

AN ANALYSIS OF THE ENVIRONMENTS OF FU ORIONIS OBJECTS WITH *HERSCHEL*¹

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ABSTRACT

We present *Herschel*-HIFI, SPIRE, and PACS 50-670 μm imaging and spectroscopy of six FU Orionis-type objects and candidates (FU Orionis, V1735 Cyg, V1515 Cyg, V1057 Cyg, V1331 Cyg, and HBC 722), ranging in outburst date from 1936-2010, from the “FOOSH” (FU Orionis Objects Surveyed with *Herschel*) program, as well as ancillary results from Spitzer-IRS and the Caltech Submillimeter Observatory. In their system properties (L_{bol} , T_{bol} , line emission), we find that FUors are in a variety of evolutionary states. Additionally, some FUors have features of both Class I and II sources: warm continuum consistent with Class II sources, but rotational line emission typical of Class I, far higher than Class II sources of similar mass/luminosity. Combining several classification techniques, we find an evolutionary sequence consistent with previous mid-IR indicators. We detect [O I] in every source at luminosities consistent with Class 0/I protostars, much greater than in Class II disks. We detect transitions of ¹³CO (J_{up} of 5 to 8) around two sources (V1735 Cyg and HBC 722) but attribute them to nearby protostars. Of the remaining sources, three (FU Ori, V1515 Cyg, and V1331 Cyg) exhibit only low-lying CO, but one (V1057 Cyg) shows CO up to $J = 23 \rightarrow 22$ and evidence for H₂O and OH emission, at strengths typical of protostars rather than T Tauri stars. Rotational temperatures for “cool” CO components range from 20-81 K, for $\sim 10^{50}$ total CO molecules. We detect [C I] and [N II] primarily as diffuse emission.

Subject headings: stars: pre-main sequence — stars: variables: T Tauri — ISM: jets and outflows — submillimeter: ISM — stars: individual (HBC 722, FU Orionis, V1057 Cyg, V1735 Cyg, V1331 Cyg, V1515 Cyg)

1. INTRODUCTION

FU Orionis-type objects (hereafter, FUors) are a class of low-mass pre-main sequence objects named after the archetype FU Orionis (hereafter, FU Ori), which produced a 6 magnitude outburst at *B*-band in 1936 and has remained close to peak brightness ever since.¹

About five more sources have been observed to flare with lightcurves broadly resembling that of FU Ori (5 mag optical flare over ~ 1 yr). Similar spectral characteristics (broad blueshifted emission lines, IR excess, near-IR CO overtone absorption) have been used to identify ~ 10 additional FUor-like objects (Hartmann & Kenyon 1996; Reipurth & Aspin 2010); a number of additional candidates have been identified (e.g. Quanz et al. 2007b; Semkov & Peneva 2010; Reipurth et al. 2012; Fischer et al. 2012, 2013). Paczynski (1976) first pro-

posed that FUors are the result of a sudden cataclysmic accretion of material from a reservoir that had built up in the circumstellar disk surrounding a young stellar object (see also Lin & Papaloizou 1985; Hartmann & Kenyon 1985). Models of FUor outbursts indicate that over only a few months the accretion rate rises from the typical rate for a T Tauri star ($\dot{M} \lesssim 10^{-7} M_{\odot} \text{ yr}^{-1}$) up to $10^{-4} M_{\odot} \text{ yr}^{-1}$, and then decays over an e-fold time of ~ 10 -100 yr (Bell & Lin 1994). Over the entire outburst the star could accrete $\sim 0.01 M_{\odot}$ of material, roughly the mass of a typical T Tauri disk (Andrews & Williams 2005). Pre-outburst optical spectra available for two FUors (V1057 Cyg and HBC 722, below) are consistent with T Tauri stars (Herbig 1977; Miller et al. 2011); they now have optical spectral types of F and K supergiants, respectively. This change in spectral type is attributed to the large increase in continuum emission in the optical/near-IR, typically dominated by the newly heated inner disk annuli rather than the central star, during outburst.

¹ *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

Only in recent years, it has become possible to produce full spectral energy distributions (SEDs) of FUors, the closest of which are several times farther away than the best-studied low mass star forming regions, first with the Infrared Space Observatory (ISO, 20–200 μm ; Lorenzetti et al. 2000; Ábrahám et al. 2004). More recently, the Spitzer Space Telescope (Werner et al. 2004) and the Herschel Space Observatory (Pilbratt et al. 2010) cover an extended spectral range (5.3–670 μm) with a greater sensitivity due in part to *Herschel*’s larger primary mirror and, at most wavelengths, higher spectral resolution. *Spitzer*-IRS (Houck et al. 2004) observations reveal that some FUors resemble embedded protostars, in that their excess emission at wavelengths greater than $\sim 30 \mu\text{m}$ cannot be accounted for by circumstellar disk models alone, while others lack a strong mid-IR excess above that of a mildly flared disk (e.g. Green et al. 2006; Quanz et al. 2007c; Zhu et al. 2008). Thus a central question is whether FUors are all surrounded by envelopes – and in a corresponding earlier stage of development than a T Tauri star – or are drawn from both embedded and disk sources. *Herschel* is particularly well-suited to distinguish between disk and envelope structure, as the SEDs of envelopes peak in the far-IR.

In Green et al. (2011), we presented *Herschel* observations of the region surrounding the newly outbursting FUor HBC 722. Here we present the first results of the “FOOSH” (FU Orionis Objects Surveyed with Herschel) Open Time program (PI: J. Green): improved PACS (Photodetector Array Camera and Spectrometer, 50–210 μm ; Poglitsch et al. 2010) and SPIRE (Spectral and Photometric Imaging REceiver, 194–670 μm ; Griffin et al. 2010) spectroscopy of HBC 722, and newly observed PACS and SPIRE spectroscopy, and imaging (70 – 500 μm), of five older FUors: V1057 Cyg, V1331 Cyg, V1515 Cyg, V1735 Cyg, and FU Ori. Additionally we include HIFI (Heterodyne Instrument for the Far Infrared; de Graauw et al. 2010) observations of CO $J = 5 \rightarrow 4$ and $J = 14 \rightarrow 13$. We re-reduce archival *Spitzer* observations of V1331 Cyg, originally presented in part by Carr & Najita (2011). Complementary to the HIFI data, we present velocity-resolved ground-based spectra of selected submillimeter lines. We analyze the spatial structure of the line and continuum emission, and consider differences amongst the sample. Lastly we compare FUors to Class 0/I embedded sources from the DIGIT Key Program (Green et al. 2013) and other protostar surveys (HOPS, Manoj et al. 2013; WISH, Karska et al. 2013) and consider the impact of the *Herschel* data on our understanding of the evolutionary state of FUors.

2. OBSERVATIONS AND DATA REDUCTION

The FOOSH program consisted of 21 hrs of *Herschel* observing time: V1057 Cyg, V1331 Cyg, V1515 Cyg, V1735 Cyg, and FU Ori were observed as part of FOOSH. In addition, HBC 722 was observed with *Herschel* as a Target of Opportunity (PI: J. Green) and the initial results were presented in Green et al. (2011). The region around each source was observed with the SPIRE-FTS and PACS spectrometers and imaged with the SPIRE bolometer and PACS cameras. Additionally, each source was observed in selected bands with HIFI. Finally, all but FU Ori were observed in Sept. 2012 with the Caltech Submillimeter Observatory heterodyne re-

ceivers. A list of the observations included in this work is presented in Table 1.

2.1. The Sample

The basic characteristics of the FUors in our sample are listed in Table 2. Although all are classified as FUors, our sources were selected to span a variety of subtypes, with regard to local extinction, dust properties, observed eruption date, and SED shape (Green et al. 2006; Quanz et al. 2007c; Kóspál et al. 2011). Pre-*Herschel* observations of HBC 722 and FU Ori suggest that they fall toward the disklike end of the spectrum (based on their mid-IR colors and low extinction values), with low (~ 0.01 – $0.02 M_{\odot}$) upper limits on envelope mass (Sandell & Weintraub 2001; Dunham et al. 2012). V1735 Cyg is most similar to an embedded source with a highly extinguished central star and a rising mid-IR SED. V1331 Cyg, V1057 Cyg, and V1515 Cyg are intermediate between the FU Ori and V1735 Cyg cases; moderate extinction values, relatively blue SEDs but clear evidence for substantial envelope material. All but V1331 Cyg were identified as FUors by their optical flare, rather than by ancillary spectral characteristics. Typically, FUors are classified based upon their optical and mid-IR spectral features; in this work we also classify the sources by their full SED properties.

FU Ori. Located in Orion, FU Ori (600 pc) is the archetype of the FUor class, the earliest observed outburst and the most extreme, rising by 6 magnitudes in B over the period of 3 months in 1936. It has since faded by only ~ 1 mag, at a rate of 14 mmag yr^{-1} (Kenyon et al. 2000). FU Ori shows a relatively blue SED (Hartmann & Kenyon 1996; Green et al. 2006; Quanz et al. 2007c), with emission from pristine silicate dust in the Spitzer bands, and a relatively low extinction ($A_V = 1.8$ mag). Models of the FU Ori system (with a central stellar mass of $0.3 M_{\odot}$) suggest that the current accretion rate remains very high ($10^{-4} M_{\odot} \text{yr}^{-1}$) and it is the northern component in a close ($0''.5$) binary system (Reipurth & Aspin 2004; Wang et al. 2004). The non-outbursting star, FU Ori S, may be a more massive companion ($> 0.5 M_{\odot}$; Pueyo et al. 2012); the binary is unresolved by *Herschel*.

V1735 Cyg. Also known as Elias 1–12, V1735 Cyg (950 pc) is the only source in our sample that does not show silicate emission features in its IRS spectrum, but instead shows absorption due to CO and H₂O ice. It is the most extinguished source in the sample, with estimates of $A_V > 6$ mag. A second submillimeter source, V1735 Cyg SM1 is nearby (~ 20 – $24''$ distant; Sandell & Weintraub 2001). The mid-IR SED suggests substantial envelope emission in the system.

V1515 Cyg. Gradually rising during the 1940s and 50s, V1515 Cyg (1000 pc) shows a flat mid-IR SED with weak silicate emission features, and $A_V = 3.0$ mag. The region around V1515 Cyg shows arclike reflection nebula structure (Kóspál 2011).

V1057 Cyg. Erupting around 1969, V1057 Cyg (600 pc) is one of two sources in the FOOSH sample with a pre-outburst optical spectrum consistent with a T Tauri star (Herbig 1977). With $A_V = 3.5$ mag and a flat-spectrum SED, weak silicate emission, and a ring-like reflection nebula, V1057 Cyg is quite similar in character to V1515 Cyg in its optical/IR spectrum.

V1331 Cyg. Classified as “pre-outburst” or between outbursts (Welin 1977), V1331 Cyg (550 pc) was identified based on its spectral similarity to the pre-outburst optical spectrum of V1057 Cyg; no optical flare was observed and thus it is possible that V1331 Cyg is not an FUor (Sandell & Weintraub 2001; Quanz et al. 2007a; Kóspál et al. 2011). V1331 Cyg is the least luminous of the Cygnus sources in our sample. Hot water vapor was observed at K-band and was attributed to the inner disk (Najita et al. 2009). We present archival IRS spectra of this source, which closely resemble the flat spectrum SEDs of V1515 Cyg and V1057 Cyg, with weak, pristine silicate emission features at 10 and 20 μm . Even if V1331 Cyg is not an FUor, it represents an important historical comparison; the accretion rate would be very high for a non-outbursting T Tauri star, at $\sim 10^{-6} M_{\odot} \text{yr}^{-1}$ (Najita et al. 2009).

HBC 722. Also known as LkH α 188-G4, PTF10qpf, and V2493 Cyg, HBC 722 (520 pc) is a newly erupted FUor (Semkov & Peneva 2010; Miller et al. 2011) located in the “Gulf of Mexico” in the southern region of the North American/Pelican Nebula, among a closely-packed set of accreting young stars. Pre-outburst observations (Cohen & Kuhl 1979) indicate that HBC 722 was a $\sim 0.5 L_{\odot}$ K7-M0 spectral type T Tauri star prior to 2009. Between 2009 and 2010, HBC 722 rose by ~ 1 mag in B , and then suddenly flared between July and Sept. 2010, rising by 3.5 mag and in luminosity by a factor of ~ 20 to $12 L_{\odot}$ over the course of 3 months (Kóspál et al. 2011). After decreasing to $5.4 L_{\odot}$ between Sept. 2010 and May 2011, HBC 722 began to rise in brightness once more, re-achieving peak brightness by June 2012 and exceeding it in optical bands by 2013. High cadence photometric observations reveal a multitude of periods and “flickering” that may be associated with a still-active inner accretion disk (Green et al. 2013). Resolved SMA observations detect little remnant envelope material close to the star (Dunham et al. 2012).

2.2. Imaging

SPIRE imaging was gathered in single-cycle integration times in the small map mode, a single (349 s) observation per source. SPIRE observed simultaneously at 250, 350 and 500 μm . The on-orbit beam sizes are $18''.1$, $25''.2$, and $36''.6$, respectively. The *Herschel* data were reduced using HIPE (*Herschel* Interactive Processing Environment; Ott 2010) pipeline.

PACS imaging was acquired using two pairs of single-cycle scan maps. For each source a “blue” (70 and 160 μm) and “green” (100 and 160 μm) simultaneous observation was taken, for two different orientation angles (70° and 110°), a total of four (279 s) observations per source. The final 160 μm image was produced by coadding all four scans; the 70 and 100 μm images were averaged from their two respective scans each. The on-orbit beam sizes for 70, 100, and 160 μm are $5''.5$, $6''.7$, and $10''.7$, respectively.

2.3. Spectroscopy

2.3.1. SPIRE-FTS

Our SPIRE-FTS data were taken in a single pointing with sparse image sampling, high spectral resolution, in 1 hr of integration time. The spectrum is di-

vided into two orders covering the spectral ranges 194 – 325 μm (“SSW”; Spectrograph Short Wavelengths) and 320 – 690 μm (“SLW”; Spectrograph Long Wavelengths), with a resolution of $\lambda/\Delta\lambda \sim 300\text{--}800$, increasing at shorter wavelengths. Each order was reduced separately within HIPEv9.0 using the standard pipeline for extended sources, including apodization. SPIRE used an onboard calibration source for flux calibration. Although SPIRE observed simultaneously along multiple spatial pixels, at this stage the off-source pixels are not yet flux-calibrated. Thus we used off-source pixels only to determine if a source is extended; all presented spectra were taken using data from the central pixel only. No background subtraction was performed, and thus the spectra should be considered an upper limit to the true flux density; we note cases where extended emission was seen in the following sections. The measured uncertainties in the continuum ranged from 0.05-0.5 Jy (an improvement over that reported in Green et al. 2011 due to pipeline improvements in the interim), or $\sim 0.5\text{--}5 \times 10^{-18} \text{ W m}^{-2}$ in linefluxes, increasing with observed frequency and several times higher in SSW compared to SLW.

2.3.2. PACS

PACS was a 5×5 array of $9''.4 \times 9''.4$ spatial pixels (also referred to as “spaxels”) covering the spectral range from 51 – 210 μm with a resolution $\lambda/\Delta\lambda \sim 1000\text{--}3000$, divided into four orders, covering $\sim 50\text{--}75$, $70\text{--}105$, $100\text{--}145$, and $140\text{--}210 \mu\text{m}$, referred to as “B2A”, “B2B”, “short R1”, and “long R1”. Each order was first reduced using a modified pipeline optimized for extended sources; a detailed description is provided in Green et al. (2013). For the FOOSH sample we utilized the full range of PACS (50-210 μm) in two linked pointed chop/nod rangescans: a blue scan covering 50-75 and 100-150 μm (SED B2A + short R1); and a red scan covering 70-105 and 140-210 μm (SED B2B + long R1). We used 6 and 4 range repetitions respectively, for integration times of 3530 and 4620 seconds (a total of ~ 8000 seconds per target and off-positions combined, for the entire 50-210 μm scan; the on-source integration time is ~ 3000 seconds). The telescope sky background was subtracted using two nod positions $6'$ from the source. The absolute flux uncertainty was estimated at 30%, typically used for multi-wavelength comparisons in extended sources (e.g. Herczeg et al. 2012). Further detail on systematic uncertainties can be found in the PACS Observer Manual.

2.3.3. HIFI

Herschel-HIFI provided high spectral resolution, down to $\sim 0.03 \text{ km s}^{-1}$. HIFI was used in single point mode, in band 6b (1578.2–1697.8 GHz) tuned to CO $J = 14 \rightarrow 13$ (1611.8 GHz), and in band 1b (562.6–628.4 GHz) tuned to CO $J = 5 \rightarrow 4$ (576.3 GHz). The data were again reduced using HIPE. The beamsize of HIFI is $43''.1$ in band 1b, and $15''$ in band 6b (Roelfsema et al. 2012). The total integration time was 622 and 677 seconds per observation in band 1b and 6b, respectively, and the uncertainty in frequency was less than 100 kHz, or 0.05 km s^{-1} . The HIFI data were reduced in a manner similar to that described in Kristensen et al. (2012). The spectra were reduced in HIPE and exported to the “CLASS” analysis package for further reduction and analysis. The

reduction consisted of subtracting linear baselines and averaging data from the H- and V-polarizations, the latter only after visual inspection of the data from the two polarizations. A main beam efficiency of 0.75 was adopted to convert from antenna temperature (T_A^*) to main beam temperature (T_{MB}), and integration over the line yielded a value for the integrated intensity ($\int T_{MB} dv$) (Roelfsema et al. 2012).

2.3.4. CSO

Observations of V1735 Cyg, V1057 Cyg, V1331 Cyg, V1515 Cyg, and HBC 722 were obtained with the 230 GHz heterodyne receiver installed on the Caltech Submillimeter Observatory (CSO) between 8 and 22 September 2012. We observed each source in $\text{HCO}^+ J = 3 \rightarrow 2$ (267.557 GHz) achieving a minimum S/N of 5 on the detected line, with τ_{225} (the atmospheric opacity at 225 GHz) varying between 0.08 and 0.13. Additionally we observed HBC 722 in $\text{CO } J = 4 \rightarrow 3$ (461.041 GHz) using the 460 GHz receiver on 17 September 2012, with $\tau_{225} = 0.070$. The beamsize is $\sim 28''$ for HCO^+ and $\sim 16.5''$ for $\text{CO } J = 4 \rightarrow 3$. The data were then exported to the CLASS package for further reduction and analysis, similar to HIFI.

2.3.5. IRS

Spitzer-IRS (InfraRed Spectrograph) was a long-slit spectrograph with $2\text{--}5''$ pixels, covering the spectral range $5.3\text{--}38\ \mu\text{m}$, with a resolution $\lambda/\Delta\lambda \sim 60\text{--}600$. V1331 Cyg (AORs 14547712 and 14549760) was processed using a custom point-source pipeline (Furlan et al. 2006), cleaned using an aggressive bad pixel mask generated from a “grand rogue mask”, in which all pixels exhibiting greater than 4σ deviations in dark frames (attributed to cosmic rays, permanent hot pixels and other effects) taken during more than one campaign were flagged and interpolated using a “nearest neighbor” technique.

