -0.058                | -0.008   | 0.011   | 0.000  | 0.001   | 0.013                                | 0.030                              | 0.033                                |
| N       | -0.096                | 0.022  | 0.007   | 0.008  | -0.007  | 0.025                                | 0.020                              | 0.032                                |
| O       | -0.112                | 0.021  | 0.006   | 0.001  | -0.007  | 0.023                                | 0.017                              | 0.028                                |
| Na      | -0.106                | 0.009  | 0.000   | -0.001   | 0.000   | 0.009                                | 0.015                              | 0.017                                |
| Mg      | -0.162                | 0.008  | -0.004  | -0.002   | -0.001  | 0.009                                | 0.026                              | 0.027                                |
| Al      | -0.178                | 0.008  | 0.001   | -0.001   | 0.000   | 0.008                                | 0.048                              | 0.048                                |
| Si      | -0.158                | 0.004  | 0.004   | -0.001   | -0.001  | 0.005                                | 0.015                              | 0.015                                |
| S       | -0.082                | -0.007   | 0.009   | -0.001   | 0.000   | 0.011                                | 0.052                              | 0.053                                |
| Ca      | -0.134                | 0.011  | -0.007  | -0.007   | -0.001  | 0.015                                | 0.015                              | 0.021                                |
| Sc      | -0.182                | 0.002  | 0.018   | -0.005   | -0.003  | 0.019                                | 0.018                              | 0.026                                |
| Ti      | -0.165                | 0.016  | 0.003   | -0.004   | 0.000   | 0.017                                | 0.015                              | 0.022                                |
| V       | -0.157                | 0.017  | 0.004   | -0.002   | -0.001  | 0.017                                | 0.015                              | 0.023                                |
| Cr      | -0.169                | 0.014  | -0.001  | -0.008   | -0.001  | 0.016                                | 0.015                              | 0.022                                |
| Mn      | -0.223                | 0.013  | -0.001  | -0.003   | 0.001   | 0.013                                | 0.015                              | 0.020                                |
| Fe      | -0.173                | 0.013  | -0.001  | -0.007   | -0.001  | 0.014                                | 0.015                              | 0.021                                |
| Co      | -0.169                | 0.012  | 0.008   | 0.002  | 0.000   | 0.015                                | 0.016                              | 0.022                                |
| Ni      | -0.213                | 0.010  | 0.002   | -0.005   | -0.001  | 0.011                                | 0.015                              | 0.019                                |
| Zn      | -0.130                | 0.002  | 0.005   | -0.008   | -0.002  | 0.009                                | 0.015                              | 0.018                                |
| Y       | 0.360                 | 0.003  | 0.008   | -0.021   | -0.022  | 0.031                                | 0.015                              | 0.035                                |
| Ba      | 0.293                 | 0.007  | 0.000   | -0.013   | -0.007  | 0.016                                | 0.026                              | 0.031                                |
| La      | 0.183                 | 0.005  | 0.029   | 0.001  | -0.005  | 0.030                                | 0.045                              | 0.054                                |
| Ce      | 0.145                 | 0.004  | 0.021   | -0.002   | -0.003  | 0.022                                | 0.015                              | 0.026                                |
| Pr      | 0.281                 | 0.031  | 0.045   | 0.025  | 0.025   | 0.066                                | 0.046                              | 0.080                                |
| Nd      | 0.143                 | 0.005  | 0.022   | -0.003   | -0.003  | 0.023                                | 0.015                              | 0.027                                |
| Eu      | -0.090                | 0.000  | 0.018   | -0.002   | -0.002  | 0.018                                | 0.017                              | 0.025                                |
| Gd      | -0.122                | 0.006  | 0.024   | -0.001   | -0.003  | 0.025                                | 0.069                              | 0.073                                |
| Dy      | -0.152                | 0.006  | 0.022   | -0.007   | -0.004  | 0.024                                | 0.035                              | 0.042                                |
| Yb      | 0.016                 | 0.003  | 0.021   | -0.007   | -0.005  | 0.023                                | 0.015                              | 0.027                                |

**Notes.** Abundances of V, Mn, Co, Y, Ba, La, Pr, Eu and Yb account for HFS.

<sup>(\*)</sup>  $A(\text{Li}) \leq 0.9$  dex,  $A(\text{Be}) \leq 0.2$  dex <sup>(a)</sup> Adding errors in stellar parameters <sup>(b)</sup> Observational errors <sup>(c)</sup> Total error (stellar parameters and observational)