

Flare stars in nearby Galactic open clusters based on *TESS* data

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Abstract. The study is devoted to search for flare stars among confirmed members of Galactic open clusters using high-cadence photometry from *TESS* mission. We analyzed 957 high-cadence light curves of members from 136 open clusters. As a result, 56 flare stars were found, among them 8 hot B-A type objects. Of all flares, 63% were detected in sample of cool stars ($T_{\text{eff}} < 5000$ K), and 29% – in stars of spectral type G, while 23% in K-type stars and approximately 34% of all detected flares are in M-type stars. Using the FLATW’RM (FLAre deTectioN With Ransac Method) flare finding algorithm, we estimated parameters of flares and rotation period of detected flare stars. The flare with the largest amplitude appears on the M3 type EQ Cha star. Statistical analysis did not reveal any direct correlation between ages, rotation periods and flaring activity.

Key words: Galaxy: open clusters and associations: general – stars: flare – stars: activity – space mission: TESS

1. Introduction

Flare stars – stars having detected at least one very short flare in their light curve – are a widely known type of objects in astrophysics. The physical mechanisms leading to the appearance of flares are associated with convective atmospheres, such as in G–M type stars on the main sequence. Thus most of the stars in our Galaxy are potentially flare stars.

In the middle of 20th century, only six flare stars were known (Lippincott, 1952). Since then, a huge progress was made in the field of detecting stellar flares. The interest of astronomers to these objects is associated not only with the incompletely studied mechanisms of formation of flares, but also with the possible influence of flares on habitable zone of potential exoplanets. The *Kepler* mission had a significant impact on the study of flare stars. As a result, *The Kepler Catalog of Stellar Flares* with more than 4000 detected flare stars was made by Davenport (2016). And even more significant impact can be expected from the ongoing *Transiting Exoplanet Survey Satellite (TESS)* mission, which already covered full celestial sphere. *TESS* satellite is a space-borne telescope, whose main objective is the detection and study of exoplanets using the technique of transits. However, its database of a high-cadence (2 minute effective sampling) light curves for hundreds of thousands of objects, as well as sequences of Full Frame Images (30 minutes effective sampling), may also provide an invaluable resource for the study of many other kinds of astrophysical objects, including flare stars.

Stellar flares are explosive magnetic reconnection events in a star’s magnetosphere (Günther et al., 2020) emitting energy through radio wavelength, X-ray, UV, optical and IR band (Lawson et al., 2019). They show a sharp rise in intensity followed by an exponential decay. Duration of these flares is typically within few minutes to hours (Doyle et al., 2018). The presence of flares indicates magnetic activity in the stellar magnetosphere. Since the magnetic activity is strongly related to the strength of the magnetic field, fast-rotating young stars are typically more active. The magnetic activity also changes over time, so one would expect the frequency and power of flares to change as well. However, it is hard to investigate these changes in G-M type stars due to their long lifetime and slow evolution on the main sequence.

There are several possible ways to estimate the age of such late type stars, such as method based on the rotation period (gyrochronology, Barnes 2007), or spatial motion (kinematic group membership). But probably the most effective and affordable one is by means of estimating an age of a parent cluster. Therefore, the aim of this study is to evaluate the effect of stellar age on the activity of flare stars in Galactic open clusters.

Open clusters are groups of stars loosely held together by their mutual gravitational attraction and formed from the same interstellar molecular cloud. Therefore the age, distance and chemical properties of members in open clusters are approximately identical. This means that the chemical composition of open

clusters also reflects a real composition of individual members (Netopil et al., 2015).

In this paper, we combine the high temporal resolution light curves acquired by *TESS* satellite with the cluster membership data based on *Gaia* DR2 (Cantat-Gaudin et al., 2018a). This allows us to reliably select the members of Galactic open clusters with known age and distance, examine their variability and extract the ones with prominent stellar flares for further statistical analysis.

The paper is organized as follows. In Section 2, we briefly describe a methodology of selection of cluster members having *TESS* light curves available and the methodology of flare detection and characterisation. In Section 3, we present statistical analysis of detected flares. A conclusion of this study is given in Section 4.

2. Selection of flare stars

The first step of our work was the selection of stars belonging to open clusters and being *TESS* targets with high temporal resolution light curves available. To do it, we took a list of Galactic open clusters from Cantat-Gaudin et al. (2018a) and selected the ones closer than 2 kpc from the Sun. Such distance cut-off allows us to discard the clusters where G0V and later type of dwarfs, the most interesting for flare stars statistics, are below *TESS* detection limit due to their intrinsic faintness. Then, we used the list of members from Cantat-Gaudin et al. (2018a) for all selected clusters and checked the availability of *TESS* light curves for every object. As a result, we acquired 859 light curves for stars located in 136 open clusters. We also considered nearby cluster Melotte 111 (or Coma Star cluster) with 98 light curves available. Basic parameters of these clusters (age, number of members, distance) are summarized in Table 2, along with the numbers of available light curves and detected flares for them.

As the next step, we performed a visual inspection of all these light curves and discarded the ones without any flares, thus significantly reduced the number of light curves. Then we performed an automated detection and characterisation of the flares using `FLATW'RM` (FLAre deTectioN With Ransac Method) code (Vida & Roettenbacher, 2018). This code uses a machine-learning algorithm to give a robust model of the light curves in order to detect flare events and uses a voting system implemented to keep false positive detections to a minimum. `FLATW'RM` detects flares and reports the times the flare starts and ends, the time of maximum flux and the maximum percent increase of flux over the light curve around the flare, along with estimated flare energy, either from the raw light curve or by fitting an analytic flare model. It also gives a crude estimation of the period of the underlying light curve, assuming that it is periodic and using the position of the strongest peak in the Lomb-Scargle periodogram.