2.4. Absolute Flux Calibration and SEDs

In order to produce a single SED for each source, we needed to calibrate the continuum flux density across *Spitzer*-IRS, *Herschel*-PACS, and *Herschel*-SPIRE, which range in spatial resolution from $2\text{--}40''$. At the shortest wavelengths of IRS, the spatial resolution was 1100 AU at a distance of 550 pc; the material emitting in the *Spitzer* bands remains unresolved. To determine the continuum level, we took the final output of each instrument and scaled, first between modules, and then between instruments. At shorter wavelengths with improved spatial resolution these sources are dominated by disk/envelope warm dust, within the IRS beam as listed above; the FWHM of the mid-IR continuum measured from IRS was typical of YSOs, with no more than a 20% flux mismatch between modules for the four sources observed in Green et al. (2006). Thus in each case we assumed that the source was essentially unresolved, and increased spatial resolution at shorter wavelengths provided a better absolute flux calibration. Where possible we then correlated this with photometry, extrapolating between missing wavelength regions. All of the scaling is described below.

The IRS data for V1515 Cyg, V1057 Cyg, and FU Ori were first presented in Green et al. (2006), and the

IRS spectra for V1735 Cyg was presented in Quanz et al. (2007c); for this work, we reduced V1331 Cyg using the same technique (the SH reduction appeared in Carr & Najita 2011). “SL” (Short-Low module; $5.3\text{--}14\ \mu\text{m}$) data were multiplied by 1.03 and “LH” (Long-High module; $20\text{--}38\ \mu\text{m}$) multiplied by 1.05 to match “SH” (Short-High module; $10\text{--}20\ \mu\text{m}$) to produce the final spectrum. Additionally we note that in V1735 Cyg, the SL spectrum was scaled by 1.35 to match SH (SL was not included in the Quanz et al. 2007c data). HBC 722 was not observed by IRS (pre- or post-outburst). As described in Green et al. (2006), the spectra were dereddened using an analytical extinction law (Mathis 1990), using an R_V of 3.1, in the cases of FU Ori, V1515 Cyg, and V1057 Cyg.

For all six sources, the PACS data were scaled, module-by-module, using a polynomial factor (as a function of wavelength), a technique that maximized continuum S/N while determining the proper absolute flux calibration. The full technique, as applied to the DIGIT dataset, is described in detail in Green et al. (2013). We modified this procedure in one way: in all but HBC 722 we included all flux in the 3×3 spaxels as part of the FUor environment, for both continuum and lines, rather than using the full array for continuum and the smaller 3×3 for lines.

For HBC 722, the SPIRE flux density was much greater than the PACS continuum. This is likely because, as noted in Green et al. (2011), the continuum flux density near HBC 722 was contaminated by an embedded source, 2MASS 20581767+4353310, $17''$ to the southeast. SPIRE could not resolve these sources, but *Spitzer* and PACS photometry could do so. We produced a PACS spectrum of the entire 3×3 central region summed together, which included the embedded source and was similar in spatial extent and measured flux density to that of the SPIRE beam. Thus the composite “HBC 722” spectrum is likely dominated by the embedded source rather than HBC 722 itself. In addition, a second epoch of observations of the same field but with a different orientation angle were taken in this program, and the absolute flux matched to within 10% for both extractions. For the HBC 722 composite spectrum, we multiplied B2B by 1.5, and both R1 modules by 0.67. In all other cases the standard DIGIT procedure (scaling to the 3×3 spaxels) produced a smooth PACS continuum. No scaling relative to IRS was performed, but the final PACS spectrum was a reasonable match to the extrapolated IRS continuum at $60\ \mu\text{m}$.

V1735 Cyg also suffered from confusion with V1735 Cyg SM1 ($24''$ to the northeast), at $\lambda > 100\ \mu\text{m}$. The spectroscopic flux density using our standard central 3×3 spaxel aperture exceeded the PACS photometry by 50% at 70 and $100\ \mu\text{m}$, and by 30% at $160\ \mu\text{m}$. We infer that an annulus of $\sim 17''$ or less is required to avoid contamination. However, the PACS spectroscopic data are not sufficiently sampled to extract the spectrum from a $17''$ aperture at all wavelengths, and thus we use the larger aperture and assume contamination at these levels.

The SPIRE data were extracted using the “extended source” calibration pipeline, as this produced a smoother continuum between modules, better S/N, and fewer spectral artifacts than the “point source” pipeline. First we multiplied SLW to match the SSW flux density in the

overlap region of the two modules (300-320 μm), and then matched the flux at 210-220 μm to an extrapolation of the final PACS spectrum. HBC 722 and V1735 Cyg required no additional scaling, either between modules, or between instruments. For both FU Ori and V1515 Cyg, we multiplied SSW by 0.9 to match PACS. For both V1331 Cyg and V1057 Cyg, we multiplied SLW by 1.55 to match SSW, with no additional scaling to match PACS.

2.5. Wavelength Calibration

For the DIGIT sample, we noted that the uncertainties in the PACS wavelength centroid were at least 30-50 km s^{-1} . With shorter exposure times in FOOSH, we found slightly more erratic centering in our line positions, up to 100 km s^{-1} . Small mispointings can masquerade as velocity shifts across the PACS IFU (see the PACS Observer Manual for further information); thus we cannot determine the gas velocities to better than 100 km s^{-1} . The SPIRE spectra have even greater uncertainty due to lower spectral sampling and bigger beamsize. Given the uncertainties, we did not observe any convincing velocity shifts in our data.

3. RESULTS

3.1. Imaging

Figure 1 shows three-color (70, 100, and 160 μm) PACS images, and Figure 2 shows three-color (250, 350, and 500 μm) SPIRE images for the five classical FUors in our sample. We consider each source below. Table 3 lists the detected flux in 10-15'' (PACS PSF corrected) and 30'' (SPIRE) apertures centered on the source coordinates. In summary, most of the FUors appear pointlike at 70 μm , but gradually the extended filaments begin to dominate at longer wavelengths.

The HBC 722 images were first presented in Green et al. (2011). All are surrounded by substantial emission from multiple sources in the far-IR which are detected within the SPIRE beam, especially at longer wavelengths. The dominant continuum source in the HBC 722 field is actually 2MASS 20581767+4353310, an embedded protostar (Green et al. 2011). They note a separate 0.41 Jy source detected at PACS 70 μm at the position of HBC 722, below the sensitivity of the spectroscopic observations.

V1515 Cyg is detected in all bands, but does not show strong continuum sources at the stellar position at SPIRE wavelengths, relative to the surrounding nebosity. At 70 μm the FUor is easily distinguished from the background, but at 500 μm the arclike emission extending 2' to the north becomes dominant.

Similarly, V1331 Cyg, V1057 Cyg, and FU Ori show clear detections with PACS but additional arclike structures generally correlating with previously analyzed mm data (e.g. Kóspál et al. 2011). In each case, the FUor appears much bluer than the filamentary material.

In the case of V1735 Cyg, the optical outbursting source is the fainter western extension of the submillimeter peak; the peak emission comes from the embedded source V1735 Cyg SM1, consistent with observations by Spitzer-MIPS (Harvey et al. 2008). The beamsize of MIPS at 70 μm is comparable to that of SPIRE at 250 μm , allowing marginal resolution of the two sources. This

is apparent from the PACS images where the sources are spatially resolved. At longer wavelengths, they are inextricably blended.

3.2. Continuum SEDs

Figure 3 shows continuum SEDs for all five FUors, including UBVR, JHK (2MASS, UKIDSS), Spitzer-IRS spectra (from Green et al. 2006; Quanz et al. 2007c, and this work) and PACS/SPIRE photometry (this work). For comparison, we also show the incomplete SED of HBC 722 from Green et al. (2011); the PACS/SPIRE spectra for this source are clearly contaminated, as evidenced by the mismatch with the 70 μm PACS photometric datapoint, and the non-detection at 1.3 mm with the Submillimeter Array (Dunham et al. 2012).

Additionally, the PACS photometry falls below the PACS spectrum for V1735 Cyg and V1515 Cyg. The former can be explained by contamination in the spectra from V1735 Cyg SM1, which is cleanly separated by the photometry but not by the spectroscopy. A similar situation is likely in V1515 Cyg, where the spectroscopic data includes additional surrounding nebosity.

The most notable slope discontinuity between PACS and SPIRE is in the FU Ori SED; the emission at SPIRE wavelengths is faint by comparison to the other sources. In V1331 Cyg, the longest wavelength SPIRE photometry, at 500 μm , seems contaminated by extended emission. In the other, brighter sources, the PACS and SPIRE photometry and spectroscopy are very well-aligned.

As FUors are variable and generally fading over time, we used photometry as close to contemporaneous as possible. We then calculated L_{bol} and T_{bol} using the technique from Dunham et al. (2010). In summary, T_{bol} is the temperature of a blackbody with the same flux-weighted mean frequency as the measured spectrum, and L_{bol} is the total integrated luminosity, using trapezoidal interpolation between datapoints.

3.2.1. Classification

The original definition of Class I and II sources for young stellar objects is derived from the ratio of the 2 μm / 25 μm flux density (Andre et al. 1993; Greene et al. 1994), the spectral index α . For HBC 722, there is a pre-outburst measurement of α at -0.77, which would fall under the definition of Class II (Miller et al. 2011). Quanz et al. (2007c) used mid-IR spectral indices and found the FUors to be indistinguishable from Class II objects in Taurus (Furlan et al. 2006). We compare the classifications of the non-confused FUors in our sample by three different classification methods, explained below, in Table 4.

In general, T_{bol} can be used as a proxy for evolutionary state (Chen et al. 1995; Robitaille et al. 2006): in this framework a Class I source has T_{bol} between 70 and 650 K; a Class II source has $T_{\text{bol}} > 650$ K. Flat-spectrum sources can be characterized as having $350 \text{ K} < T_{\text{bol}} < 950 \text{ K}$ (Evans et al. 2009; Fischer et al. 2013). T_{bol} is quite sensitive to the near-IR flux and is thus highly uncertain in variable sources such as FUors. With that caveat in mind, we find that FU Ori, V1515 Cyg, and V1057 Cyg exhibit Class II SEDs, while V1735 Cyg and V1331 Cyg exhibit typical flat spectrum SEDs, near the Class I/II boundary. This sequence

agrees with the sequence from SED color, extinction, and strength of silicate emission found in Green et al. (2006) and Quanz et al. (2007c). The lone surprise is V1331 Cyg which might be expected to be the second-most blue/evolved source, given its silicate emission feature and flat/slightly declining mid-IR SED. In summary, the T_{bol} values suggest that FUors are more evolved than Class I sources. Unfortunately, the value of T_{bol} as an evolutionary proxy is debatable. In the case of an FUor flare, models suggest that the far-IR luminosity will increase within a few months of outburst (Johnstone et al. 2013), although Ábrahám et al. (2004) find, in contrast, that changes to the SED occur primarily in the optical/IR rather than the submillimeter on the 1-10 yr timescale. If the far-IR lagged significantly behind, this would increase T_{bol} without any change in the submillimeter SED or derived envelope mass. However, this scenario would require the dust to be 1-10 ly (0.3 - 3.0 pc) distant, and hence unrelated to the FUor.

Thus we also consider a criteria related to the physical stage of the object, using the Herschel spectra. We calculate the envelope mass using the “OH5” opacity law (Ossenkopf & Henning 1994) as described in e.g., Dunham et al. (2011). First we determine an approximate dust temperature from the shape of the PACS and SPIRE continuum from 100-600 μm . We compare dust at this temperature, assuming an emissivity $\alpha \lambda^{-1.8}$, normalized to the 350 μm continuum, to confirm that the model slope is similar to the observed spectrum. Then we calculate an envelope mass from the SPIRE photometry. The results for the four isolated sources in our sample are shown in Table 5. Using this criteria, Crapsi et al. (2008) consider 0.1 M_{\odot} as the envelope mass cutoff between Stage I and II sources. Their boundary value was derived from Robitaille et al. (2006) models, converting their boundary accretion rate value of $10^{-6} M_{\odot} \text{yr}^{-1}$ to 0.07 M_{\odot} , with the additional consideration that disks greater than 0.1 M_{\odot} were gravitationally unstable. However, there are two dominant sources of uncertainty in this calculation. First, the dust temperature is not perfectly constrained, and varies between 38 and 54 K, causing an uncertainty in the derived envelope mass of $\sim 35\%$. Second, the estimated mass increases significantly if derived using longer wavelengths, due to the inclusion of increasingly larger areas of colder dust from the more distant parts of the envelope, suggesting that values calculated at 350 μm generally underestimate the envelope mass.

The results for each classification method are summarized in Table 4. The α and T_{bol} classifications agree very well, while the envelope mass method indicates an earlier evolutionary stage than the other methods in the cases of V1057 Cyg and V1515 Cyg. If we crudely average the classifications of the three different methods (assigning values of 1, 1.5, and 2 to Class/Stage I, Flat Spectrum, and Class/Stage II, respectively), we find the following sequence in increasing order of Class/Stage: V1057 Cyg (1.5), V1515 Cyg (1.67), V1331 Cyg (1.67), V1735 Cyg (1.75), and HBC 722 (2.0). This ordering is broadly in agreement with the mid-IR dust feature sequence from Quanz et al. (2007c). Therefore, we find that the umbrella group of “FUors” includes both Stage I and II objects.

3.2.2. Bolometric luminosity

We also update L_{bol} for our sample. First it is important to consider the observational epochs of the data in our SEDs: the optical/near-IR data has been updated to match the 2005-2006 *Spitzer* spectra; the *Herschel* data are from 2011. Thus the computed L_{bol} values are upper limits to the current L_{bol} , if all are declining. Ábrahám et al. (2004) found that the flux for $\lambda > 20 \mu\text{m}$ held constant in time; thus the SEDs including the *Herschel* data may closely represent the SED from the 2005-6 timeframe. However, Johnstone et al. (2013) find that while the wavelength-dependent variability during outburst is led on day timescales at mid-IR and shorter wavelengths, at wavelengths close to 100 μm the effect on the SED occurs over months, and the $\lambda \lesssim 200 \mu\text{m}$ *Herschel* spectra could be sensitive to changes.

The addition of the Herschel data affects the derived L_{bol} and T_{bol} systematically, but this is overshadowed by a larger effect attributable to beamsize. First, we compare the SEDs with and without the *Herschel* data: including PACS/SPIRE data along with ancillary photometry results in larger L_{bol} , and smaller T_{bol} , by 10-20% for each source, compared to the values derived only from the ancillary photometry. The reason for this systematic increase in L_{bol} is that the broad extrapolation from *Spitzer* (35 μm) data to the 850 μm or 1.1 mm data, in the absence of the *Herschel* data, slightly understates the total flux.

However, the beamsize of far-IR (IRAS) data used in older calculations of L_{bol} likely caused significant overestimation of the bolometric luminosity. Comparing our measurements to earlier values of L_{bol} , V1057 Cyg peaked at 800 L_{\odot} at peak and dropped to 250 L_{\odot} (Kenyon & Hartmann 1991), but we calculate an L_{bol} of 119 L_{\odot} . The large beamsize of the far-IR data used in the Kenyon & Hartmann (1991) estimate causes a significant overestimate: we calculated L_{bol} for V1057 Cyg using only the optical/near-IR and IRAS fluxes from Kenyon & Hartmann (1991) and found an L_{bol} of 215 L_{\odot} , comparable to their derived value of 250 L_{\odot} . This is much greater than the 119 L_{\odot} we calculate using the *Herschel* fluxes; thus we conclude that the discrepancy is due to the greater fluxes at 25 and 60 μm from IRAS compared with PACS. In the case of FU Ori, we observed a similar but smaller effect; although the IRAS fluxes are a factor of 2-3 higher than the PACS fluxes, L_{bol} only changes from 274 to 210 L_{\odot} with the updated data. FU Ori is a much bluer object than V1057 Cyg, and thus its L_{bol} is less sensitive to changes in submillimeter flux. Whether this is a beamsize effect or an actual decrease in luminosity is not clear, in either source.

With these caveats in mind, we find that FU Ori is currently the brightest source at 210 L_{\odot} . V1057 Cyg, V1515 Cyg, and V1735 Cyg all show $L_{\text{bol}} \sim 100\text{-}120 L_{\odot}$, while V1331 Cyg is much fainter, at 16.7 L_{\odot} . In general, we find lower L_{bol} than previously reported values in the literature, but we cannot separate the large uncertainty in older measurements of far-IR/submm flux from an actual decrease in the far-IR/submm SED. Resolving this ambiguity is beyond the scope of this work but can be addressed with multi-epoch submillimeter continuum measurements.

3.3. Non-Detection of Dust and Ice Features

We do not observe any evidence for cold crystalline Mg-rich or Fe-rich dust at $69\ \mu\text{m}$, which has been seen in some T Tauri and Herbig Ae/Be star disks (Sturm et al. 2010; Sturm et al. 2013). However, our sensitivity is substantially lower than in sources with detections of this feature, due to shorter integration times, and greater distances to the source (lower S/N). Furthermore, the sources in Sturm et al. (2013) typically show crystalline silicate dust in the IRS bands, whereas all of the FUor IRS spectra are pristine (Quanz et al. 2007c). At this time, calibration uncertainties in spectral shape prevent a search for $60\text{--}70\ \mu\text{m}$ H_2O crystalline ice features.

3.4. Line Detections

For the DIGIT embedded objects we used separate extractions to determine line spatial extent as distinct from continuum (Green et al. 2013). This was unnecessary for the FOOSH sources, which show no evidence of extended line emission, except in cases of contamination. Differences in the continuum and line extent resolved with *Herschel* in FUors is attributed to large scale structure, multiplicity, or outflows and is discussed in the following section. For this work we follow a slightly simpler procedure than that used in Green et al. (2013), and assume the line and continuum emitting regions are the same, and apply the same scaling factors.

The list of detected lines in PACS and SPIRE for all six source regions is reported in Table 6. In summary, the PACS observations show few lines overall, while the SPIRE spectra are richer. The only lines detected in all six sources are [O I] $63\ \mu\text{m}$ and [C I] 370 and $610\ \mu\text{m}$. Other observed fine structure lines include [O I] $145\ \mu\text{m}$, [N II] 205 and $122\ \mu\text{m}$, and [C II] $158\ \mu\text{m}$. We detect CO $J = 5 \rightarrow 4$ and $J = 4 \rightarrow 3$ in all but FU Ori (although CO $J = 5 \rightarrow 4$ is detected only in the HIFI data, below the detection threshold of SPIRE due to the narrow feature width). CO $J = 7 \rightarrow 6$ and higher are seen only in V1057 Cyg, V1735 Cyg, and HBC 722; in V1057 Cyg we detect CO up to $J = 23 \rightarrow 22$. We detect ^{13}CO $J = 8 \rightarrow 7$ to $J = 5 \rightarrow 4$ in V1735 Cyg and HBC 722. We detect H_2O $174.6\ \mu\text{m}$, and tentatively detect OH $84.41\ \mu\text{m}$, in one source (V1057 Cyg). OH $119\ \mu\text{m}$, the lowest state of the OH $3/2 \rightarrow 3/2$ ladder, is the only line in absorption (seen in V1735 Cyg).