Figure 1 illustrates the Hertzsprung–Russell (H–R) diagram of several clusters from our analysis with locations of found flare stars and cluster members

with *TESS* high temporal resolution light curves available. We also plotted the isochrones calculated using PARSEC (Chen et al., 2014) for metallicities $Z=0.017$ and $Z=0.020$ updated to the latest transmission curve calibrated on *Gaia* DR2 data (Evans et al., 2018)¹. The figure shows that the majority of the found flare stars have masses between $0.5 - 2 M_{\odot}$. Figure 2 presents several clusters where no flare star was detected.

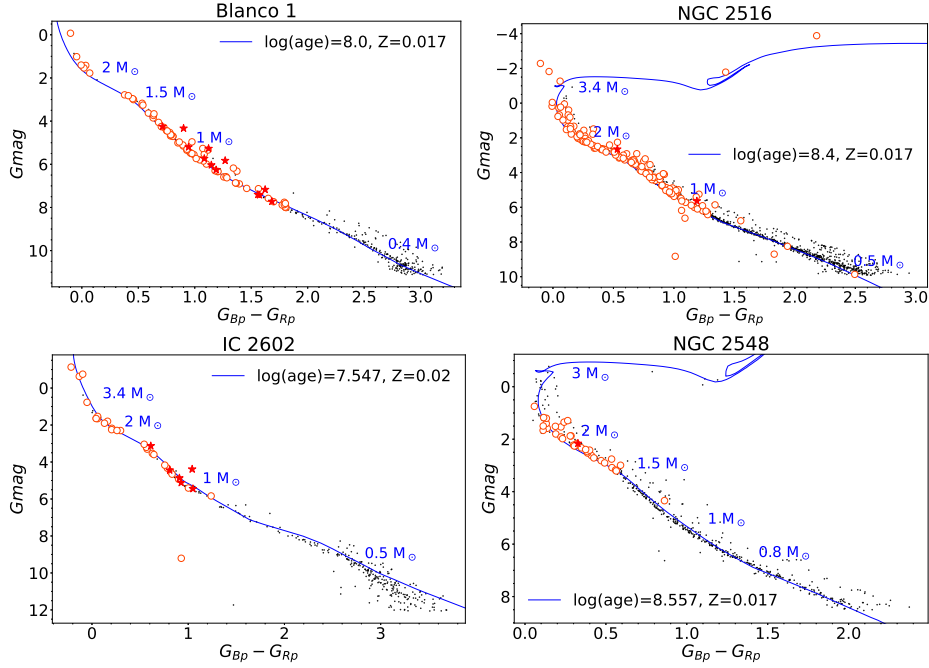


Figure 1.: H–R diagram for Blanco 1, NGC 2516, IC 2602 and NGC 2548 clusters based on *Gaia* photometric data and compared with the PARSEC isochrones (see text for details). Red circles are TESS objects without flares, while red stars – with flares.

It is important to note that two of the clusters – ASCC 19 and Gulliver 6 – are located very close on the sky and also on similar distance. Due to this, the star HD 290527 is considered as a member of both clusters simultaneously (Cantat-Gaudin et al., 2018a). On the other hand, the star HD 290674 is considered to be a member of Gulliver 6 according to Cantat-Gaudin et al. (2018a) and a member of ASCC 19 – according to Cantat-Gaudin et al. (2018b). This leads to the ambiguities of the ages for these stars.

¹PARSEC isochrones in *Gaia* DR2 passbands are available at <http://stev.oapd.inaf.it/cgi-bin/cmd>

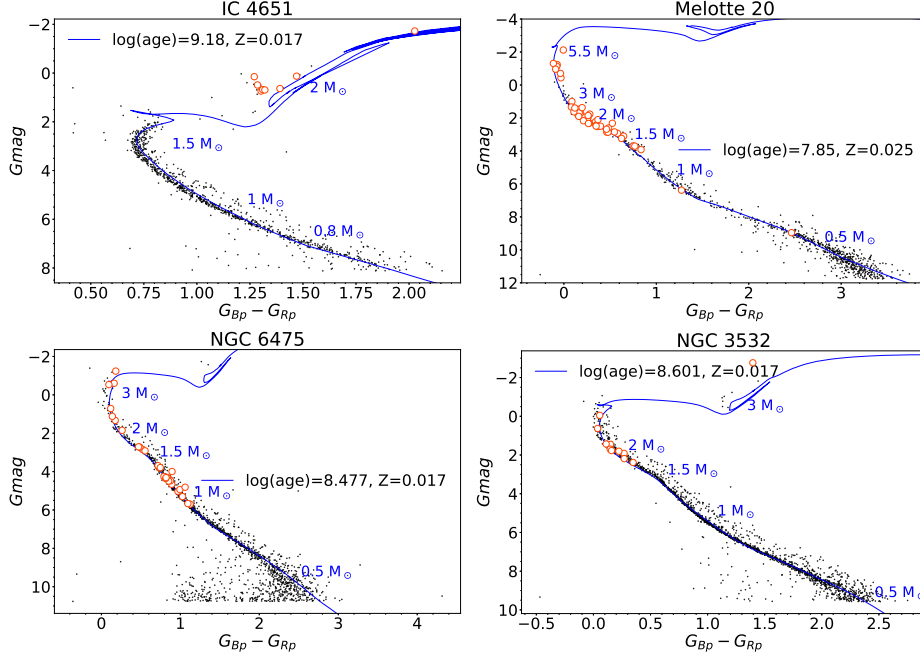


Figure 2.: H–R diagram for IC 4651, Melotte 20, NGC 6475 and NGC 3532 clusters based on *Gaia* photometric data and compared with PARSEC isochrones. Red circles are TESS objects without flares.

A full list of flare stars we found is given in Table 1. Specified spectral types are taken from the SIMBAD database. As well as spectral types are not known for all star, 5th column gives effective temperatures T_{eff} according to *Gaia* DR2 (Gaia Collaboration et al., 2016, 2018). The 6th column contains the number of flares detected during the period of observations, while the last column is the period of rotation. Both values were estimated using FLATW’RM. The code estimates the periods based on the strongest peak in the Lomb-Scargle periodogram, and thus may provide bogus or halved periods in some cases. Therefore we also checked all measured period values by eye and corrected them where necessary so that they clearly correspond to the main periodic components. For one of the stars in our sample – EO Cha – the determined period is close to total duration of observations and might be caused by systematic effects in the data. However, our value is in agreement with the one reported in AAVSO VSX database², and thus we decided to keep it in our analysis.