None of the sources was detected in CO $J = 14 \rightarrow 13$ with HIFI, but all were detected in CO $J = 5 \rightarrow 4$. Figure 4 (top) shows the CO $J = 5 \rightarrow 4$ HIFI (black) compared with the HCO^+ $J = 3 \rightarrow 2$ ($267.558\ \text{GHz}$; blue) line observed with the CSO. FU Ori was not observed in HCO^+ in our observing run; all of the others – V1515 Cyg, V1735 Cyg, V1057 Cyg, V1331 Cyg, and HBC 722 – were observed to emit in the HCO^+ $J = 3 \rightarrow 2$. In the case of HBC 722, we also observed CO $J = 4 \rightarrow 3$ and $J = 2 \rightarrow 1$. The line profiles are compared in Figure 4 (bottom), placed on the same vertical scale for comparison. (Note that the CO integrated intensity is scaled by the factors listed in each subfigure.)

As an exemplar, in Figure 5 we present the $50\text{--}670\ \mu\text{m}$ continuum-subtracted spectrum of V1057 Cyg (top), rebinned to low resolution for clarity, vs. V1735 Cyg (bottom), which is partly contaminated by V1735 Cyg SM1. In this form, a few details become apparent. The

peak CO lineflux appears at much longer wavelengths in V1735 Cyg, compared to V1057 Cyg, or to a typical embedded source (e.g. Fig. 15, Green et al. 2013). The faint ^{13}CO ladder appears from $J = 5 \rightarrow 4$ to $J = 8 \rightarrow 7$ only in the case of V1735 Cyg. The only detection of H_2O (at $174.6\ \mu\text{m}$) is in V1057 Cyg (the stronger nearby line is CO $J = 15 \rightarrow 14$ at $173.6\ \mu\text{m}$). V1735 Cyg shows a relatively weak [O I] feature in comparison to the CO.

The remaining FUor spectra (including the contaminated HBC 722 from Green et al. 2011) are shown in Figures 6 and 7.

3.5. Spatial Distribution of Continuum and Lines

Although we detect substantial line emission from three sources (V1057 Cyg, V1735 Cyg, and HBC 722), we must consider whether the lines are local to each FUor, using the spatial information provided by the PACS array. It is more difficult to determine this with SPIRE due to the larger beam, sparse spatial coverage, and lack of off-position data.

3.5.1. PACS

Figures 8 and 9 show 10% contours for all detected lines (grayscale) and local continuum (red), with PACS. The local continuum is selected from a line-free set of channels surrounding each line. In HBC 722, we see a clear demarcation between the FUor (central spaxel) and the embedded source (two spaxels to the SE). The submm continuum is strongly dominated at all wavelengths by the embedded source. For V1735 Cyg, V1735 Cyg SM1 appears as a continuum source cleanly detected two spaxels to the east, at the edge of the array. In the other four sources (V1515 Cyg, V1057 Cyg, V1331 Cyg, and FU Ori) the central spaxel dominates the continuum emission, and the pointing accuracy can be estimated as within 0.2 spaxels from the centroid at each wavelength. The continuum footprint grows with wavelength, consistent with the point-spread function (Green et al. 2013); thus there is no evidence for extended PACS continuum emission in these four sources.

The line emission differs from the continuum profiles in a few cases. In the HBC 722 composite, the [O I] $63.18\ \mu\text{m}$ emission is co-spatial with the FUor, while the CO $J = 16 \rightarrow 15$ emission extends along both sources in a SE-NW band roughly correlated with the CO $J = 2 \rightarrow 1$ map (Green et al. 2011). The marginally detected H_2O emission at $179.53\ \mu\text{m}$ is associated with the embedded object to the SE. We attribute the [O I] to the FUor, although we note that the [O I] could be excited by shocks driven by other millimeter sources identified in SMA observations (Fig. 1, MMS1 and MMS2; Dunham et al. 2012). In V1735 Cyg, the [O I] emission is equally prominent around both the FUor and the SM1 source, but CO $J = 16 \rightarrow 15$ is compact and associated only with SM1. In V1057 Cyg, V1515 Cyg, V1331 Cyg, and FU Ori, the [O I] and CO emission is compact and aligned with the FUor, when detected (although there is a hint of extended emission in FU Ori [O I]). There is very little line emission detected in V1515 Cyg at all. [N II] $122\ \mu\text{m}$ is detected in off-center spaxels in V1515 Cyg, suggesting that the emission is diffuse. [O III] ($88.35\ \mu\text{m}$) and [N III] ($57.0\ \mu\text{m}$) appear only in absorption, indicating emission lines in the off-position. In HBC 722, the

rotation of the off-position between our two data epochs dominated change in flux of the [C II], [O III], and [N III] absorption features, clear evidence that they derive from the off-position and are not associated with the sources in the PACS field.

3.5.2. SPIRE

We were able to utilize software within HIPE to determine the contributions from individual pixels of the SPIRE-FTS detector array. The spacing between pixels in sparse mode is $\sim 33''$ and $51''$ for the short (SSW) and long wavelength (SLW) modules, with 37 and 19 hexagonal detectors respectively. A simple analysis of the spatial distribution of bright lines reveals some interesting trends.

The [N II] 205.4 μm line appears to be diffuse and present in nearly every pixel surrounding the HBC 722, V1515 Cyg, and V1331 Cyg fields. There is some hint of a central peak around V1057 Cyg. [N II] is absent in the FU Ori and V1735 Cyg fields. As SPIRE uses an onboard calibrator, and no off-position, we attribute the [N II] emission to diffuse background emission in Cygnus. This is consistent with fine structure emission ([S III], [Si II]) seen in IRS spectra of these sources prior to off-position subtraction (Green et al. 2006).

The mid- J CO emission is much more compact, with the low-lying ($J_{\text{up}} < 9$) states of CO originating in the central pixel for V1515 Cyg, V1057 Cyg, V1331 Cyg, and FU Ori, the sources which do not suffer from confusion.

3.5.3. Summary of Results

Both HBC 722 and V1735 Cyg are contaminated by nearby protostars at longer wavelengths. In the case of HBC 722, the [O I] line flux appears compact and local to the FUor; the CO is either extended in the foreground, or a chain of emission knots including HBC 722; the rest of the line emission in Table 6 and all continuum emission originates from the spaxels containing 2MASS 20581767. In V1735 Cyg, the contributions of the two sources are comparable at PACS wavelengths. The [O I] and PACS continuum are separable from V1735 Cyg SM1. However, at SPIRE wavelengths they are blended. Thus we *cannot* attribute the CO or ^{13}CO emission to V1735 Cyg alone.

In the other four FUors, we attribute all [O I], CO, H_2O , and continuum to the FUor itself. V1057 Cyg is an outlier, as the only non-confused FUor with detectable high- J CO, H_2O and OH emission; the other three (V1515 Cyg, V1331 Cyg, and FU Ori) show only strong [O I] and weak detections of CO up to $J = 6 \rightarrow 5$ with no H_2O or OH observed.

The final calibrated line fluxes appear in Table 6.

4. ANALYSIS

4.1. [O I]

HBC 722 and V1735 Cyg are the sources with the strongest [O I] 63 μm lines. They are also the only two sources with [O I] 145 μm detections. The 63/145 μm [O I] flux ratios are 21 and 11 respectively. The ratios fall slightly outside the range of the DIGIT embedded sample (14–20), for HBC 722 and V1735 Cyg. For the other sources, the lower limit to the 63/145 μm ratio is 15 (V1057 Cyg), 7.6 (FU Ori), and 1.5 (V1331 Cyg and V1515 Cyg).

Figure 10 shows the distribution of [O I] 63 μm line luminosity for the DIGIT embedded and FOOSH samples. A K-S test shows that the FUors are significantly different from the protostars as a group. The average (median) [O I] luminosity is 4.3 (3.1) $\times 10^{-3} L_{\odot}$ for the FUors, compared to 1.6 (0.5) $\times 10^{-3} L_{\odot}$ for the Class 0/I sources. If we compare to more evolved sources, the strongest [O I] emission in any Herbig Ae/Be star in the DIGIT sample is $1.4 \times 10^{-3} L_{\odot}$, and most Herbig/Class II sources fall below $10^{-4} L_{\odot}$ (Meeus et al., Fedele et al., subm.), for cases where the [O I] line is detected at all.

Figure 11 shows the correlation between [O I] line strength and L_{bol} across the DIGIT embedded and FOOSH samples. The large [O I] line strength in the FOOSH sample can be attributed to greater L_{bol} , albeit with considerable scatter.

We consider whether shocks contribute significantly to [O I] emission in our sample. The [O I] can be excited by either UV radiation from the central star or shocks (Hollenbach 1985). The earliest outburst in the FOOSH sample, FU Ori, occurred 75 years prior to the *Herschel* observations. At 300 km s^{-1} , the wind-lofted material would have propagated $\sim 4800 \text{ AU}$, slightly less than the beamsize of $\sim 6000 \text{ AU}$ for these sources. The emitting area for [O I] increases over time as the shock propagates outward, so a propagating shock model might predict higher [O I] line fluxes from the older outbursts, which we do not observe.

Moreover, a slower wind speed of 100 km s^{-1} is more consistent with the accretion rate in five of the six sources. The accretion luminosity is calculated (e.g. Hartmann 1998):

$$L_{\text{acc}} = \frac{GM\dot{M}}{2R_{\star}} \quad (1)$$

where R_{\star} is taken as $2R_{\odot}$. The [O I] line luminosity is related to the instantaneous accretion rate, $\dot{M} = 10^4 L([\text{O I}] 63 \mu\text{m}, \text{ in } L_{\odot})$, in units of $10^{-5} M_{\odot} \text{ yr}^{-1}$ (Hollenbach 1985). We compare these two mass loss indicators in Figure 12. The mass loss rates are consistent with a wind speed of 100 km s^{-1} , except in the case of HBC 722, in which the accretion luminosity *underestimates* the [O I] luminosity. The shock had only \sim three months to propagate since peak outburst ($\sim 5 \text{ AU}$ at a shock speed of 100 km s^{-1}), and thus the physical region of [O I] enhancement due to shocks should be small, yet we observe the opposite effect (albeit at weak significance).

In contrast, the [O I] line flux from UV radiation should track L_{bol} . Instead, the mass loss rate indicated by [O I] is constant as a function of accretion luminosity, between $\sim 10^{-5}$ and $10^{-6} M_{\odot} \text{ yr}^{-1}$ for all six sources. This suggests that either the [O I] emission is dominated by the current UV field, or that contributions from outflow-driven [O I] emission predate the current outburst. This is particularly interesting given that HBC 722 appears to be a relatively evolved burst source with little remnant envelope to fuel multiple burst cycles at this stage.

4.2. CO

In addition to [O I], CO is detected in a few transitions in FU Ori, V1515 Cyg, and V1331 Cyg, and in 20

transitions in V1057 Cyg. One useful tool in analyzing molecular emission is the rotational diagram; a detailed review can be found in Goldsmith & Langer (1999); a brief review in the context of *Herschel* spectroscopy is in Green et al. (2013). As with samples of protostars, we can fit components to distinct ranges of CO J_{up} : to reduce artifacts due to absolute flux calibration, we consider separately CO from SPIRE ($J_{\text{up}} \leq 13$), PACS R1 ($14 \leq J_{\text{up}} \leq 24$), and PACS B2A/B2B ($25 \leq J_{\text{up}}$; no detections in this sample). The results are presented in Figure 13.

Most Class 0/I (embedded and non-outbursting) sources observed with *Herschel*-PACS exhibit CO emission up to $J = 24 \rightarrow 23$ or even higher (Manoj et al. 2013; Karska et al. 2013; Green et al. 2013). These sources consistently show a 300-400 K component in this frequency range (van Kempen et al. 2010), attributed to UV irradiation of the outflow cavity walls and C-shocks (Visser et al. 2012). This component is usually referred to as “warm”. Many embedded sources also exhibit a “hot” component detected in CO transitions $J_{\text{up}} \geq 25$, ranging from 700-900 K. We use the same definitions in this work, and add the “cool” component to identify fits to CO transitions $J_{\text{up}} \leq 13$; the source of the cool emission is usually attributed to passive heating (from gas particle collisions with dust grains heated by stellar radiation; Visser et al. 2012).

We do not observe a warm component in most of these systems. Of the FUors, only V1057 Cyg shows robust cool and warm components at 81 and 368 K, respectively. Lorenzetti et al. (2000) noted a detection of CO $J = 17 \rightarrow 16$ for V1057 Cyg only, but no higher or lower- J transitions, using ISO-LWS. We also fit cool components to V1331 Cyg and V1515 Cyg and find T_{rot} of 30 and 19 K, respectively. We detect only one CO line in this range in FU Ori ($J = 5 \rightarrow 4$) from HIFI, and thus cannot compute T_{rot} .

We also fit a single component to the CO and ^{13}CO rotational diagrams for V1735 Cyg. We find T_{rot} of 67 K and 45 K, respectively, but do not attribute the emission solely to the FUor.

As the highest observed rotational level of CO emission decreases, the rotational temperature of gas decreases as well, from 81 K in V1057 Cyg, to 19 K in V1515 Cyg. The total amount of CO gas, $\mathcal{N}(\text{CO})$, dominated in the FUors by the cool CO, does not follow this trend but instead varies more like L_{bol} . This is consistent with $\mathcal{N}(\text{CO})$ derived from the warm component in embedded sources (Manoj et al. 2013; Karska et al. 2013; Green et al. 2013). If the gas around V1057 Cyg is well-approximated by a two-temperature (cool/warm) fit, optical depth effects would be most apparent at low- J , where the line flux would fall below the fit, for certain density regimes (Goldsmith & Langer 1999, Fig. 12). Instead, the CO continues to exhibit positive curvature to the longest SPIRE wavelengths. This suggests that either the gas is populated at low densities (Neufeld 2012; Manoj et al. 2013), that optical depth effects are not substantially affecting the V1057 Cyg line fluxes even at low- J , or that there is a larger reservoir of cooler gas.

The best constraint on opacity arises from the ratio of ^{13}CO to ^{12}CO : in V1057 Cyg, the upper limit of the isotopic ratio of the $J = 5 \rightarrow 4$ line intensity is ~ 10 , requiring an optical depth less than ~ 7 . For comparison,

the CO optical depth measured from the $^{12}\text{CO}/^{13}\text{CO}$ ratio declines with upper state J for both confused sources V1735 Cyg and HBC 722, from τ of ~ 7 at $J = 5 \rightarrow 4$ down to ~ 3 at $J = 7 \rightarrow 6$.

In the limit of moderate or higher density, the observed curvature suggests that additional “cold” components, or even a power-law distribution of components may be needed to fit the rotational diagram, in order to produce sufficient low- J emission. It has already been suggested that a power law distribution is a good fit to the CO $J = 14 \rightarrow 13$ up to $J = 49 \rightarrow 48$ PACS lines detected in the Orion protostar sample (Manoj et al. 2013); whether this applies to the FUors is not clear.

4.2.1. Velocity resolved mid- J emission

Using HIFI, CO $J = 14 \rightarrow 13$ is weakly detected in HBC 722, likely contaminated by extended emission; it is not detected in any other object. We also observed $\text{HCO}^+ J = 3 \rightarrow 2$, a dense gas tracer, using heterodyne observations at the CSO, resolving the line profiles (Figure 4), for all sources except FU Ori. We compared these profiles to CO $J = 5 \rightarrow 4$ resolved lines from HIFI for each of V1057 Cyg, V1735 Cyg, V1331 Cyg, V1515 Cyg, and HBC 722. Additionally, CO $J = 3 \rightarrow 2$ and $^{13}\text{CO } J = 2 \rightarrow 1$ profiles for these sources appear in McMudroch et al. (1993), Evans et al. (1994), and McMudroch et al. (1995). In the case of V1331 Cyg, the CO and HCO^+ exhibited similar profiles, with narrow features. In the other four sources, the HCO^+ narrow feature peaked at the velocity of self-absorption in CO; the CO additionally showed wings much broader than those seen in HCO^+ . In fact, the self-absorption dip is the clearest indicator of the source velocity, and the HCO^+ peak is slightly redshifted in comparison to the absorption, indicating that some of the contributing gas in the red part of the wings is at the source velocity; the wings are typical of embedded sources (e.g. Jørgensen et al. 2009).

The sources with stronger high- J CO also exhibit broader CO $J = 5 \rightarrow 4$ profiles; the HCO^+ is always a narrow single peak in all sources. In V1735 Cyg and HBC 722, the broad CO traces contaminating outflows, while the HCO^+ measures the velocity of the ambient dense gas. To the extent that HCO^+ aligns with the absorption in CO, we presume that low excitation ambient gas is present at that velocity.

In the case of V1515 Cyg, we see little excited gas, consistent with earlier observations (Evans et al. 1994; Kóspál 2011). In the cases of FU Ori and V1331 Cyg, we interpret the weak, single-peaked CO as tracing the source. In the case of V1057 Cyg, we observe a narrower but still somewhat broad feature. V1057 Cyg also has the richest PACS spectrum; in this case the CO emission may originate from material at the source, rather than contaminating outflows. In this case, the outflow and the high- J CO may be driven from V1057 Cyg itself.

Thus the mid-IR non-*Herschel* SEDs are not good predictors of which emission lines would be detected at longer wavelengths. Despite similar mid-IR spectral features, shape, and extinction parameters, V1515 Cyg and V1057 Cyg are quite different from the *Herschel* perspective.

4.2.2. Where is the Hot CO Emission in FUors?

We might posit that CO emission from hot gas would be a transitory signature of heating from a recent outburst. How then do we explain the lack of high- J CO emission in these FUors?

Analysis of the *non-detections* reveals that the absence of high- J CO and all but one H₂O line in our sources is of only modest significance. Figure 15 shows histograms of the line luminosities of the CO $J = 16 \rightarrow 15$ (left) and H₂O 174.63 μm (right) for the DIGIT sample (Green et al. 2013). The vertical dashed lines show the strength of these lines in V1057 Cyg; the vertical dotted lines show 3σ upper limits for the non-detections in the FOOSH sample. In both cases, the detected CO $J = 16 \rightarrow 15$ line (in V1057 Cyg) would be among the brightest in the DIGIT protostar sample, but not a notable outlier. The upper limits from our sample for CO and H₂O are not low enough to rule out significant emission in any source.

If the CO emission originates in a tenuous envelope, there may not be enough emitting material to produce a detectable signature in the warmer regions. If the CO emission originates in a disk, then we require a thermal inversion to produce emission lines in the upper layer. Models suggest that the midplane of FUors may become hotter than the upper disk layers during outburst, which would mask the signature of CO (Zhu et al. 2010). Our data do not sufficiently constrain the line luminosities to distinguish the FUors from ordinary protostars.