²American Association of Variable Star Observers Variable Star Index is available in <https://www.aavso.org/vsx/>

Table 1.: List of flare stars selected in this work. Spectral types are taken from the SIMBAD database. T_{eff} is temperature taken from *Gaia* DR2 (Gaia Collaboration et al., 2016, 2018). Numbers of flares (N_{fl}) and rotation periods (P_{rot}) are estimated as described in Section 2.

Cluster	<i>Gaia</i> DR2 ID	Name	Sp. type	T_{eff}	N_{fl}	P_{rot} [days]
ASCC 19 ^a	3220225706694228096	HD 290527	A3	7470	3	6.69
ASCC 21	3236083864117227264	HD 36030	B9V	9777	1	–
Alessi 13	4861032719915154176	–	–	3922	1	2.27
Alessi 13	4860643905115917312	CD-36 1309	F8	5922	4	1.20
Alessi 13	4854771001195018752	CD-37 1263	G5	5637	5	2.06
Blanco 1	2320842551136015872	–	K4-5V	4954	1	5.55
Blanco 1	2320838943363481984	–	K5.3	4296	5	2.17
Blanco 1	2320874054720005248	–	F-K	4131	1	5.84
Blanco 1	2320933157765826560	CD-30 19826	F-K	5496	3	1.03
Blanco 1	2320596123092361088	–	F-K	5031	2	0.34
Blanco 1	2320862200611102848	–	F-K	5499	7	2.34
Blanco 1	2320860074602561408	–	F-K	4987	2	0.41
Blanco 1	2320881377640416640	–	F-K	4301	3	1.18
Blanco 1	2320869897192498432	CD-30 19800	F8	6070	1	2.65
Blanco 1	2320795031617857152	–	F-K	4977	3	3.15
Blanco 1	2320872650266808192	–	F-K	4873	3	4.14
Blanco 1	2320847533296914560	–	F-K	4579	4	6.47
Gulliver 6	3217389348947353728	HD 290674	A0	9034	1	6.41
IC 2602	5239841340702856960	HD 93405	F3/5V	6383	5	0.75
IC 2602	5239660372284736896	–	–	5362	4	3.80
IC 2602	5239626420542800512	V570 Car	G9Ve	5071	2	1.21
IC 2602	5239498744077038976	HD 310131	G6V	5583	2	0.62
IC 2602	5253546997989686912	HD 307772	G7Ve	5161	3	0.32
IC 2602	5251470948229949568	–	–	5768	1	3.02
Mamajek 1	5209082129256747008	EG Cha	K4Ve	4365	8	4.35
Mamajek 1	5209135352491538432	EO Cha	M0	4089	3	22.26
Mamajek 1	5209038423669475712	EL Cha	M2	3526	6	1.84
Mamajek 1	5215178848217868288	–	M3	3673	2	2.96
Mamajek 1	5209118305765620096	EQ Cha	M3	3367	14	1.25
NGC 1662	3294740954034417792	–	–	4454	1	–
NGC 1662	3295557822453425792	–	–	4447	1	–
NGC 2422	3030122727536859008	–	–	4826	1	6.97
NGC 2451A	5538425822158520576	–	–	3761	2	0.52

Table 1.: Continued.

Cluster	<i>Gaia</i> DR2 ID	Name	Sp. type	T_{eff}	N_{fl}	P_{rot} [days]
NGC 2451B	5538722449779987072	–	–	4438	1	0.50
NGC 2451B	5538818176010943488	–	–	5326	5	1.95
NGC 2516	5290732721731067136	–	–	4882	2	2.07
NGC 2516	5290719115275116288	–	A7III	7038	1	1.02
NGC 2548	3064487035741401984	–	A2/4	8160	1	0.80
NGC 6281	5976330585929049600	V948 Sco	ApSi	8553	1	1.64
Platais 8	5305106426785000448	CPD-55 1885	G5V	5511	3	0.92
Platais 8	5303854293507991680	HD 78027	A1V	8521	1	3.31
Platais 8	5310127655876188544	CD-55 2543	G8V	5406	1	3.84
Platais 9	5427272240335502848	–	–	3316	1	–
Platais 9	5427469843189439360	–	–	3759	1	3.49
Platais 9	5423893028786327168	HD 80484	A1V	8884	1	0.69
Platais 9	5327362504934951040	–	–	5732	1	1.23
Melotte 111	3959841756787187456	–	M3.0	3395	2	–
Melotte 111	4002505586787874560	–	M3.5	3421	5	–
Melotte 111	3960130997064687616	–	M3.5	3783	5	–
Melotte 111	4008440617411511552	Sand 64	M3.9	3921	2	–
Melotte 111	4009049575054518400	–	M2.4	3789	1	–
Melotte 111	3954074577781239936	Sand 63	M2.2	3628	2	–
Melotte 111	4009398051520907008	HK Com	M4.2	3848	4	–
Melotte 111	4009440382718550272	–	M3.7	3937	1	–
Melotte 111	4009455054326835968	–	M2.6	3899	3	3.39
Melotte 111	4002543008838439296	–	M3.8	–	1	1.87

^a HD 290527 is probably a member of Gulliver 6 cluster, see text for details.

3. Discussion

3.1. Spectral distribution of selected flare stars

Most of stellar flares occur in red dwarfs, more often in cool M dwarfs (Güdel & Nazé, 2009). Data from the first data release of *TESS* mission confirm it: among 1228 flare stars detected by Günther et al. (2020), 673 objects were classified as M dwarfs. Our data is in a good agreement with these previous studies. Figure 3 presents the histogram of the number of flare stars versus spectral type based

on the data from Table 1. Indeed, the majority of selected objects are stars of late spectral type.