In the DIGIT sample, the CO emission from disks is much weaker. The highest line luminosity from CO $J = 16 \rightarrow 15$ detected in the DIGIT disks sample is in HD 100546, a Herbig Ae/Be star of similar luminosity to the FUors (Sturm et al. 2010). The CO line is a factor of 20 weaker in luminosity than V1057 Cyg, at $3 \times 10^{-5} L_{\odot}$ (Meeus et al., *subm.*). Thus it is unlikely that the disk is a significant contributor to the CO line luminosity in V1057 Cyg.

4.3. Other Molecules

We consider here a few other marginal detections. We report the detection of a single line of H₂O at 174.63 μm in V1057 Cyg. The H₂O line luminosity in V1057 Cyg is $0.47 \times 10^{-3} L_{\odot}$, which would place this single line among the most luminous of the DIGIT sample. The 174.63 μm line (excitation energy of 197 K) is not even the lowest excitation line within the PACS bands; in embedded sources the lowest energy detected H₂O line is at 179.53 μm (114 K); however the 174.63 μm line is typically among the most luminous H₂O lines detected in Class 0/I protostars. We also do not detect the 557 GHz $1_{10} - 1_{01}$ line in the SPIRE bands. We do not detect H₂O in any other sources. Although the detection in V1057 Cyg is robust, the upper limits on this line in the other FUors do not rule out substantial emission lines as well.

We tentatively detect OH 84.41 μm , half of a doublet in the $3/2 \rightarrow 3/2$ ladder, also in V1057 Cyg. Although this is the most commonly detected line in the embedded sample, the 84.61 μm half of the doublet is typically just as strong. More prominently, OH 119 μm is seen *in absorption* in V1735 Cyg. Only V1735 Cyg and 2 of 30 Class 0/I sources in the DIGIT sample show this line in absorption. There is also a hint of OH emission at 163 μm , possibly blended with a weak detection of the CO

$J = 16 \rightarrow 15$ line, but both detections are inconclusive.

No other molecules are detected in the FOOSH sample with *Herschel*. In contrast, at least some embedded sources show rich SPIRE spectra (Goicoechea et al. 2012, Green et al., *in prep.*). However, the upper limits for H₂O and OH line luminosity are consistent with the detections in embedded sources.

4.4. Additional Lines

[C II] 157.74 μm is detected in V1735 Cyg, V1331 Cyg, and V1515 Cyg, but is not clearly associated with the sources, similar to the DIGIT embedded sources.

[C I] is observed at 370 and 610 μm in all sources. The [C I] 370/610 μm flux ratio falls between 2.1 and 3.0 across the sample, consistent with ISM values (Jenkins & Tripp 2001), and the [C I] emission is associated with cooler extended gas rather than the FUors themselves. The SPIRE maps show that the lines are extended (at beamsize $\sim 40''$); we do not attribute the emission to the FUors.

[N II] 205 μm (SPIRE) is observed to emit strongly in HBC 722, V1331 Cyg, V1057 Cyg, and V1515 Cyg, but not V1735 Cyg or FU Ori. The [N II] 122 μm (PACS) line is detected along with the 205 μm line around V1515 Cyg; no other sources in our sample show the 122 μm line. This is in contrast with ISO-LWS results, in which [N II] was marginally detected at 122 μm with ISO-LWS in V1735 Cyg, V1331 Cyg, V1057 Cyg (Lorenzetti 2005). As noted earlier, the difference in detection rates of [N II] 122 μm is likely due to the use of an off-position for our PACS data, but an onboard calibration source for SPIRE.

We detected both lines around V1515 Cyg, and examined the PACS and SPIRE maps to see if there was any hint of centrally peaked emission. The flux of each line shows no significant variation ($< 50\%$) across the entire PACS and SPIRE field-of-view, and is uncorrelated with proximity to the central position. The [N II] 122/205 μm flux ratio is diagnostic of the electron density (Hudson & Bell 2004; Oberst et al. 2011); The ratio in the central pixel is 0.31, suggesting an extremely low electron density ($n_e < 1 \text{ cm}^{-3}$). We note that the observed ratio should be considered an upper limit, as the PACS spectra are observed using an off-position for sky subtraction, while the SPIRE spectra are not.

Further evidence of the diffuse nature of the [N II] emission comes from the fact that there is no clear trend between [N II] 205 μm and the strength or position of the CO emission. Thus we conclude that among fine structure lines, only the [O I] emission is localized to the FUor.

5. DISCUSSION

5.1. Evolutionary Stage

It has been posited from occurrence rate (Hartmann & Kenyon 1996) and from simulation (e.g. Dunham & Vorobyov 2012) that FUors are Stage 0/I stars undergoing episodic outburst cycles with replenished material from an infalling envelope. However, FUors have much in common with Class II sources: in Spitzer-IRS bands all but V1735 Cyg show silicate emission and optical-IR continuum indices consistent with relatively low extinction (e.g. Hartmann & Kenyon 1996; Green et al. 2006; Quanz et al. 2007c). It has been

suggested that the overall group of FUors fall into two different categories, those consistent with flared disks, and those requiring envelopes (e.g. Kenyon & Hartmann 1991; Zhu et al. 2008).

Zhu et al. (2008) modeled the dust from Spitzer-IRS observations of FU Ori, V1057 Cyg, and V1515 Cyg and found that V1057 Cyg and V1515 Cyg required envelopes typical of protostars, while FU Ori did not require an envelope. The *Herschel* spectra generally fall above their model predictions, indicating an underestimate of the reservoir of cold dust surrounding these systems. Nonetheless, our estimates of the envelope mass from the *Herschel* dust emission broadly agrees with this classification; we find that V1057 Cyg and V1515 Cyg have envelopes greater than $0.1 M_{\odot}$, while FU Ori has a much smaller envelope mass of $\sim 0.02 M_{\odot}$. HBC 722 was considered a T Tauri star prior to outburst, and measurements of the envelope mass upper limit suggest a tenuous envelope at best.

We noted some inconsistencies between the different classification methods for the continuum, but found broad agreement with the mid-IR sequence in Quanz et al. (2007c). However, FUors may not be well-characterized by the Stage I/II sequence, because the *Herschel* line observations provide a different picture from the continuum. In general, the presence of weaker silicate features and larger mid-IR α is correlated with increased submillimeter dust and CO/H₂O emission, but is not well-correlated with [O I]. Class 0/I sources typically exhibit hotter and more excited CO rotational lines than Class II sources (cf. Karska et al. 2013, Meeus et al. *subm.*); by this standard FUors would *seem* to resemble disk sources; however, these lines may be too faint to detect in our observations. The DIGIT (and HOPS) samples show a tight correlation between CO $J = 16 \rightarrow 15$ line luminosity and derived $\mathcal{N}(\text{warm})$. In V1057 Cyg, the only FUor where the warm gas is detected, it falls in the DIGIT embedded source correlation. Additionally the DIGIT sample showed a lack of correlation between $T_{\text{rot}}(\text{warm})$ and $\mathcal{N}(\text{CO})$; again, the lone FUor (V1057 Cyg) is consistent with the embedded sample. Finally, the DIGIT sample showed a (weaker) trend between L_{bol} and $\mathcal{N}(\text{warm})$; once again V1057 Cyg fits this trend. It is clear that the spectra of V1057 Cyg, as characterized by PACS, would not distinguish it as an outlier in the DIGIT Class 0/I sample, even though it has a Class II SED, suggesting it is a kind of hybrid object. However, the high upper limits on line emission do not provide strong constraints on the evolutionary state of the other FUors.

6. CONCLUSIONS

We present analysis of six FUors with *Herschel*-PACS, SPIRE, and HIFI, *Spitzer*-IRS, and CSO heterodyne observations, from the “FOOSH” sample.

- We present PACS and SPIRE imaging of V1057 Cyg, V1331 Cyg, V1515 Cyg, V1735 Cyg, and FU Ori, and their surrounding regions. Complementary to pre-existing near-IR and submillimeter ground-based imaging, all show filamentary morphologies. In the case of V1515 Cyg, the optical FUor is only marginally detected at SPIRE wavelengths. Most fall in complicated far-IR fields, typ-

ical of active star-forming regions, but only two (V1735 Cyg and HBC 722) are confused. V1735 Cyg is blended with V1735 Cyg SM1, a nearby submillimeter source, at wavelengths greater than $100 \mu\text{m}$, and HBC 722, as noted previously, is blended with 2MASS 20581767+4353310 at wavelengths greater than $70 \mu\text{m}$.

- The 1-1000 μm SEDs for the five older FUors indicate that FU Ori, V1515 Cyg, and V1057 Cyg are Class II, while V1735 Cyg and V1331 Cyg are borderline Class I/II, or flat spectrum sources, classified by T_{bol} . We derive envelope masses of $0.07\text{--}0.3 M_{\odot}$ from the long wavelength data, and find FU Ori and V1331 Cyg to be Stage II, V1515 Cyg to be borderline Stage I/II, and V1057 Cyg to be Stage I by the Crapsi et al. (2008) criterion (envelope mass $> 0.1 M_{\odot}$). Their derived L_{bol} values are somewhat lower than previously measured values, likely a complication of the replacement of the large IRAS beamsize with the smaller PACS/SPIRE beam in our SEDs.
- All exhibit [O I] $63 \mu\text{m}$ and [C I] 370 and $610 \mu\text{m}$ lines. [N II] 122 and $205 \mu\text{m}$ and [C II] $158 \mu\text{m}$ are detected as well; the latter two lines are detected in four of the sources while the former is only detected in one of those (V1515 Cyg). Of these lines, only [O I] ($63 \mu\text{m}$ and $145 \mu\text{m}$, when detected) is attributed to the FUors. HBC 722 is detected only weakly in continuum at $70 \mu\text{m}$, and cleanly detected in line emission only in [O I] $63 \mu\text{m}$.
- Analysis of V1735 Cyg is complicated by the presence of a nearby submillimeter source (V1735 Cyg SM1), and the contributions from the two sources to gas lines other than [O I] cannot be separated. This includes four mid- J transitions of ^{13}CO as well as nine transitions of CO. Mapping indicates that the detected CO emission is dominated by SM1. The low-lying OH $119 \mu\text{m}$ doublet is observed in absorption toward V1735 Cyg; by comparison, this is seen in 2 of 30 protostars in the DIGIT sample. The sources are separable in continuum emission out to $200 \mu\text{m}$ with *Herschel*, and we use ground-based observations to help constrain longer wavelengths.
- V1057 Cyg is an outlier among the FOOSH sample. We detect CO up to $J = 23 \rightarrow 22$ and determine that the emitting source is compact. We fit cool (81 K) and warm (368 K) rotational temperature components to the CO. We detect H₂O in one transition at $174.63 \mu\text{m}$, which would place this source among the most luminous protostars in line emission in the DIGIT sample, but not an extreme outlier. V1057 Cyg also shows a tentative detection of OH $84.41 \mu\text{m}$.
- In the other three sources (V1331 Cyg, V1515 Cyg, and FU Ori), no H₂O or OH lines are detected, but upper limits are still consistent with the detections for protostars. CO is detected up to $J = 6 \rightarrow 5$ in these sources. We fit cool components of 30 and 19 K to V1331 Cyg and V1515 Cyg, respectively;

no component can be fit for FU Ori, which is detected in only a single CO transition ($J = 5 \rightarrow 4$). Although V1057 Cyg and V1515 Cyg are very similar from the *Spitzer* perspective (in terms of their SED, extinction values, and dust properties), they are differentiated by their gas properties in the *Herschel* range: the former exhibits high- J CO, H₂O and luminous [O I]; the latter shows no molecular emission and only weak [O I] traceable to the source.

- Dense gas is present at the source velocity in all four sources where it was observed, based on HCO⁺ $J = 3 \rightarrow 2$. In one case, V1331 Cyg, the HCO⁺ profile matches the CO $J = 5 \rightarrow 4$ from HIFI suggesting that the CO may originate on-source. In V1057 Cyg, V1735 Cyg, and HBC 722, the CO $J = 5 \rightarrow 4$ line is broader than the HCO⁺, with hints of self-absorption at the velocity of the dense gas, suggesting that the CO may partially trace outflows in the region. CO $J = 14 \rightarrow 13$ is marginally detected in HBC 722 but contaminated by extended emission; it is not detected in any other objects. Only in V1057 Cyg is the high- J CO is suspected to originate at the source.
- The FUors all exhibit large [O I] line luminosities, which we attribute to increased UV radiation and outflow rates, via episodic accretion. The mass loss rate derived from [O I] assuming a 100 km s⁻¹ wind and the accretion luminosity from continuum are consistent (assuming a conversion factor of 10 to 1 from infall to outflow), except in the case of HBC 722, whose [O I] implied mass accretion rate exceeds the rate derived from the accretion luminosity by two orders of magnitude. This suggests that HBC 722 may have experienced previous outbursts, as the [O I] could not have propagated far enough for a significant enhancement in first three months of outburst prior to observations.

From the *Herschel* perspective, FUors outbursts are associated with both Stage I and II sources; the evolutionary state of the sources is somewhat variable depending upon the criteria used. They are invariably surrounded by considerable nebulosity and cold dust, typically associated with protostars rather than Class II sources. Their

continuum SEDs are not uniform, with some exhibiting “flat-spectrum” and others Class II features. The sources with the lowest IR excess and most prominent mid-IR silicate dust features – FU Ori, V1331 Cyg, and V1515 Cyg – exhibit little or no excited CO gas, typical of most disk sources. V1057 Cyg shows a cool and a warm gas component, consistent with embedded protostars, although not a detectable hot component ($T_{\text{rot}} > 500$ K). The H₂O and OH detections in V1057 Cyg are typical for protostars but much more luminous than those seen in Herbig Ae/Be stars (Fedele et al., *subm.*) and T Tauri stars (Salyk et al., *in prep.*). The defining characteristic from *Herschel* is that all FUors exhibit large [O I] 63 μm line luminosities, indicative of high accretion rates.

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REFERENCES

- Ábrahám, P., Kóspál, Á., Csizmadia, S., et al. 2004, *A&A*, 428, 89
 Andre, P., Ward-Thompson, D., & Barsony, M. 1993, *ApJ*, 406, 122
 Andrews, S. M., & Williams, J. P. 2005, *ApJ*, 631, 1134
 Bell, K. R., & Lin, D. N. C. 1994, *ApJ*, 427, 987
 Carr, J. S., & Najita, J. R. 2011, *ApJ*, 733, 102
 Chen, H., Myers, P. C., Ladd, E. F., & Wood, D. O. S. 1995, *ApJ*, 445, 377
 Crapsi, A., van Dishoeck, E. F., Hogerheijde, M. R., Pontoppidan, K. M., & Dullemond, C. P. 2008, *A&A*, 486, 245
 Cohen, M., & Kuhl, L. V. 1979, *ApJS*, 41, 743
 de Graauw, T., Helmich, F. P., Phillips, T. G., et al. 2010, *A&A*, 518, L6
 Dunham, M. M., Arce, H. G., Bourke, T. L., et al. 2012, *ApJ*, 755, 157
 Dunham, M. M., Evans, N. J., Terebey, S., Dullemond, C. P., & Young, C. H. 2010, *ApJ*, 710, 470
 Dunham, M. K., Rosolowsky, E., Evans, N. J., II, Cyganowski, C., & Urquhart, J. S. 2011, *ApJ*, 741, 110
 Dunham, M. M., & Vorobyov, E. I. 2012, *ApJ*, 747, 52
 Elias, J. H. 1978, *ApJ*, 223, 859
 Evans, N. J., Dunham, M. M., Jørgensen, J. K., et al. 2009, *ApJS*, 181, 321
 Evans, II, N. J., Balkum, S., Levreault, R. M., Hartmann, L., & Kenyon, S. 1994, *ApJ*, 424, 793
 Fischer, W. J., Megeath, S. T., Tobin, J. J., et al. 2012, *ApJ*, 756, 99
 Fischer, W. J., Megeath, S. T., Stutz, A. M., et al. 2013, *Astronomische Nachrichten*, 334, 53
 Furlan, E., Hartmann, L., Calvet, N., et al. 2006, *ApJS*, 165, 568