M-type dwarfs are very faint objects, an absolute visual magnitude of M0V star is $M_V = 9.0$ mag and for M5V is $M_V = 14.6$ mag (Schmidt-Kaler, 1982). In our work, we found M-type flare stars only in two nearby open clusters, Mamajek 1 (distance $D = 98.72$ pc) and Mel 111 ($D = 96$ pc). Figure 4 shows that flares were detected for almost 50 % of all considered M-type stars.

Stellar flares have been previously detected in some hot B–A type stars (in optical range (Schaefer, 1989; Balona, 2012); in X-ray (Schmitt et al., 1994; Yanagida et al., 2007)). Several objects from our sample are also early spectral type flare stars. In this study, we detected flares in seven A-type stars (Table 1). One of them – V948 Sco – belongs to α^2 CVn variables, chemically peculiar main sequence stars with a strong magnetic field. For all other A-type stars, the variability was registered for the first time.

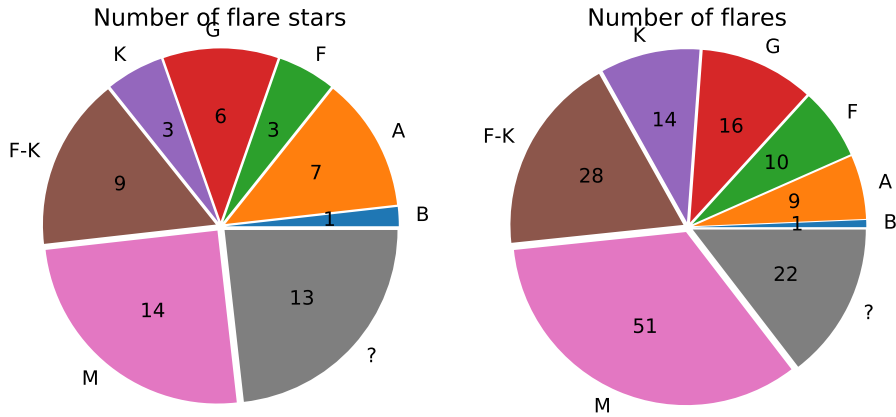


Figure 3.: Distribution of flare stars in our sample (left panel), and detected flares from them (right panel) over spectral types from Table 1.

Another selected hot star is HD 36030, classified as B9V star by Houk & Swift (1999). Figure 5 presents the light curve of HD 36030 and a clearly detected flare. As in the case with A-type stars, we did not find any previously published work about variability of this object.

Flares of hot stars may be explained through binarity. As mostly A-type stars are probably double or multiple systems, it is natural to expect flares to originate in a cool companion. However, Balona (2019) gives strong arguments that flares originate in A-type stars themselves. These arguments are based on ratio of luminosity: an M- or K-type companion is about 50–100 times less luminous than an A-type star. The flare amplitude originating in a cool companion should therefore, on average, appears to be about 50–100 times smaller when

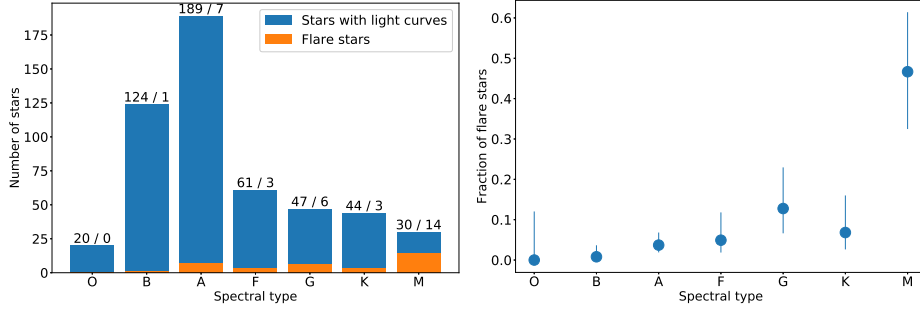


Figure 4.: Left panel – histogram of total number of light curves and flare stars among them per spectral type (where it was reliably determined). Right panel – fraction of flare stars as a function of spectral type. Error bars depict the 90 % confidence intervals for the corresponding values.

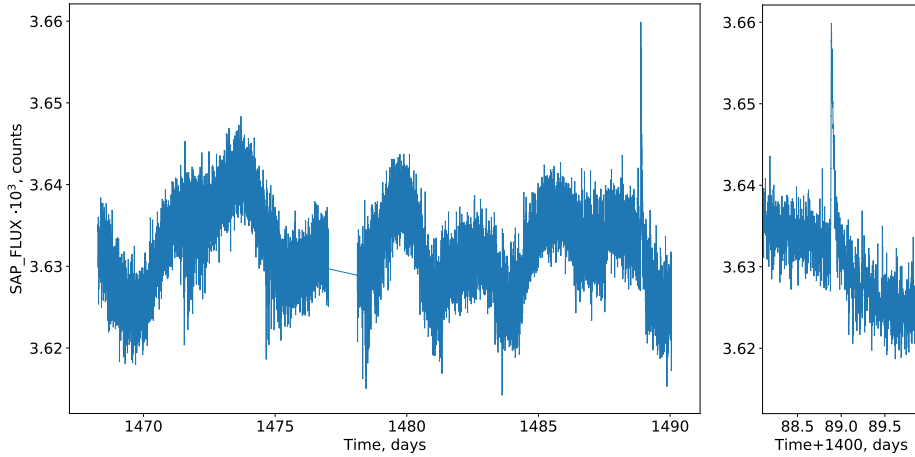


Figure 5.: Light curve of HD 36030 (B9V) with a clearly detected flare.

observed together with A-type star. However, flares detected in hot stars are comparable with their luminosity (Balona, 2019). Flares in A-stars of our sample have amplitudes of couple percents only, and may be in principle, at least partially, explained by a giant flares with amplitudes of several magnitudes on a late type companion, like the ones reported in Beskin et al. (2017). However, such giant flares are quite rare, and are not detected in our sample of M stars. Thus, we may not expect them to be abundant among the possible components of our A stars too. On the other hand, magnetic interaction with the companion star could provide a plausible explanation for observed flaring activity (Balona, 2012).