- Goicoechea, J. R., Cernicharo, J., Karska, A., et al. 2012, *A&A*, 1209.3966
- Goldsmith, P. F., & Langer, W. D. 1999, *ApJ*, 517, 209
- Green, J. D., Hartmann, L., Calvet, N., et al. 2006, *ApJ*, 648, 1099
- Green, J. D., Evans, II, N. J., Kóspál, Á., et al. 2011, *ApJ*, 731, L25
- Green, J. D., Robertson, P., Baek, G., et al. 2013, *ApJ*, 764, 22
- Green, J. D., Evans, N. J., II, Jørgensen, J. K., et al. 2013, arXiv:1304.7389, *ApJ*, in press
- Greene, T. P., Wilking, B. A., Andre, P., Young, E. T., & Lada, C. J. 1994, *ApJ*, 434, 614
- Griffin, M. J., Abergel, A., Abreu, A., et al. 2010, *A&A*, 518, L3+
- Hartmann, L. 1998, *Accretion Processes in Star Formation (Accretion processes in star formation / Lee Hartmann. Cambridge, UK ; New York : Cambridge University Press, 1998. (Cambridge astrophysics series ; 32) ISBN 0521435072.)*
- Hartmann, L., & Kenyon, S. J. 1985, *ApJ*, 299, 462
- Hartmann, L., & Kenyon, S. J. 1996, *ARA&A*, 34, 207
- Harvey, P. M., Huard, T. L., Jørgensen, J. K., et al. 2008, *ApJ*, 680, 495
- Herbig, G. H. 1977, *ApJ*, 217, 693
- Herczeg, G. J., Karska, A., Bruderer, S., et al. 2012, *A&A*, 540, A84
- Hollenbach, D. 1985, *Icarus*, 61, 36
- Houck, J. R., Roellig, T. L., van Cleve, J., et al. 2004, *ApJS*, 154, 18
- Hudson, C. E., & Bell, K. L. 2004, *MNRAS*, 348, 1275
- Jenkins, E. B., & Tripp, T. M. 2001, *ApJS*, 137, 297
- Johnstone, D., Hendricks, B., Herczeg, G. J., & Bruderer, S. 2013, *ApJ*, 765, 133
- Jørgensen, J. K., van Dishoeck, E. F., Visser, R., et al. 2009, *A&A*, 507, 861
- Karska, A., Herczeg, G. J., van Dishoeck, E. F., et al. 2013, arXiv:1301.4821
- Kenyon, S. J., & Hartmann, L. W. 1991, *ApJ*, 383, 664
- Kenyon, S. J., Kolotilov, E. A., Ibragimov, M. A., & Mattei, J. A. 2000, *ApJ*, 531, 1028
- Kóspál, Á. 2011, *A&A*, 535, A125
- Kóspál, Á., Abraham, P., Acosta-Pulido, J. A., et al. 2011, *A&A*, 527, A133
- Kristensen, L. E., van Dishoeck, E. F., Bergin, E. A., et al. 2012, *A&A*, 542, A8
- Kurosawa, R., Harries, T. J., & Symington, N. H. 2006, *MNRAS*, 370, 580
- Laugalys, V., Straizys, V., Vrba, F. J., et al. 2006, *Baltic Astronomy*, 15, 483
- Lin, D. N. C., & Papaloizou, J. 1985, in *Protostars and Planets II*, ed. D. C. Black & M. S. Matthews, 981–1072
- Lorenzetti, D. 2005, *Space Sci. Rev.*, 119, 181
- Lorenzetti, D., Giannini, T., Nisini, B., et al. 2000, *A&A*, 357, 1035
- Manoj, P., Watson, D. M., Neufeld, D. A., et al. 2013, *ApJ*, 763, 83
- Mathis, J. S. 1990, *ARA&A*, 28, 37
- McMuldach, S., Sargent, A. I., & Blake, G. A. 1993, *AJ*, 106, 2477
- McMuldach, S., Blake, G. A., & Sargent, A. I. 1995, *AJ*, 110, 354
- Miller, A. A., Hillenbrand, L. A., Covey, K. R., et al. 2011, *ApJ*, 730, 80
- Najita, J. R., Doppmann, G. W., Carr, J. S., Graham, J. R., & Eisner, J. A. 2009, *ApJ*, 691, 738
- Neufeld, D. A. 2012, *ApJ*, 749, 125
- Oberst, T. E., Parshley, S. C., Nikola, T., et al. 2011, *ApJ*, 739, 100
- Ossenkopf, V., & Henning, T. 1994, *A&A*, 291, 943
- Ott, S. 2010, in *Astronomical Society of the Pacific Conference Series, Vol. 434, Astronomical Data Analysis Software and Systems XIX*, ed. Y. Mizumoto, K.-I. Morita, & M. Ohishi, 139
- Paczynski, B. 1976, in *IAU Symposium, Vol. 73, Structure and Evolution of Close Binary Systems*, ed. P. Eggleton, S. Mitton, & J. Whelan, 75–+
- Pilbratt, G. L., Riedinger, J. R., Passvogel, T., et al. 2010, *A&A*, 518, L1
- Poglitich, A., Waelkens, C., Geis, N., et al. 2010, *A&A*, 518, L2
- Pueyo, L., Hillenbrand, L., Vasisht, G., et al. 2012, *ApJ*, 757, 57
- Quanz, S. P., Apai, D., & Henning, T. 2007a, *ApJ*, 656, 287
- Quanz, S. P., Henning, T., Bouwman, J., Linz, H., & Lahuis, F. 2007b, *ApJ*, 658, 487
- Quanz, S. P., Henning, T., Bouwman, J., et al. 2007c, *ApJ*, 668, 359
- Reipurth, B., & Aspin, C. 2004, *ApJ*, 608, L65
- Reipurth, B., & Aspin, C. 2010, in *Evolution of Cosmic Objects through their Physical Activity*, ed. H. A. Harutyunian, A. M. Mickaelian, & Y. Terzian, 19–38
- Reipurth, B., Aspin, C., & Herbig, G. H. 2012, *ApJ*, 748, L5
- Robitaille, T. P., Whitney, B. A., Indebetouw, R., Wood, K., & Denzmore, P. 2006, *ApJS*, 167, 256
- Roelfsema, P. R., Helmich, F. P., Teyssier, D., et al. 2012, *A&A*, 537, A17
- Sandell, G., & Weintraub, D. A. 2001, *ApJS*, 134, 115
- Semkov, E., & Peneva, S. 2010, *The Astronomer's Telegram*, 2801, 1
- Semkov, E. H., Peneva, S. P., Munari, U., et al. 2012, *A&A*, 542, A43
- Sturm, B., Bouwman, J., Henning, T., et al. 2010, *A&A*, 518, L129
- Sturm, B., Bouwman, J., Henning, T., et al. 2013, *A&A*, 553, A5
- van Kempen, T. A., Kristensen, L. E., Herczeg, G. J., et al. 2010, *A&A*, 518, L121
- Visser, R., Kristensen, L. E., Bruderer, S., et al. 2012, *A&A*, 537, A55
- Wang, H., Apai, D., Henning, T., & Pascucci, I. 2004, *ApJ*, 601, L83
- Weintraub, D. A., Sandell, G., & Duncan, W. D. 1991, *ApJ*, 382, 270
- Welin, G. 1977, *Information Bulletin on Variable Stars*, 1314, 1
- Werner, M. W., Roellig, T. L., Low, F. J., et al. 2004, *ApJS*, 154, 1
- Zhu, Z., Hartmann, L., Calvet, N., et al. 2008, *ApJ*, 684, 1281
- Zhu, Z., Hartmann, L., Gammie, C. F., et al. 2010, *ApJ*, 713, 1134

TABLE 1
HERSCHEL OBSERVING LOG

Object	Mode	AOR	Date
HBC 722	HIFI CO 14-13	1342210779	3 Dec 2010
HBC 722	HIFI CO 5-4	1342210806	4 Dec 2010
HBC 722	SPIRE Spec.	1342210857	6 Dec 2010
HBC 722	SPIRE Phot.	1342210915	8 Dec 2010
HBC 722	PACS Phot. Blue1	1342211094	13 Dec 2010
HBC 722	PACS Phot. Blue2	1342211095	13 Dec 2010
HBC 722	PACS Phot. Green1	1342211096	13 Dec 2010
HBC 722	PACS Phot. Green2	1342211097	13 Dec 2010
HBC 722	PACS Spec. Red	1342211173	14 Dec 2010
HBC 722	PACS Spec. Blue	1342211174	14 Dec 2010
V1515 Cyg	SPIRE Phot.	1342211353	20 Dec 2010
V1057 Cyg	SPIRE Phot.	1342211357	20 Dec 2010
V1735 Cyg	SPIRE Spec.	1342219560	24 Apr 2011
V1735 Cyg	SPIRE Phot.	1342219973	6 May 2011
V1331 Cyg	SPIRE Phot.	1342220630	8 May 2011
V1735 Cyg	HIFI CO 14-13	1342220505	12 May 2011
HBC 722	HIFI CO 5-4	1342221418	20 May 2011
V1515 Cyg	HIFI CO 14-13	1342221450	20 May 2011
V1331 Cyg	HIFI CO 14-13	1342221451	20 May 2011
V1057 Cyg	HIFI CO 14-13	1342221452	20 May 2011
HBC 722	HIFI CO 14-13	1342221453	20 May 2011
V1515 Cyg	SPIRE Spec.	1342221685	25 May 2011
V1331 Cyg	SPIRE Spec.	1342221694	25 May 2011
V1057 Cyg	SPIRE Spec.	1342221695	25 May 2011
V1057 Cyg	PACS Phot. Green1	1342223184	23 Jun 2011
V1057 Cyg	PACS Phot. Blue1	1342223185	23 Jun 2011
V1057 Cyg	PACS Phot. Blue2	1342223186	23 Jun 2011
V1057 Cyg	PACS Phot. Green2	1342223187	23 Jun 2011
V1735 Cyg	PACS Phot. Green1	1342225246	22 Jul 2011
V1735 Cyg	PACS Phot. Blue1	1342225247	22 Jul 2011
V1735 Cyg	PACS Phot. Blue2	1342225248	22 Jul 2011
V1735 Cyg	PACS Phot. Green2	1342225249	22 Jul 2011
V1331 Cyg	PACS Phot. Green1	1342225252	22 Jul 2011
V1331 Cyg	PACS Phot. Blue1	1342225253	22 Jul 2011
V1331 Cyg	PACS Phot. Blue2	1342225254	22 Jul 2011
V1331 Cyg	PACS Phot. Green2	1342225255	22 Jul 2011
FU Ori	HIFI CO 5-4	1342228625	15 Sep 2011
FU Ori Cyg	SPIRE Phot.	1342219654	23 Sep 2011
FU Ori	SPIRE Spec.	1342230412	09 Oct 2011
V1515 Cyg	PACS Phot. Blue1	1342232439	17 Nov 2011
V1515 Cyg	PACS Phot. Green1	1342232440	17 Nov 2011
V1515 Cyg	PACS Phot. Green2	1342232441	17 Nov 2011
V1515 Cyg	PACS Phot. Blue2	1342232442	17 Nov 2011
V1735 Cyg	HIFI CO 5-4	1342232695	23 Nov 2011
V1057 Cyg	HIFI CO 5-4	1342232698	23 Nov 2011
V1331 Cyg	PACS Spec. Red	1342233445	2 Dec 2011
V1331 Cyg	PACS Spec. Blue	1342233446	2 Dec 2011
V1515 Cyg	PACS Spec. Red	1342235690	28 Dec 2011
V1515 Cyg	PACS Spec. Blue	1342235691	28 Dec 2011
V1735 Cyg	PACS Spec. Red	1342235848	28 Dec 2011
V1735 Cyg	PACS Spec. Blue	1342235849	31 Dec 2011
V1057 Cyg	PACS Spec. Red	1342235852	1 Jan 2012
V1057 Cyg	PACS Spec. Blue	1342235853	1 Jan 2012
FU Ori	PACS Phot. Blue1	1342242658	29 Mar 2012
FU Ori	PACS Phot. Green1	1342242659	29 Mar 2012
FU Ori	PACS Phot. Blue2	1342242660	29 Mar 2012
FU Ori	PACS Phot. Green2	1342242661	29 Mar 2012
V1515 Cyg	HIFI CO 5-4	1342244405	14 Apr 2012
V1331 Cyg	HIFI CO 5-4	1342244406	14 Apr 2012
FU Ori	PACS Spec. Blue	1342250907	11 Sep 2012
FU Ori	PACS Spec. Red	1342250908	11 Sep 2012
HBC 722	HIFI CO 14-13	1342255774	22 Nov 2012

NOTE. — Each SPIRE image and spectrum includes a single observation ID (OBSID). Each PACS spectral scan consists of a “blue” (50-75 and 100-150 μm) and “red” (70-100 and 140-200 μm) pair of observations. Each PACS image consists of a pair of cross-scans (1 and 2) for each of the “blue” (70 and 160 μm) and “green” (100 and 160 μm) bands.

TABLE 2
SAMPLE DESCRIPTION

Source	RA (J2000)	Dec	Dist. (pc)	Date of Eruption	Mid-IR Class	Binary?	Ref.
FU Ori	05:45:22.4	+09:04:12	600	1936	disk	few AU binary	Herbig (1977)
V1515 Cyg	20:23:48.0	+42:12:26	1000	~ 1950	some env.	—	Herbig (1977)
HBC 722	20:53:17.0	+43:53:43	520	2009-2010	—	—	Laugalys et al. (2006)
V1057 Cyg	20:58:53.7	+44:15:29	600	1969	some env.	—	Herbig (1977)
V1735 Cyg	21:47:20.7	+47:32:04	950	1952-1965	env.	—	Elias (1978)
V1331 Cyg	21:01:09.2	+50:21:45	550	—	disk	—	Quanz et al. (2007a)

NOTE. — Sample description. A portion of the table is from Reipurth & Aspin (2010). The Mid-IR class is described in the text.

TABLE 3
PACS/SPIRE FLUX DENSITY

Object	70 μm	100 μm	160 μm	250 μm	350 μm	500 μm
	Jy	Jy	Jy	Jy	Jy	Jy
V1057 Cyg	30.0 \pm 3.0	30.1 \pm 3.0	28.2 \pm 2.8	15.8 \pm 1.6	8.1 \pm 1.2	2.8 \pm 0.6
V1331 Cyg	7.5 \pm 0.8	9.7 \pm 1.0	11.6 \pm 1.2	9.9 \pm 1.1	6.4 \pm 1.0	2.7 \pm 0.6
V1515 Cyg	4.6 \pm 0.5	3.9 \pm 0.4	5.0 \pm 0.5	4.3 \pm 0.8	2.1 \pm 0.6	0.9 \pm 0.4
V1735 Cyg	17.1 \pm 1.7	19.7 \pm 2.0	27.1 \pm 2.7	< 49.4 \pm 4.0	< 27.3 \pm 2.8	< 11.0 \pm 1.4
FU Ori	6.5 \pm 0.7	5.9 \pm 0.6	4.9 \pm 0.5	2.8 \pm 0.5	2.5 \pm 0.6	0.9 \pm 0.3
HBC 722 ¹	0.41 \pm 0.02	—	—	—	—	—

NOTE. — PACS 70, 100, and 160 μm , and SPIRE 250, 350, and 500 μm photometry of the five FUors in our sample, with the extended source correction applied, in 30'' apertures around the source coordinates. The SPIRE uncertainty includes a 7% calibration uncertainty, and sky subtraction from an annulus between 90-120'' from the source. The PACS formal uncertainty is smaller, but we assume a 10% minimum uncertainty to account for extended emission. V1057 Cyg, V1331 Cyg, V1515 Cyg, and FU Ori were resolved as separate sources with PACS and were extracted using an annulus that included all flux. V1735 Cyg includes a significant contribution from V1735 Cyg SM1 at SPIRE wavelengths (e.g. > 200 μm), a nearby submm source to east of V1735 Cyg, and thus the flux densities for $\lambda > 200 \mu\text{m}$ should be considered upper limits; in addition, there is some contamination at PACS wavelengths, where we find contributions from SM1 for annuli greater than 17''.

¹HBC 722 (Green et al. 2011) was not detected as a separate continuum source with PACS/SPIRE except at 70 μm .

TABLE 4
CLASSIFICATION

Object	α	T_{bol}	Env. Mass
	Class	Class	Stage
V1057 Cyg	FS	II	I
V1515 Cyg	II	II	I
V1331 Cyg	II	FS	I/II ¹
FU Ori	II	II	II
V1735 Cyg	II	FS	—
HBC 722	II ²	—	—

NOTE. — Classification of the FOOSH sources by three different techniques: mid-IR spectral index α , T_{bol} , and envelope mass (see text). “FS” indicates a flat-spectrum classification. Contaminated measurements are reported as “—”.

¹ V1331 Cyg can be classified as Stage I or II depending upon the wavelength used to measure the envelope mass; the inferred mass is close to the boundary of 0.10 M_{\odot} defined in Crapsi et al. (2008).

² α was measured pre-outburst for HBC 722 (Miller et al. 2011)

TABLE 5
 DUST-DERIVED ENVELOPE MASS

Object	$L_{\text{bol}}/L_{\text{smm}}$	T_{dust}	Env. Mass (250 μm)	(350 μm)	(500 μm)
		K	M_{\odot}	M_{\odot}	M_{\odot}
V1057 Cyg	975	43	0.16	0.25	0.30
V1515 Cyg	505	42	0.13	0.19	0.28
V1331 Cyg	159	38	0.11	0.07	0.46
FU Ori	5858	54	0.02	0.05	0.07

 NOTE. — Ratio of bolometric over submillimeter luminosity, dust temperature, and envelope mass calculated at each of the three SPIRE photometric bands. The large envelope mass derived at 500 μm in V1331 Cyg is attributed to extended/diffuse emission.

 TABLE 6
 LINE FLUXES

Line	Freq.	λ	HBC 722	V1735 Cyg	V1057 Cyg	V1331 Cyg	V1515 Cyg	FU Ori
			10^{-18}	10^{-18}	10^{-18}	10^{-18}	10^{-18}	10^{-18}
Units:	GHz	μm	W m^{-2}	W m^{-2}	W m^{-2}	W m^{-2}	W m^{-2}	W m^{-2}
CO $J = 4 \rightarrow 3$	461.0	650.25	33.99	31.54	9.44	6.73	7.58	< 1.72
[C I] $^3\text{P}_1 \rightarrow ^3\text{P}_0$	491.4	610.9	20.13	11.15	4.81	6.46	4.56	3.75
$^{13}\text{CO } J = 5 \rightarrow 4$	551.0	544.1	4.04	5.72	< 1.65	< 1.65	< 1.65	< 1.65
CO $J = 5 \rightarrow 4$	576.3	520.23	50.84	58.27	16.30	4.50	5.40	< 3.72 ¹
$^{13}\text{CO } J = 6 \rightarrow 5$	661.4	453.5	5.14	5.46	< 3.06	< 3.06	< 3.06	< 3.06
CO $J = 6 \rightarrow 5$	691.5	433.56	89.72	111.02	21.62	8.25	< 2.85	< 2.85
$^{13}\text{CO } J = 7 \rightarrow 6$	771.3	388.74	4.78	8.54	< 3.15	< 3.15	< 3.15	< 3.15
CO $J = 7 \rightarrow 6$	806.7	371.65	125.88	157.59	23.98	3.80	< 2.55	< 2.55
[C I] $^3\text{P}_2 \rightarrow ^3\text{P}_1$	809.9	370.42	43.07	31.56	10.39	8.74	8.98	4.91
$^{13}\text{CO } J = 8 \rightarrow 7$	881.5	340.18	< 3.15	4.42	< 2.61	< 2.61	< 2.61	< 2.61
CO $J = 8 \rightarrow 7$	921.8	325.23	138.10	162.27	28.94	< 9.06	< 9.06	< 9.06
CO $J = 9 \rightarrow 8$	1036.9	289.12	141.13	101.74	< 5.13	< 5.13	< 5.13	< 5.13
CO $J = 10 \rightarrow 9$	1152.0	260.24	158.23	76.49	32.30	< 18.18	< 18.18	< 18.18
CO $J = 11 \rightarrow 10$	1267.0	236.61	129.02	53.28	28.30	< 20.19	< 20.19	< 20.19
CO $J = 12 \rightarrow 11$	1382.0	216.93	106.28	13.17	< 15.61	< 15.61	< 15.61	< 15.61
[N II] $^3\text{P}_1 \rightarrow ^3\text{P}_0$	1461.2	205.18	132.77	< 39.55	119.38	69.09	289.03	< 39.55
CO $J = 13 \rightarrow 12$	1496.9	200.27	88.12	< 40.94	< 40.94	< 40.94	< 40.94	< 40.94
CO $J = 14 \rightarrow 13$	1611.7	186.01	81.56	< 28.74	37.50	< 28.74	< 28.74	< 28.74
o-H ₂ O $3_{03} \rightarrow 2_{12}$	1716.7	174.63	< 15.75	< 15.75	42.04	< 15.75	< 15.75	< 15.75
CO $J = 15 \rightarrow 14$	1726.6	173.63	82.01	50.79	64.06	< 15.75	< 15.75	< 15.75
CO $J = 16 \rightarrow 15$	1841.4	162.81	54.03	< 27.66	46.75	< 27.66	< 27.66	< 27.66
[C II] $^2\text{P}_{3/2} \rightarrow ^2\text{P}_{1/2}$	1900.6	157.74	< 22.44	137.33	< 22.44	25.57	471.85	< 22.44
CO $J = 17 \rightarrow 16$	1956.0	153.27	48.31	< 41.64	53.02	< 41.64	< 41.64	< 41.64
[O I] $^3\text{P}_0 \rightarrow ^3\text{P}_1$	2060.0	145.53	39.62	38.13	< 24.22	< 24.22	< 24.22	< 24.22
CO $J = 18 \rightarrow 17$	2070.7	144.78	29.20	< 24.22	62.14	< 20.67	< 20.67	< 20.67
CO $J = 19 \rightarrow 18$	2185.1	137.20	41.15	< 15.06	41.24	< 15.06	< 15.06	< 15.06
CO $J = 20 \rightarrow 19$	2299.6	130.37	26.11	< 11.85	38.01	< 11.85	< 11.85	< 11.85
CO $J = 21 \rightarrow 20$	2414.0	124.19	< 33.72	< 33.72	60.58	< 33.72	< 33.72	< 33.72
[N II] $^3\text{P}_2 \rightarrow ^3\text{P}_1$	2459.3	121.90	< 17.67	< 17.67	< 17.67	< 17.67	89.11	< 17.67
CO $J = 22 \rightarrow 21$	2528.2	118.58	< 16.92	< 16.92	42.94	< 16.92	< 16.92	< 16.92
CO $J = 23 \rightarrow 22$	2642.3	113.46	< 42.44	< 42.44	96.10	< 42.44	< 42.44	< 42.44
OH $3/2 \rightarrow 3/2$ ($7/2+ \rightarrow 5/2-$)	3551.6	84.41	< 39.11	< 39.11	70.95	< 39.11	< 39.11	< 39.11
[O I] $^3\text{P}_1 \rightarrow ^3\text{P}_2$	4745.1	63.18	837.08	417.61	355.54	38.76	31.51	183.45

 NOTE. — Spectral lines detected with SPIRE and PACS. Where no 3σ detection is found, we provide 3σ upper limits from local continuum. All line fluxes for HBC 722 are attributed primarily to contamination, with the exception of [O I] 63 μm ; we update the line fluxes from Green et al. (2011). Line fluxes for V1735 Cyg are partially contaminated by V1735 Cyg SM1. CO $J = 7 \rightarrow 6$ and [C I] 371 μm are partially blended. The CO $J = 23 \rightarrow 22$ may be blended with the H₂O 113.53 μm line.