According to Balona (2015), the relative number of flare stars from cool M dwarfs to hot A-type stars is probably the same. According to our results (see Figure 4), the fraction of flare stars among G-M type objects is significantly higher.

There are already no published data about spectral types for fourteen flare stars in our sample (see Table 1, Figure 3). Future spectral observations of these objects could improve statistics.

3.2. Statistical analysis of flares

The mean duration of flares for stars in Table 1 varies between 20 to 70 minutes. This timescale is comparable with the one of GOES X-class solar flares, quite often visible in the white light (Harra et al., 2016). Only 4 stars among 56 analysed in this paper have more than 5 flares during period of *TESS* observations. One of them is EQ Cha (M3) – variable star of Orion type, close visual binary (Sicilia-Aguilar et al., 2009) and a member of Mamajek 1 open cluster. It showed 14 flares during 25 days of observations. Flaring activity of EQ Cha star has been already reported in X-ray by López-Santiago et al. (2010).

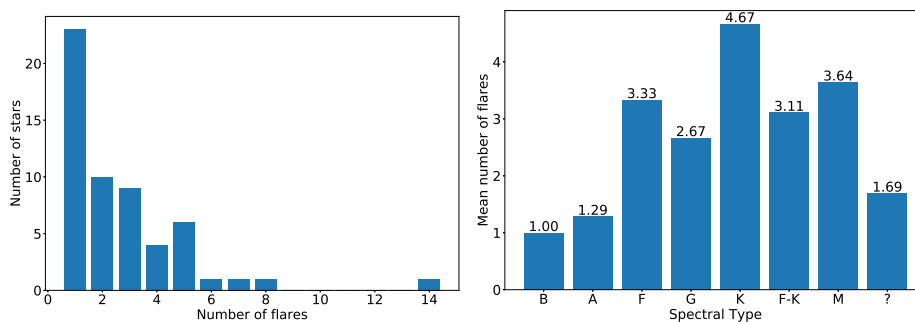


Figure 6.: Histogram of number of flares detected from a star (left panel). Mean number of flares per star depending on spectral class from Table 1 (right panel).

On the right panel of Figure 6, we can see that K-type stars have the largest mean flare activity in their light curves. The largest number of flares belongs to M-type stars, which are all earlier than M4 in our sample (see Table 1) and thus are cool stars with convective envelope, as well as to K-type stars. These K and early M dwarfs, which retain a radiative core, presumably all have a quite efficient solar-type dynamo powering their magnetic activity. On the other hand, M dwarfs of spectral type later than M4 are fully convective, and have a turbulent dynamo at work (Wright et al., 2011; Stassun et al., 2011). The activity of G-type stars could be also associated with their convective envelopes. Figure 7 shows the dependence of a flare rate on the star’ effective temperature.

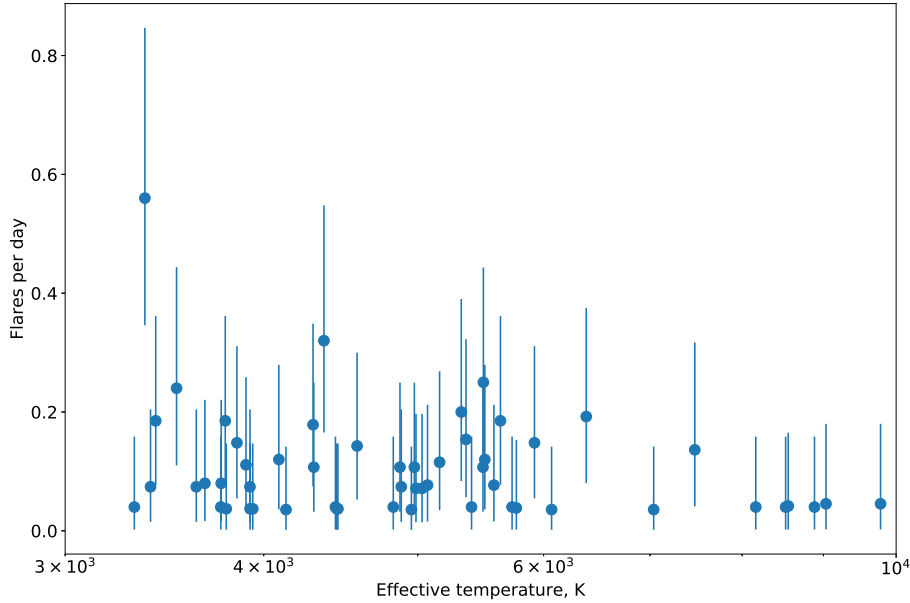


Figure 7.: Rate of stellar flares as a function of temperature. Error bars correspond to the 90 % confidence intervals computed assuming Poissonian statistics.

The appearance of spots on the surface of stars and stellar flares are clear evidences of stellar magnetism. The connection between stellar rotation and magnetic activity has been studied in many works. For example, Guo et al. (2014) analyzed X-class flares of Sun during the 22nd and 23rd solar cycles, and they found that flares closely follow the same 11 year cycle as sunspots. Doyle et al. (2018) did not find any correlation between the rotation phase and the number of flares for M-dwarfs based on *Kepler* data, but clearly detected the decrease of flaring activity for rotational periods longer than 10 days. The same dependence is seen in X-ray activity-rotation relation (Pizzocaro et al., 2019).

Figure 8 shows the dependence of flaring activity in our sample, estimated as a number of flares per day of observations on the rotational period of the star. We may spot just a slight trend of decreasing activity towards slower rotation. While we may not stress it as a significant dependence, let's note that it agrees with the expected behaviour of magnetic activity being as a result of internal magnetic dynamo, arising from the combination of stellar differential rotation and convection in the sub-photospheric layers.