¹Although not detected with SPIRE, we detect this line with HIFI, at an integrated line flux of $2.95 \times 10^{-18} \text{ W m}^{-2}$.

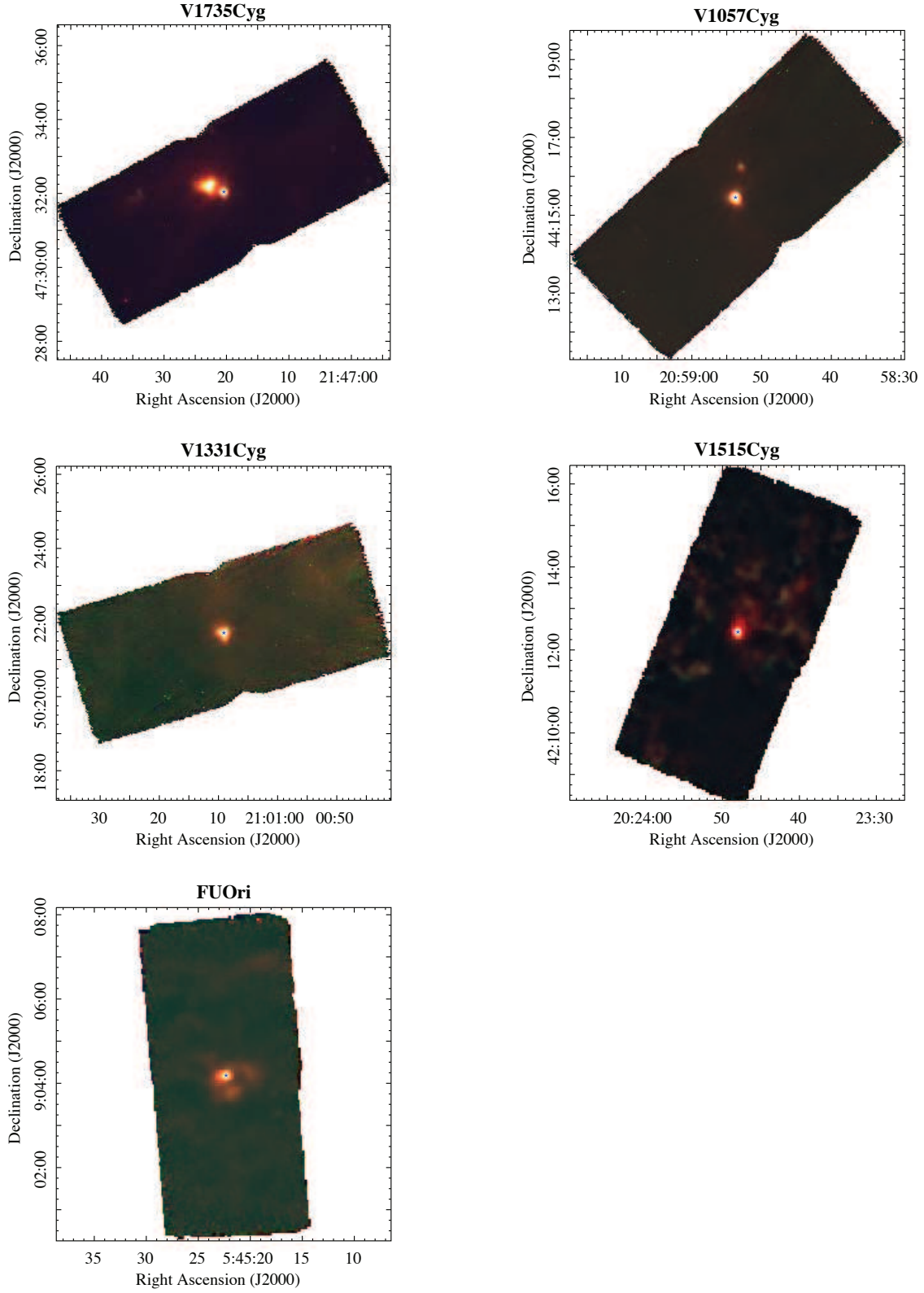


FIG. 1.— 70 (blue), 100 (green), and 160 μm (red) composite PACS imaging of the FUors in our sample. The positions of the optical FUors are marked in each image with a blue *.

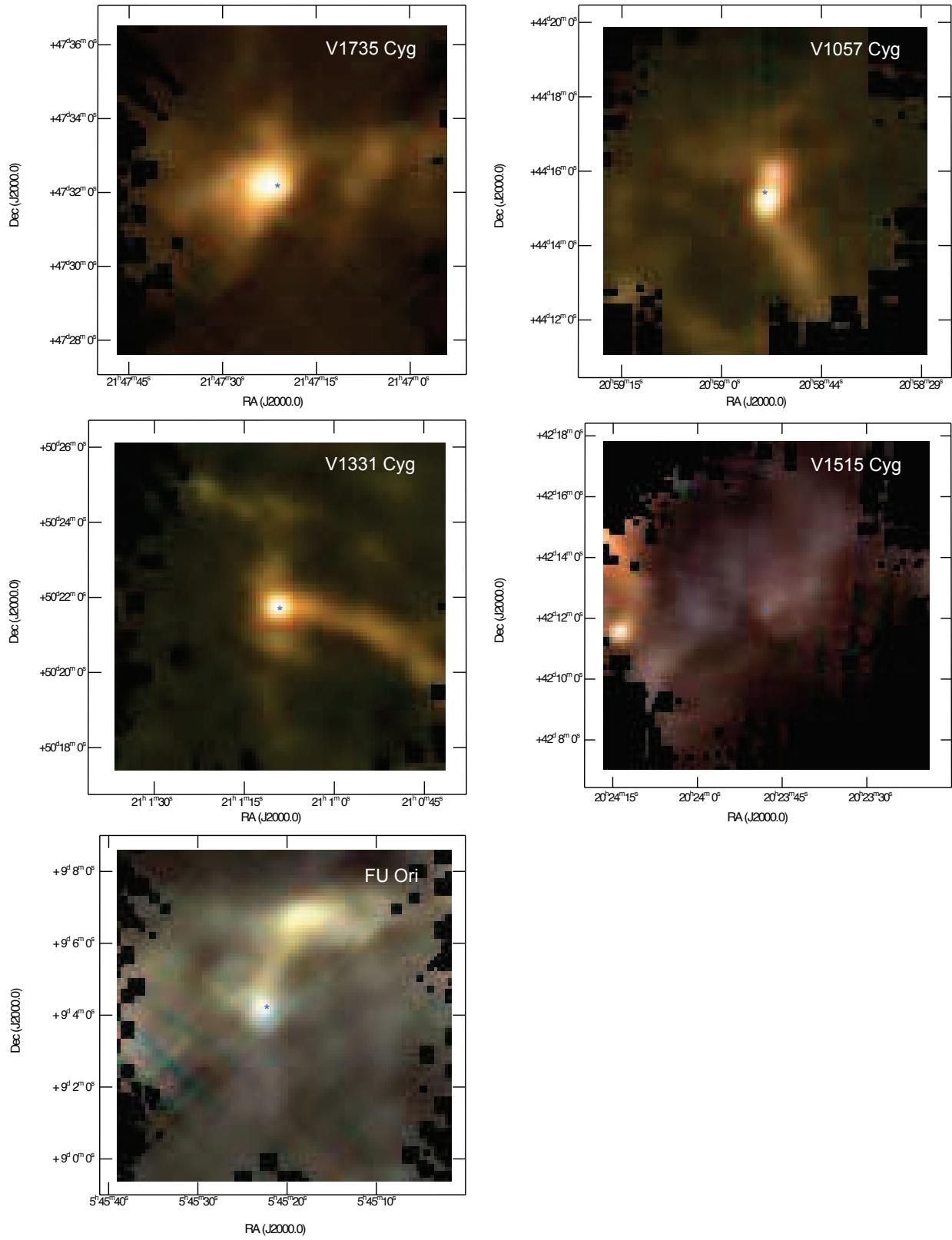


FIG. 2.— 250 (blue), 350 (green), and 500 μm (red) composite SPIRE imaging of the FUOrs in our sample. The positions of the optical FUOrs are marked in each image with a blue *.

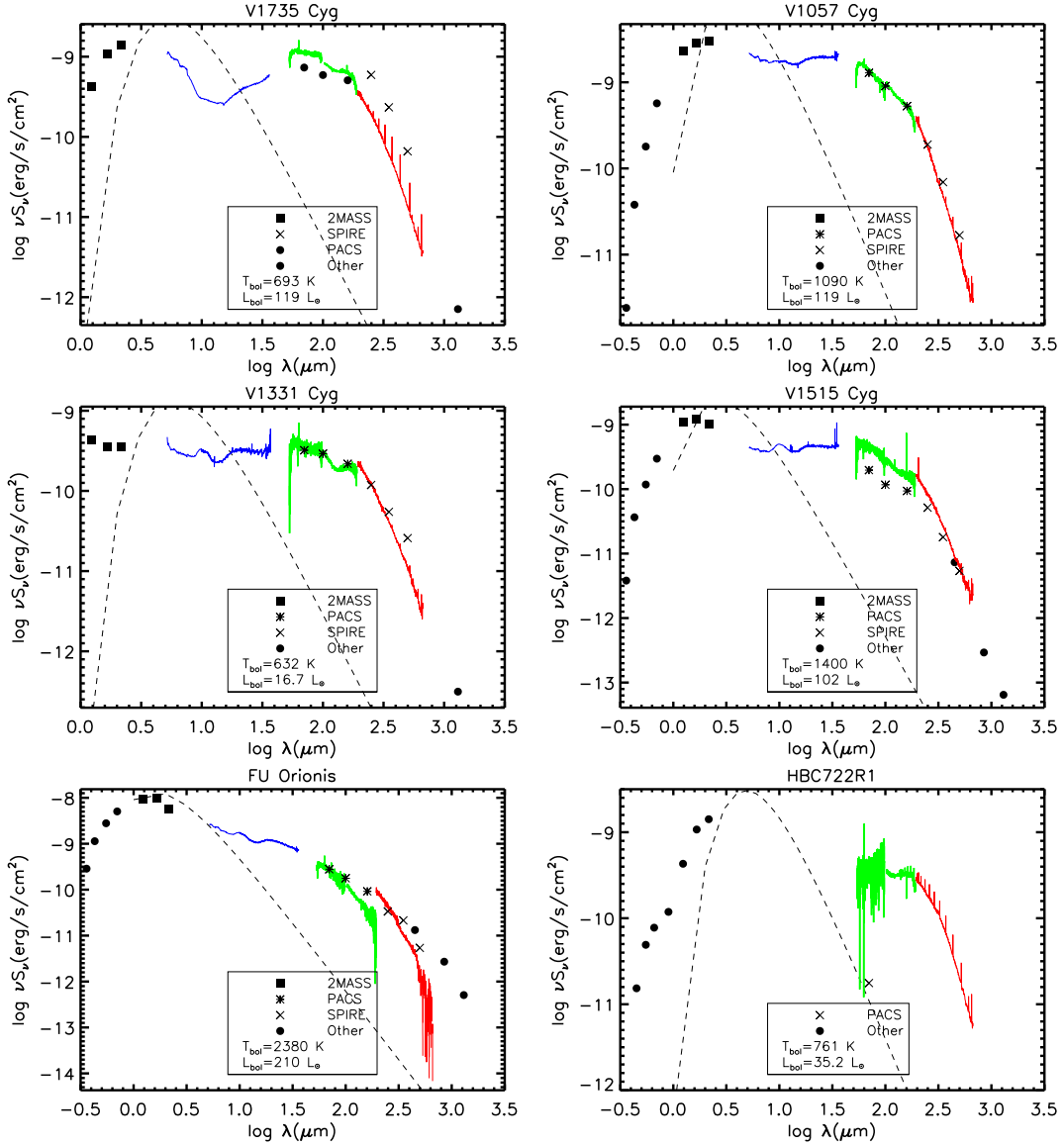


FIG. 3.— 0.5-1000 μm SEDs. Both the UBV R (Maidanak Observatory) photometric points and IRS spectra are circa 2005, analytically dereddened (Mathis 1990) in the cases of FU Ori ($A_V = 1.8$), V1515 Cyg ($A_V = 3.0$) and V1057 Cyg ($A_V = 3.5$). The other sources were not corrected for extinction. The data is taken from Green et al. (2006) and Quanz et al. (2007c). The near-IR points are primarily from 2MASS (JHK; filled squares; circa 2000) and the dereddening correction was again applied. The submm (‘*’) points are PACS 70, 100, and 160 μm , and the (‘X’) points are SPIRE 250, 350, and 500 μm (this work; circa 2011-12). When available we include 0.45-1.3 mm SCUBA data. The dashed line is a blackbody spectrum with the fitted bolometric luminosity and temperature. We include the contaminated SED of the HBC 722 composite for comparison (BVRIJHK from Kóspál et al. 2011, contemporaneous with our Herschel observations); note that only the optical/near-IR points and PACS 70 μm are attributed to HBC 722 itself. The large discrepancy between the PACS photometry and spectroscopy indicates that the PACS spectra are dominated by nearby extended sources.

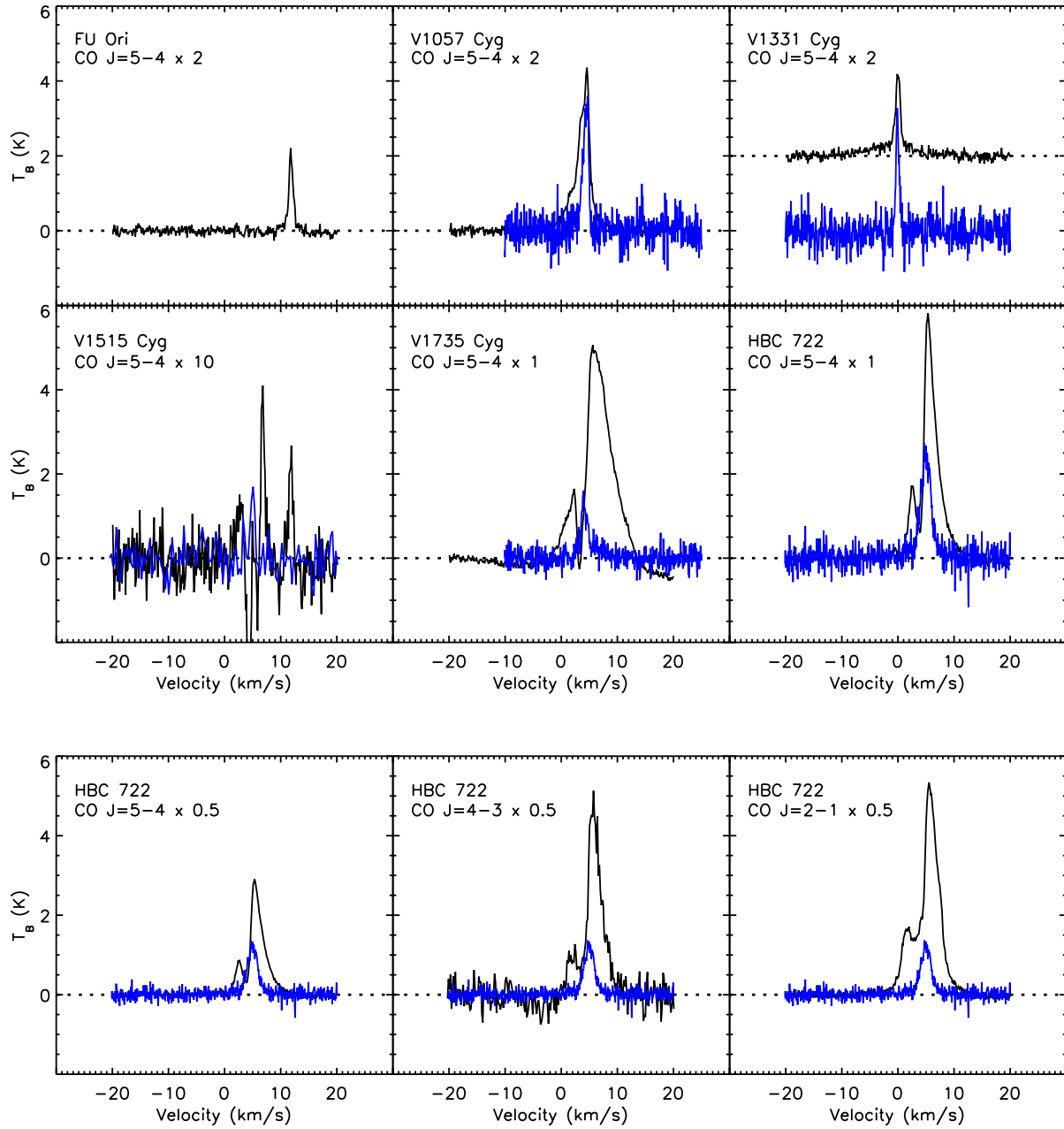


FIG. 4.— **Top:** HIFI observations of CO $J=5 \rightarrow 4$ (black) and HCO^+ $J=3 \rightarrow 2$ (blue). The dashed horizontal line indicates the baseline offset; the CO $J=5 \rightarrow 4$ in V1331 Cyg is offset for clarity. **Bottom:** CO $J=5 \rightarrow 4$ (HIFI), $J=4 \rightarrow 3$ and $J=2 \rightarrow 1$ (CSO) vs. HCO^+ $J=3 \rightarrow 2$, for HBC 722. Note that in each subfigure the black spectrum is scaled by the factors listed in its respective title.

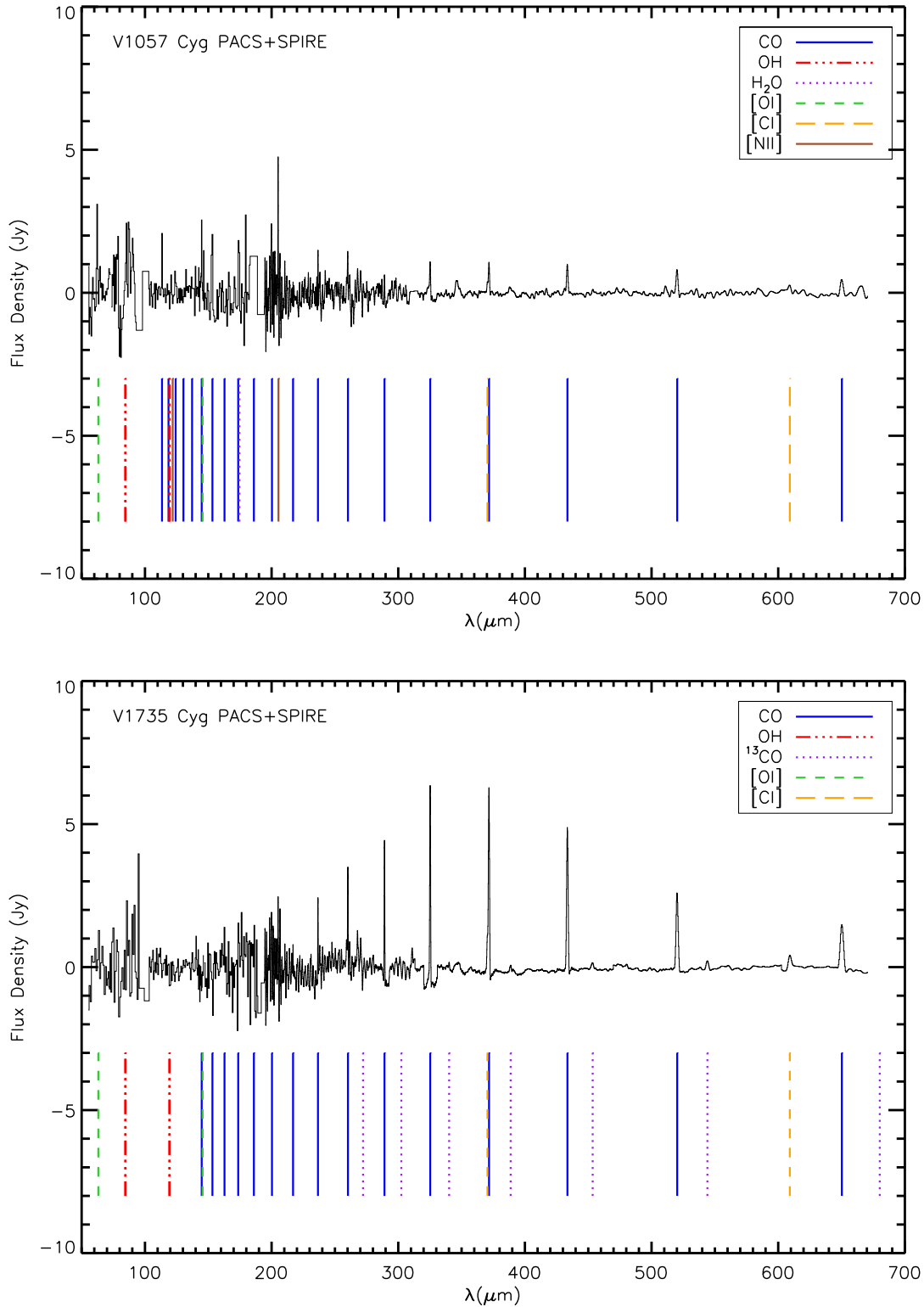


FIG. 5.— **Top:** Continuum-subtracted PACS/SPIRE spectrum of V1057 Cyg, rebinned to lower resolution for clarity. **Bottom:** Continuum-subtracted spectrum for V1735 Cyg, contaminated by V1735 Cyg SM1.

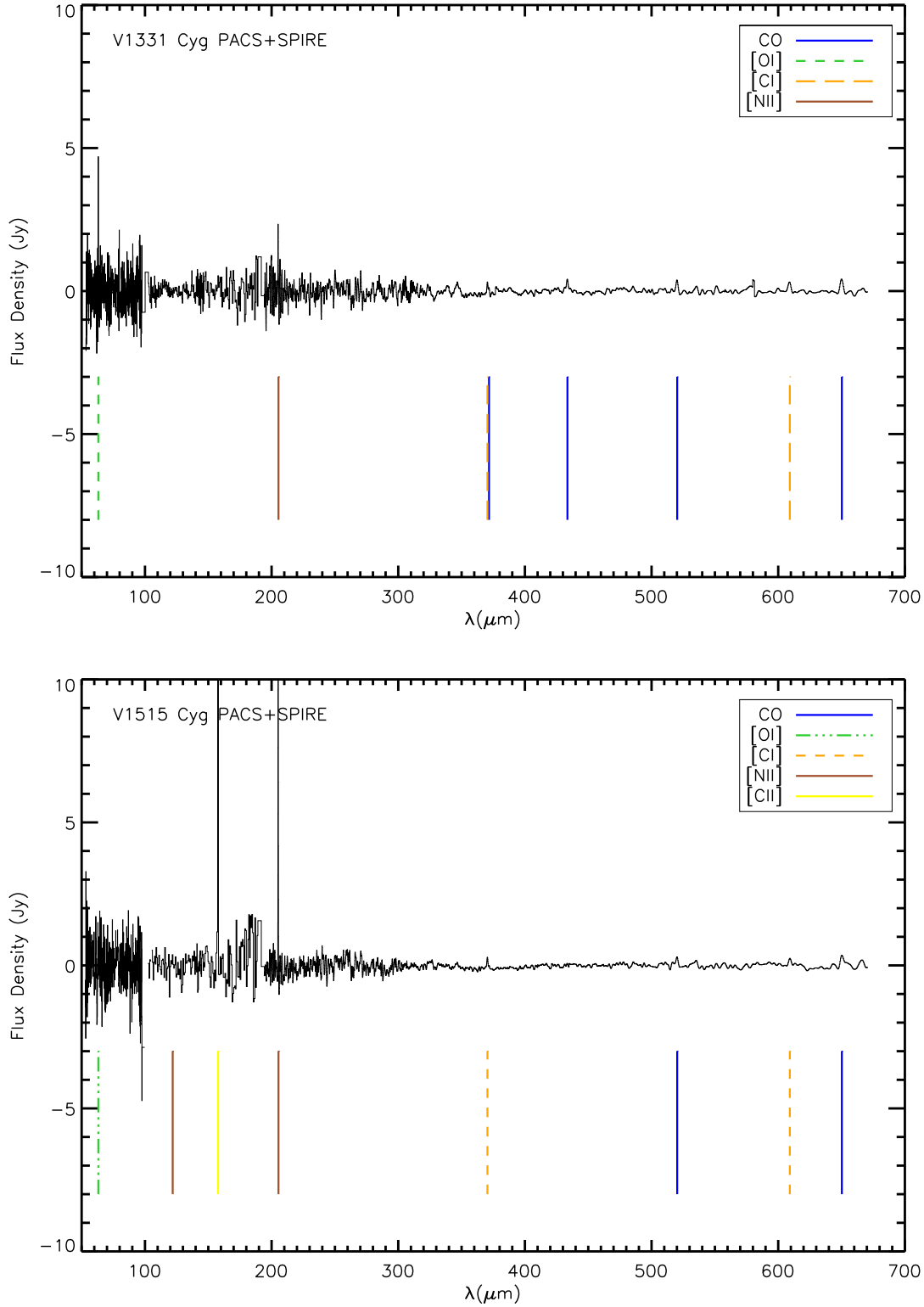


FIG. 6.— **Top:** Continuum-subtracted PACS/SPIRE spectrum of V1331 Cyg, rebinned to lower resolution for clarity. **Bottom:** Continuum-subtracted spectrum for V1515 Cyg.

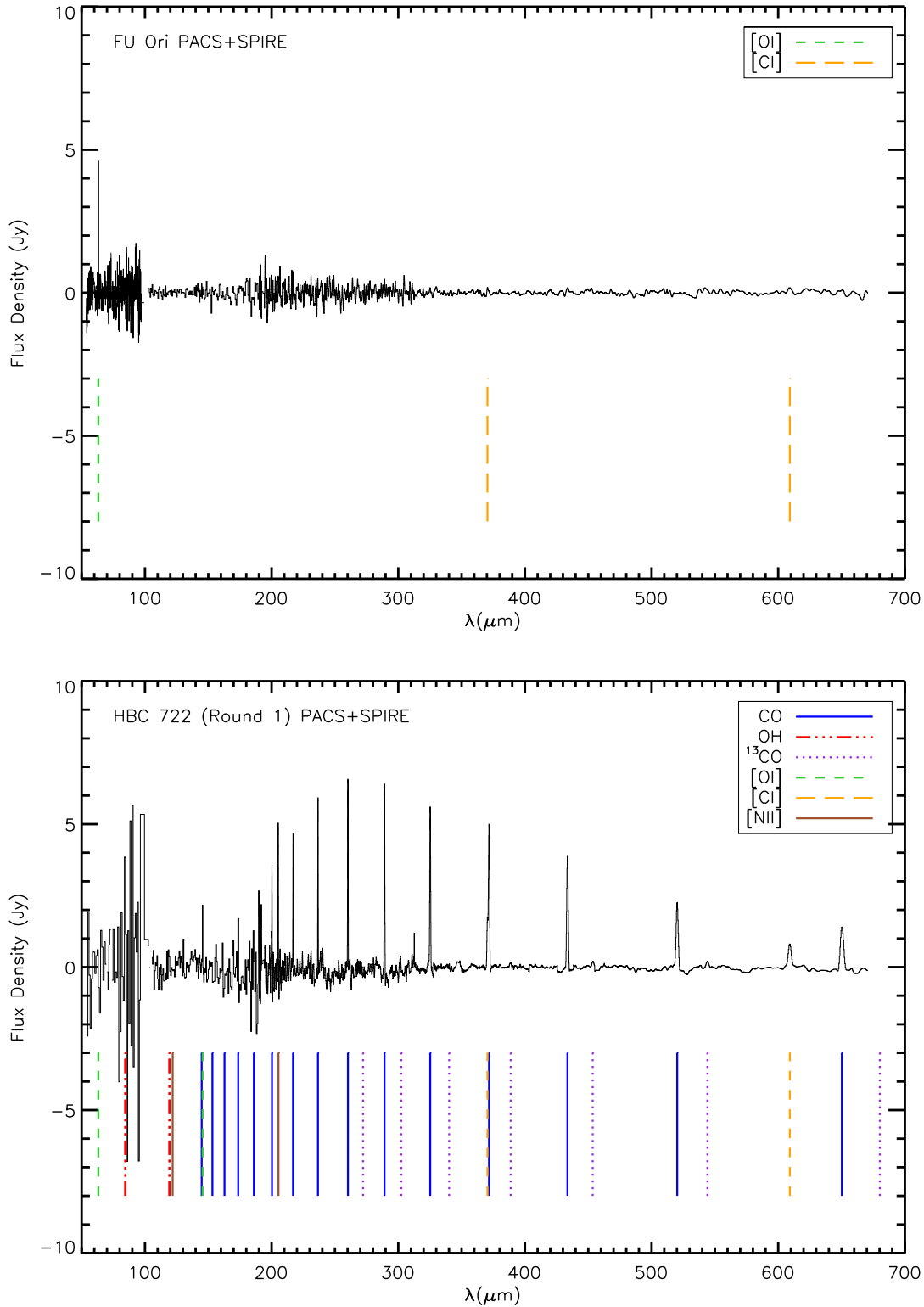


FIG. 7.— **Top:** Continuum-subtracted PACS/SPIRE spectrum of FU Ori, rebinned to lower resolution for clarity. **Bottom:** Continuum-subtracted spectrum for HBC 722 (taken during December 2010), contaminated by nearby Class 0/I protostars.

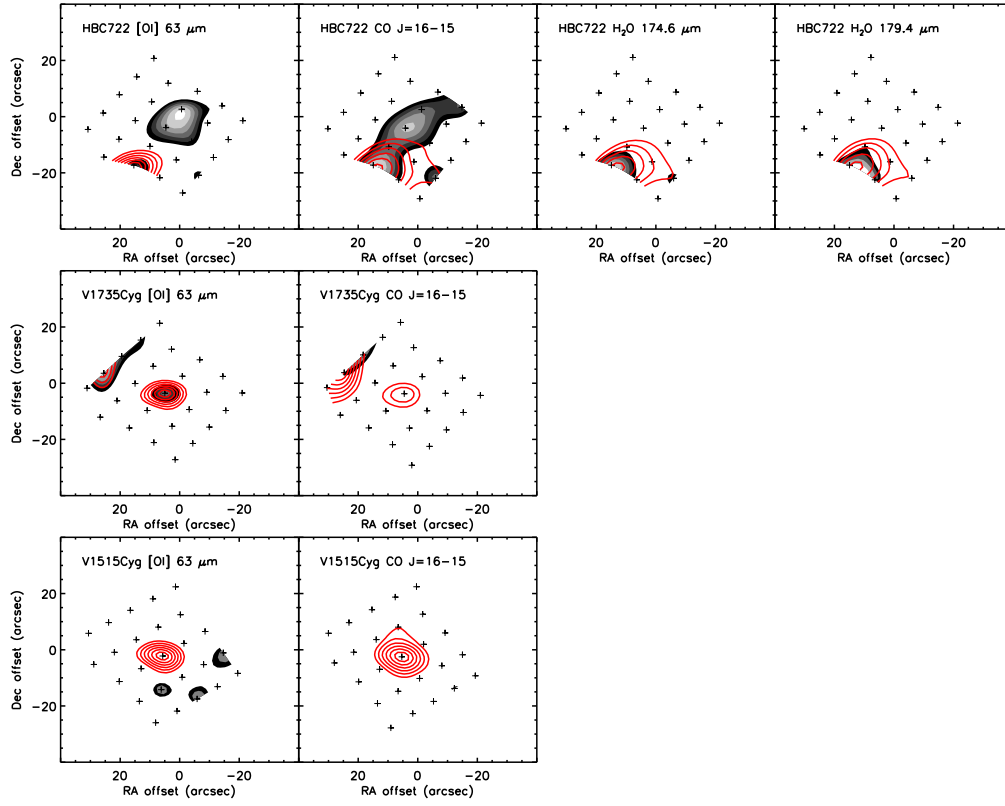


FIG. 8.— Spatial distribution of PACS lines (grayscale) vs. local continuum (red); the ‘+’ indicate the spaxel positions. The contours are in increments of 10% from the peak flux (usually in the central spaxel), plotted down to the noise limit. The noise limit is computed as $2\times$ the average flux of the outer ring of 16 spaxels, with the exception of the V1735 Cyg and HBC 722 maps, for which we avoid the edge spaxels containing a likely Class 0/I protostar (respectively, V1735Cyg SM1 and 2MASS 20581767+4353310).

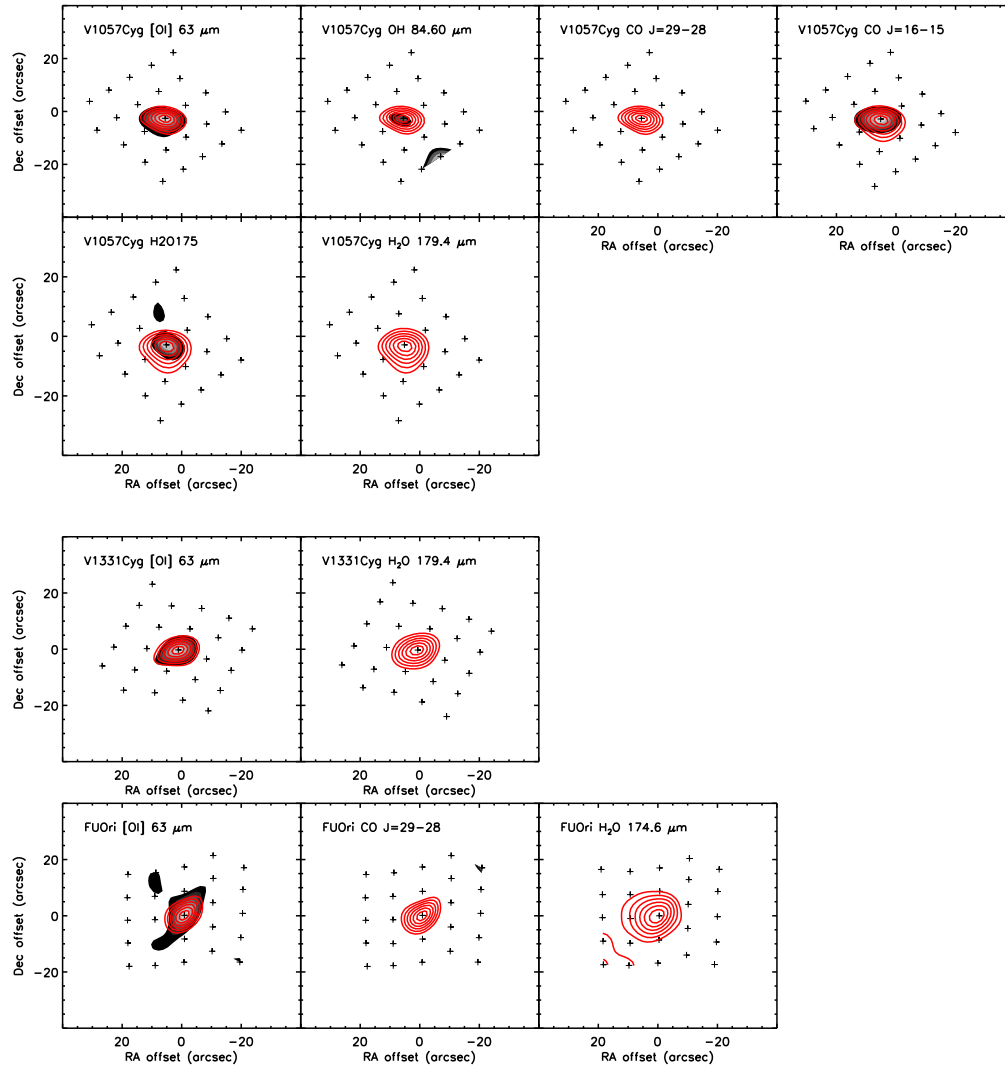


FIG. 9.— See Figure 8 for a description of this figure.