3.3. Age distribution of flare stars

Mamajek 1 is the youngest open cluster among ones where we have found flare stars. Its age is 6 Myr. The oldest one where we have found flare is the cluster

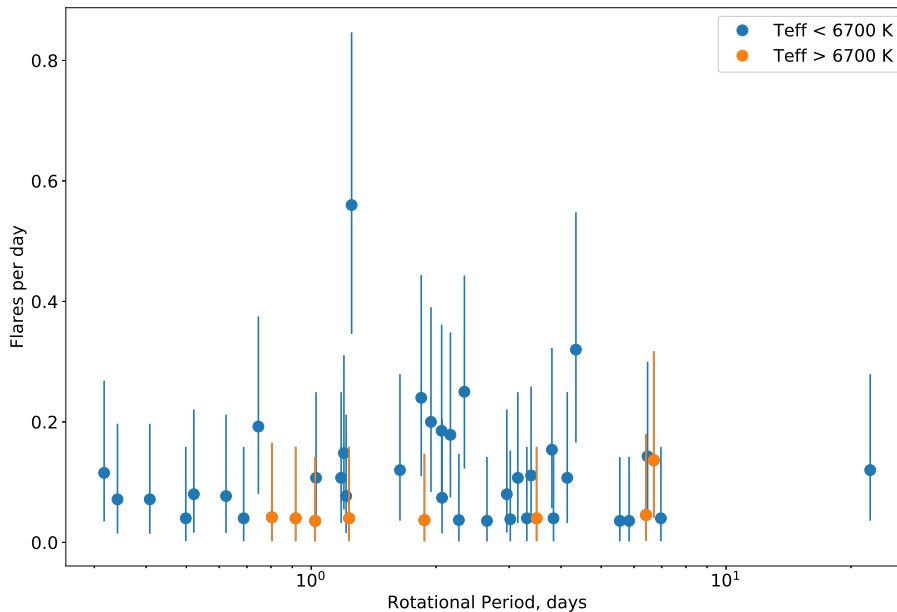


Figure 8.: Rate of stellar flares as a function of rotational period. Error bars correspond to the 90 % confidence intervals computed assuming Poissonian statistics.

NGC 1662. Its age 1 Gyr is comparable with the maximal lifetime of open clusters (Fujii & Portegies Zwart, 2016). The ages of all examined clusters span from 1 Myr (IC 5146, the youngest) to 4.3 Gyr (NGC 188, the oldest one). Therefore, the collected data do not allow us to confidently determine the age after which flare stars in clusters disappear.

Figure 9 shows how flaring activity changes with stellar age. The resulting distribution is mostly flat, with just a single deviating point with 14 detected flares from EQ Cha star. We therefore cannot reliably see any dependence of flaring activity on the age.

4. Conclusion

In order to study the dependence of stellar flaring activity on the age, we performed a uniform study of large sample of Galactic open clusters from catalogue of Cantat-Gaudin et al. (2018a). We analyzed a total amount of 957 high-cadence *TESS* light curves for stars from 136 open clusters. By visual inspection and later using a `FLATWRM` code, it finally turned out that 56 of them are flare stars with 151 flares detected.

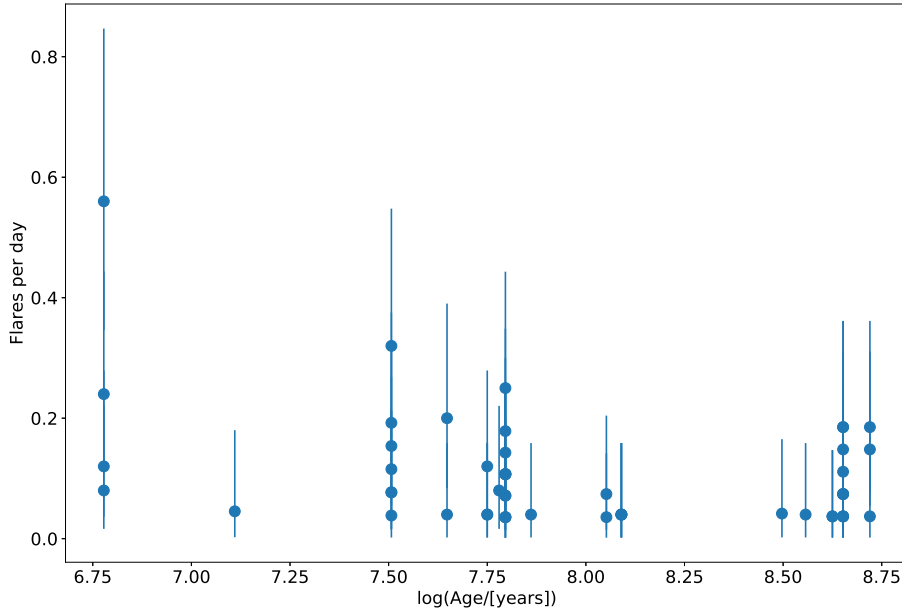


Figure 9.: Rate of flares as a function of age. The rate was estimated from the number of detected flares divided by the effective light curve duration for every individual star. Error bars depict the 90 % confidence intervals computed assuming Poissonian statistics.

41.5 % of all flares were detected in cool stars. 8 flare stars belong to B- and A-type stars. The flare activity was detected for the first time for all these hot stars. The flare with the largest amplitude appears on M3-type EQ Cha star. From total amount of detected flare stars in our sample, 25 % of them are members of spectral type M, 11 % are G-type stars and about 5.5% F- and K-type stars.

Statistical analysis of detected flare stars shows that:

- among cool stars up to 50 % have flares, while among A-type stars it is only 3.7 %.
- duration of flares varies between 20 to 70 minutes.
- we found no significant link between rotation rate and flare activity.
- we found no significant link between stellar age and flare activity.

In the study, we used *TESS* data recorded at 2-minute cadence light curves, which are an invaluable source of photometric data for various kinds of stars, pre-selected before the mission start. Although *TESS* full-frame images have a

much wider coverage of objects, limited only by crowding, their utility for studies of stellar flares are significantly lower due to much worse temporal resolution (30 minutes, which is comparable to total duration of flares). As TESS mission now continues into its third year, re-visiting the regions of the sky already observed before, we expect to get more data for the objects we studied in the future. In the future, we plan to do spectroscopic observations of B-A type stars with flare activity we had detected, in order to investigate their possible multiplicity.

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³<https://www.cosmos.esa.int/gaia>

⁴<https://www.cosmos.esa.int/web/gaia/dpac/consortium>

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A. Clusters with *TESS* light curves available

Table 2.: Complete list of Galactic open clusters investigated in this work, with ages τ taken from various sources (see the last column), distances D and number of members N_{mem} taken from Cantat-Gaudin et al. (2018a), number of available light curves N_{lc} in the *TESS* database and number of flare stars N_{fs} estimated in this work.

Name	$\log(\tau/[\text{yr}])$	D [pc]	N_{mem}	N_{lc}	N_{fs}	Reference for age
Alessi 13	8.72	104.34	48	10	3	Yen et al. (2018)
Alessi 19	7.38	584.80	74	1	0	Bossini et al. (2019)
Alessi 20	8.22	432.71	124	2	0	WEBDA ^a
Alessi 21	7.816	581.40	137	2	0	Bossini et al. (2019)
Alessi 24	7.945	485.67	163	1	0	Bossini et al. (2019)
Alessi 3	8.90	280.43	178	8	0	Yen et al. (2018)
Alessi 37	8.125	722.54	142	1	0	Bossini et al. (2019)
Alessi 43	7.48	968.99	318	2	0	WEBDA
Alessi 5	7.723	399.84	302	3	0	Bossini et al. (2019)
Alessi 9	8.42	207.38	194	5	0	Yen et al. (2018)
Alessi Teutsch 12	7.977	612.75	44	1	0	Bossini et al. (2019)
Alessi Teutsch 5	7.02	900.09	158	1	0	WEBDA
Alessi Teutsch 8	8.65	1034.13	341	2	0	WEBDA
ASCC 105	7.994	560.85	127	1	0	Bossini et al. (2019)
ASCC 10	8.599	685.40	71	1	0	Bossini et al. (2019)
ASCC 12	8.42	1062.70	162	1	0	WEBDA
ASCC 16	7.047	352.36	226	9	0	Bossini et al. (2019)
ASCC 19	7.086	361.27	188	3	1	Bossini et al. (2019)
ASCC 21	7.032	348.92	131	3	1	Bossini et al. (2019)
ASCC 32	7.404	813.01	259	2	0	Bossini et al. (2019)
ASCC 58	7.826	485.44	137	2	0	Bossini et al. (2019)
ASCC 79	6.86	850.34	129	1	0	Bossini et al. (2019)
ASCC 85	7.42	893.66	119	1	0	WEBDA
Aveni Hunter 1	8.26	425.89	82	3	0	WEBDA
Basel 8	8.102	1600.00	33	1	0	WEBDA
Berkeley 86	7.116	1792.11	31	2	0	WEBDA
Berkeley 87	7.152	1745.20	131	1	0	WEBDA
BH 164	7.81	423.01	209	2	0	Bossini et al. (2019)
BH 221	8.01	1131.22	162	1	0	WEBDA
BH 23	7.14	443.07	94	1	0	WEBDA
BH 56	7.24	927.64	122	3	0	WEBDA

Table 2.: Continued.

Name	$\log(\tau/[\text{yr}])$	D [pc]	N_{mem}	N_{ic}	N_{fs}	Reference for age
BH 99	7.908	449.44	389	1	0	Bossini et al. (2019)
Biurakan 2	7.011	1845.02	68	2	0	WEBDA
Blanco 1	7.975	237.53	381	82	12	Bossini et al. (2019)
Bochum 13	6.823	1763.67	73	1	0	WEBDA
Collinder 106	6.74	1574.80	114	1	0	WEBDA
Collinder 107	7.00	1620.75	159	1	0	WEBDA
Collinder 132	7.08	666.22	99	1	0	WEBDA
Collinder 135	7.407	305.06	352	1	0	WEBDA
Collinder 197	7.128	967.12	243	4	0	WEBDA
Collinder 258	7.834	1308.90	124	1	0	Bossini et al. (2019)
DBSB 104		1101.32	27	1	0	
Dolidze 8		1007.05	43	3	0	
FSR 0771	8.195	1618.12	48	1	0	Bossini et al. (2019)
FSR 0951	8.589	1808.32	195	1	0	Bossini et al. (2019)
Gulliver 11		942.51	64	1	0	
Gulliver 21	8.472	664.89	126	1	0	Bossini et al. (2019)
Gulliver 6		422.48	343	6	2	
Gulliver 9	7.154	503.78	265	1	0	Bossini et al. (2019)
Haffner 13	7.497	574.05	210	1	0	Bossini et al. (2019)
Harvard 10	7.899	705.22	164	1	0	Bossini et al. (2019)
Harvard 5	7.812	1308.90	52	1	0	Bossini et al. (2019)
IC 1396	7.054	938.09	460	2	0	WEBDA
IC 2391	7.561	151.93	224	7	0	Bossini et al. (2019)
IC 2395	7.223	725.16	297	4	0	WEBDA
IC 2488	8.2	1392.76	435	1	0	Bossini et al. (2019)
IC 2602	7.547	152.42	311	33	6	Bossini et al. (2019)
IC 2714	8.55	1390.82	934	1	0	Bossini et al. (2019)
IC 4651	9.057	946.97	854	8	0	WEBDA
IC 5146	6.00	824.40	108	2	0	WEBDA
King 6	8.582	742.94	234	1	0	Bossini et al. (2019)
L 1641S		437.06	101	1	0	
Mamajek 1	6.778	98.72	20	9	5	WEBDA
Mamajek 4	8.824	449.64	199	3	0	Bossini et al. (2019)
Melotte 111	8.652	96 ^b	859	98	10	WEBDA
Melotte 20	7.938	176.43	764	50	0	Bossini et al. (2019)
Melotte 22	7.937	136.13	992	5	0	Bossini et al. (2019)
Muzzio 1		1926.78	40	1	0	

Table 2.: Continued.