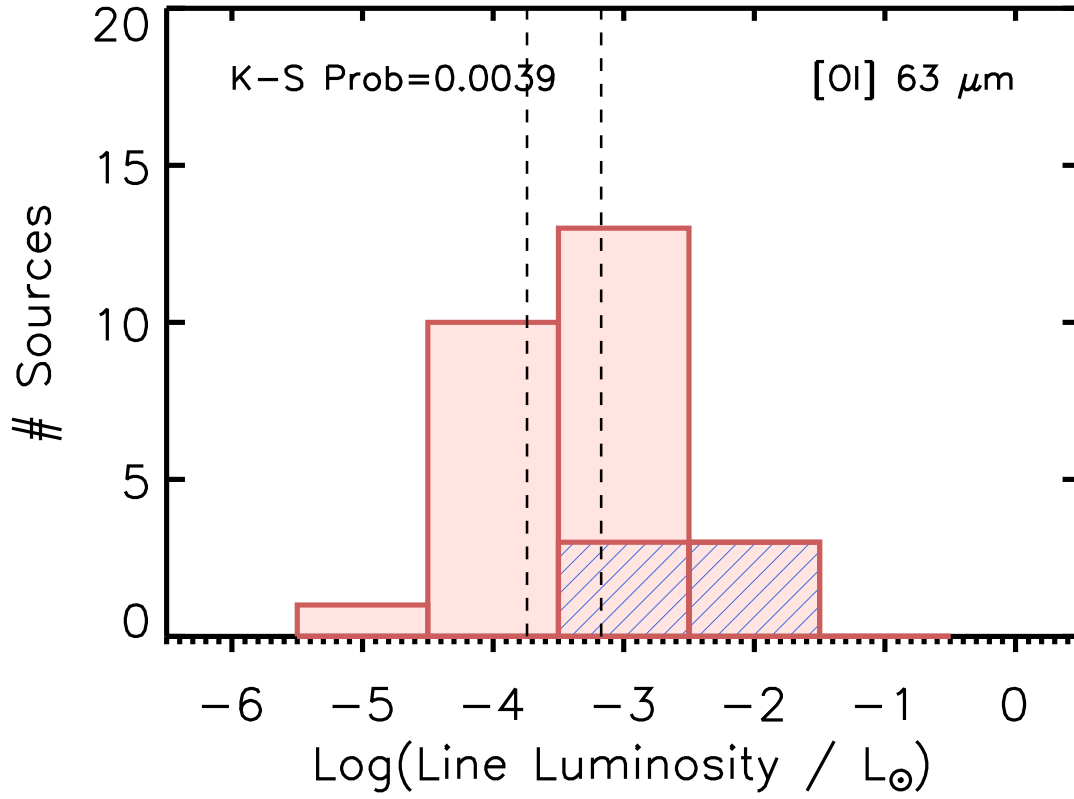


FIG. 10.— Histogram of [O I] 63 μm line luminosity for the full DIGIT embedded sample (Green et al. 2013). The FOOSH sample is superimposed in shaded bars. The probability of the distributions arising from the same source function is included as the K-S probability, and is less than 1%. The dashed vertical lines are the detection thresholds for the closest (520 pc) and most distant (1000 pc) sources in the FOOSH sample.

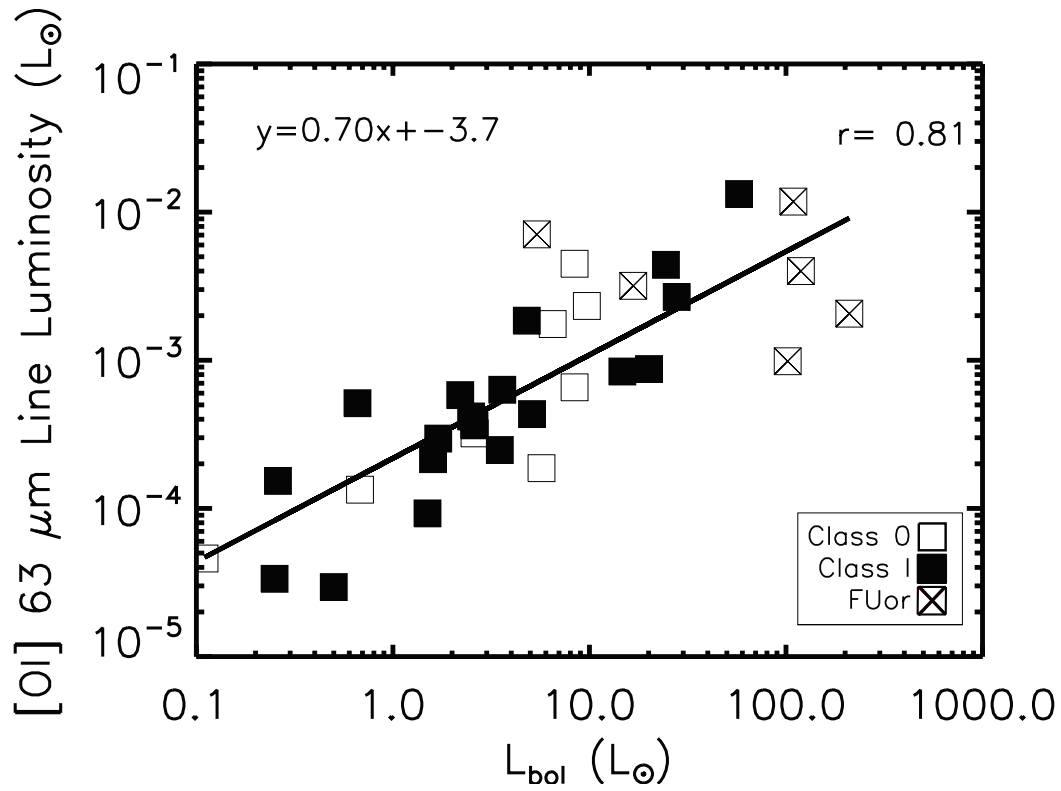


FIG. 11.— [O I] 63 μm vs. L_{bol} for the DIGIT and FOOSH samples. The open squares are Class 0 sources, the filled squares Class I, and the X squares are the FUors. The line is a best-fit log-log correlation characterized by the equation in the upper left portion of the plot, with $r=0.81$.

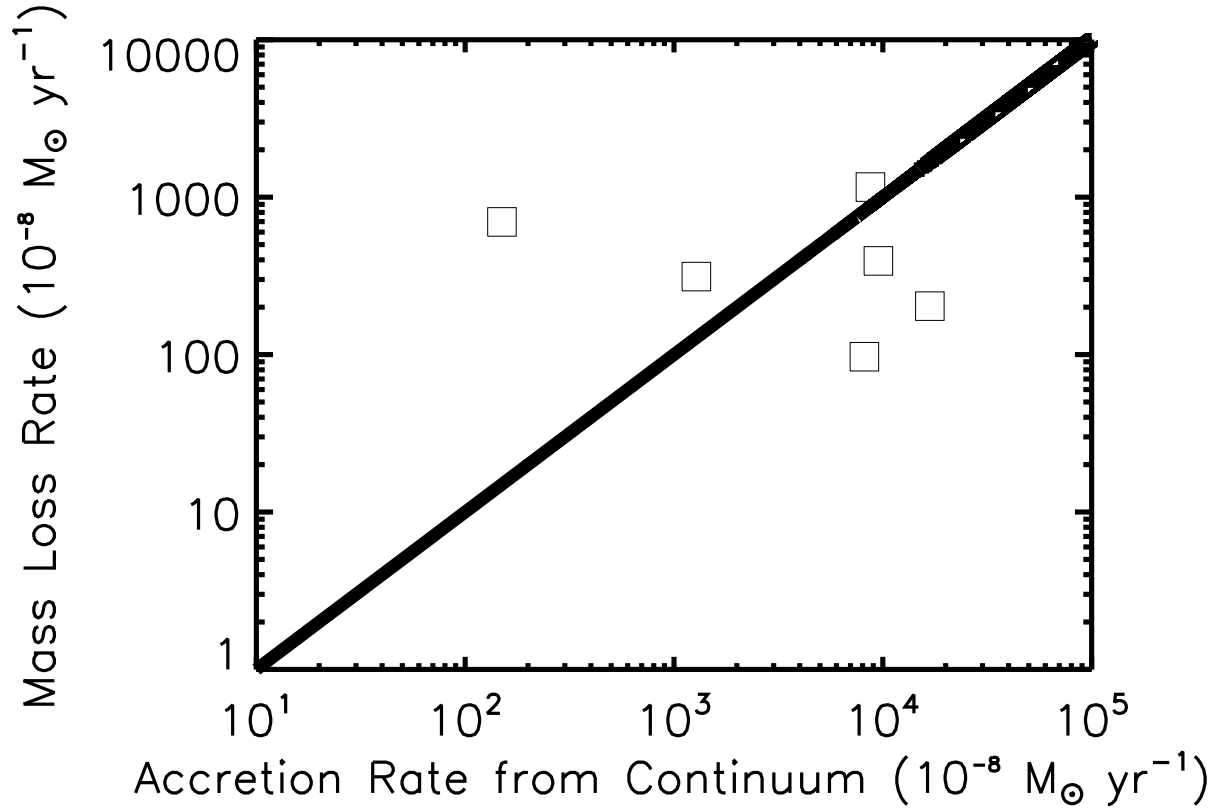


FIG. 12.— Comparison of the accretion rate (calculated from the bolometric luminosity) and the instantaneous mass loss rate as derived from [O I], for the six FUors, assuming a wind velocity of 100 km s^{-1} (Hollenbach 1985). The points would fall along the bold line if the [O I] mass loss rate and accretion luminosity tracers were in perfect correlation, assuming a conversion rate of 10% from accretion to outflow (e.g. Kurosawa et al. 2006); the greatest outlier is HBC 722.

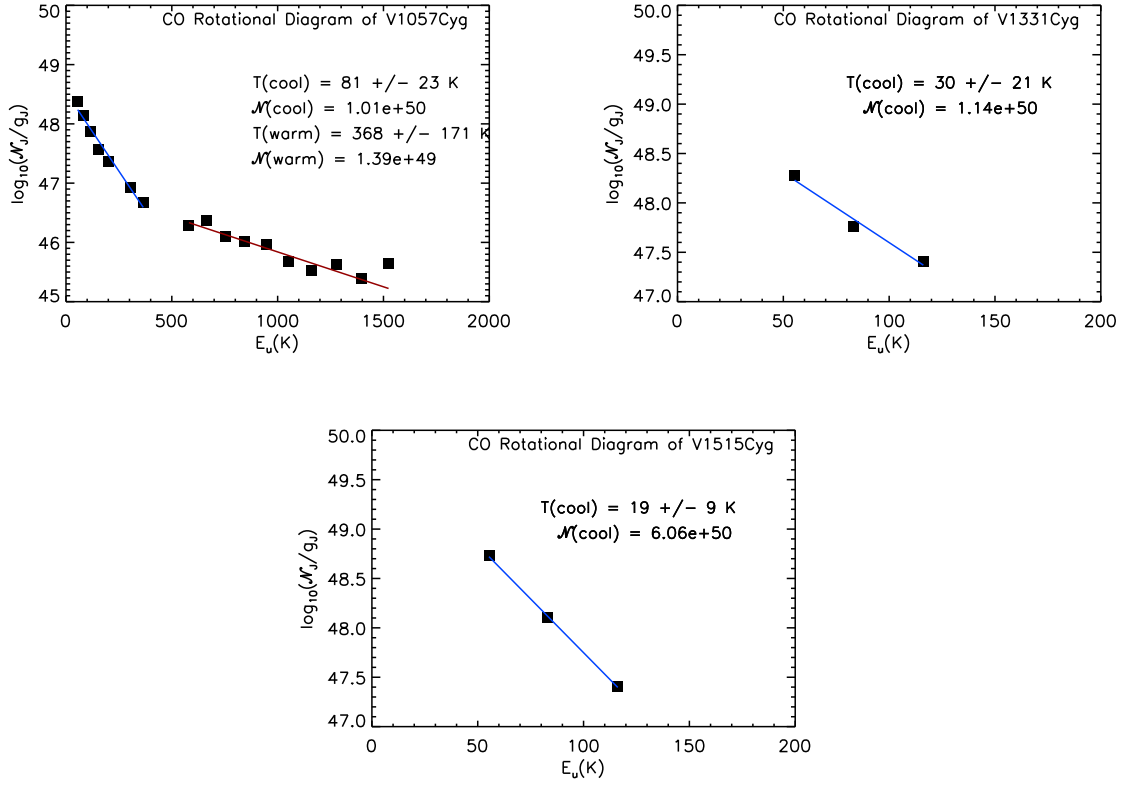


FIG. 13.— CO rotational diagrams for the unconfused FUors in our sample. CO $J = 23 \rightarrow 22$ is likely blended with H_2O and is not considered in the fit for V1057 Cyg. FU Ori is not detected in any CO transitions with SPIRE; however a narrow CO $J = 5 \rightarrow 4$ component was detected with HIFI, below the detection threshold of SPIRE.

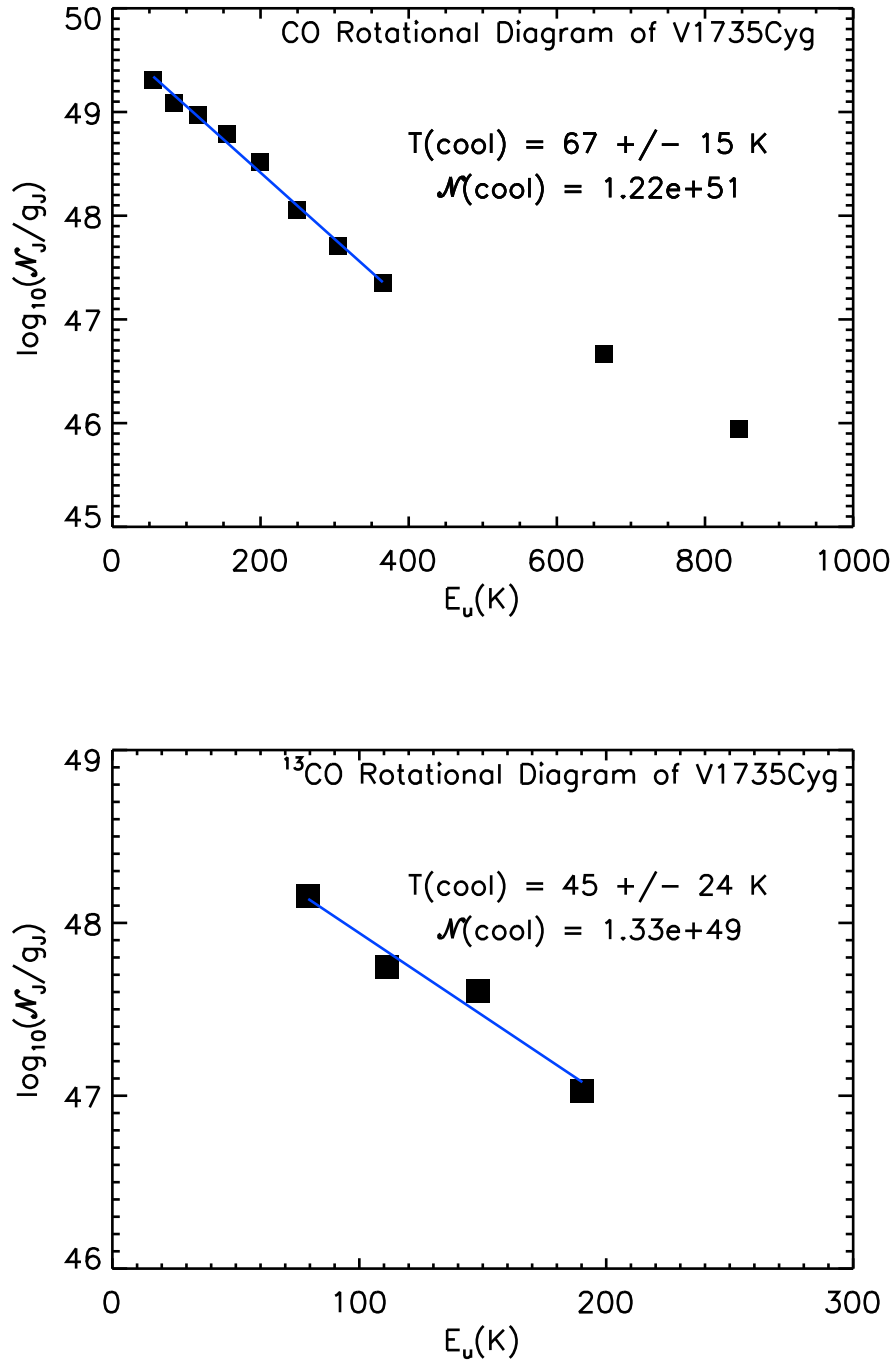


FIG. 14.— CO (top) and ^{13}CO (bottom) rotational diagram for the contaminated V1735 Cyg SPIRE spectrum.

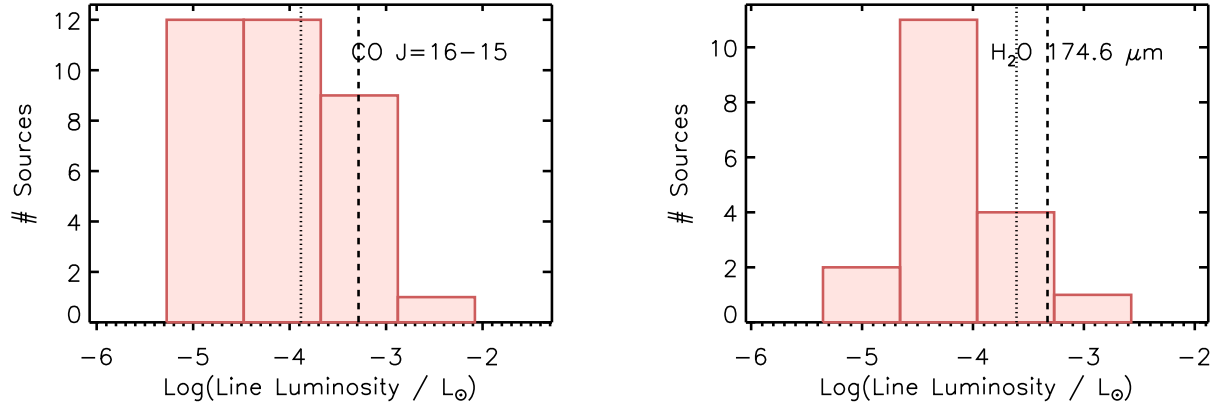


FIG. 15.— **Left:** Histogram of luminosities of the CO $J = 16 \rightarrow 15$ line for the DIGIT sample from (Green et al. 2013). The vertical dashed line indicates the line strength for V1057 Cyg. The vertical dotted line marks the 3σ upper limit for sources with non-detections in our sample. **Right:** The same analysis for the H₂O 174.63 μm line, detected in V1057 Cyg.