Name	$\log(\tau/[\text{yr}])$	D [pc]	N_{mem}	N_{lc}	N_{fs}	Reference for age
NGC 129	7.886	1956.95	392	1	0	WEBDA
NGC 1333		299.04	50	3	0	
NGC 1545	8.019	731.53	139	1	0	Bossini et al. (2019)
NGC 1579		552.49	56	2	0	
NGC 1662	8.957	416.67	238	24	2	Bossini et al. (2019)
NGC 188	9.632	1972.39	864	3	0	WEBDA
NGC 1901	8.918	424.09	75	15	0	Bossini et al. (2019)
NGC 2232	7.87	326.05	198	2	0	WEBDA
NGC 2244	6.896	1620.75	623	13	0	WEBDA
NGC 2287	8.486	735.29	645	1	0	Bossini et al. (2019)
NGC 2323	7.975	1003.01	866	1	0	Bossini et al. (2019)
NGC 2354	8.126	1328.02	276	2	0	WEBDA
NGC 2422	8.104	483.09	442	43	1	Bossini et al. (2019)
NGC 2423	8.991	956.94	430	1	0	Bossini et al. (2019)
NGC 2437	8.390	1658.37	1797	1	0	WEBDA
NGC 2451A	7.647	193.61	337	26	1	Bossini et al. (2019)
NGC 2451B	7.592	367.78	298	15	2	Bossini et al. (2019)
NGC 2516	8.4	413.74	798	144	1	Bossini et al. (2019)
NGC 2546	7.962	967.12	165	1	0	Bossini et al. (2019)
NGC 2547	7.432	391.70	233	4	0	Bossini et al. (2019)
NGC 2548	8.557	775.80	479	37	1	WEBDA
NGC 3114	8.093	1048.22	1296	6	0	WEBDA
NGC 3532	8.601	484.03	1889	11	0	Bossini et al. (2019)
NGC 3680	9.077	1071.81	100	8	0	WEBDA
NGC 5662	7.968	776.40	255	1	0	WEBDA
NGC 5822	8.95	842.46	667	1	0	Bossini et al. (2019)
NGC 6124	8.147	642.26	1327	1	0	WEBDA
NGC 6178	7.248	909.92	45	1	0	WEBDA
NGC 6193	6.775	1231.53	465	2	0	WEBDA
NGC 6231	6.843	1697.79	653	11	0	WEBDA
NGC 6242	7.767	1324.50	523	1	0	Bossini et al. (2019)
NGC 6250	7.415	1012.15	85	1	0	WEBDA
NGC 6281	8.497	533.90	513	6	0	WEBDA
NGC 6475	8.477	279.96	941	29	0	Bossini et al. (2019)
NGC 6811	8.936	1149.43	306	4	0	Bossini et al. (2019)
NGC 6866	8.89	1457.73	104	1	0	Bossini et al. (2019)
NGC 6871	6.958	1945.53	594	11	0	WEBDA

Table 2.: Continued.

Name	$\log(\tau/[\text{yr}])$	D [pc]	N_{mem}	N_{lc}	N_{fs}	Reference for age
NGC 6910	7.127	1834.86	159	2	0	WEBDA
NGC 6913	7.111	1808.32	89	2	0	WEBDA
NGC 6940	9.01	1055.97	593	4	0	Bossini et al. (2019)
NGC 7092	8.491	299.67	161	6	0	Bossini et al. (2019)
NGC 7129		920.81	36	1	0	
NGC 7243	8.006	896.06	313	2	0	Bossini et al. (2019)
NGC 743		1135.07	62	1	0	
NGC 752	9.05	446.63	240	9	0	WEBDA
Pismis 5	7.197	974.66	92	1	0	WEBDA
Platais 3	8.319	177.94	90	12	0	Bossini et al. (2019)
Platais 8	7.75	134.86	211	13	3	WEBDA
Platais 9	7.894	183.05	128	18	4	Bossini et al. (2019)
Pozzo 1	7.117	350.51	390	3	0	Bossini et al. (2019)
Roslund 3	7.643	1709.40	112	1	0	Bossini et al. (2019)
RSG 5	7.70	339.56	173	1	0	Röser et al. (2016)
RSG 7	8.30	427.90	133	1	0	Röser et al. (2016)
RSG 8	8.50	452.49	211	1	0	Röser et al. (2016)
Ruprecht 91	6.90	1078.75	214	2	0	WEBDA
Stephenson 1	7.435	359.84	97	3	0	Bossini et al. (2019)
Stock 12	8.188	442.28	109	2	0	Bossini et al. (2019)
Stock 1	8.676	408.66	183	2	0	Bossini et al. (2019)
Stock 23	7.973	619.96	89	3	0	Bossini et al. (2019)
Stock 2	8.23	378.64	1190	1	0	WEBDA
Stock 5	7.73	987.17	115	1	0	WEBDA
Stock 7	7.214	693.96	141	1	0	WEBDA
Teutsch 38	8.14	665.78	109	1	0	WEBDA
Trumpler 10	7.74	437.25	478	2	0	Bossini et al. (2019)
Trumpler 2	7.962	698.81	179	2	0	Bossini et al. (2019)
Trumpler 3	7.83	684.46	222	2	0	WEBDA
Turner 5	8.49	421.41	28	1	0	WEBDA
vdBergh 1	8.025	1912.05	73	1	0	WEBDA

^a WEBDA database <https://webda.physics.muni.cz/>^b Distance for Melotte 111 are taken from WEBDA.